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The central dark matter content of early-type galaxies: scaling relations and connections with star formation histories

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ABSTRACT

We examine correlations between masses, sizes and star formation histories for a large sample of low-redshift early-type galaxies, using a simple suite of dynamical and stellar population models. We confirm an anticorrelation between the size and stellar age and go on to survey for trends with the central content of dark matter (DM). An average relation between the central DM density and galaxy size of $\langle \rho_{\text{DM}} \rangle \propto R_{\text{eff}}^{-2}$ provides the first clear indication of cuspy DM haloes in these galaxies – akin to standard Λ cold dark matter haloes that have undergone adiabatic contraction. The DM density scales with galaxy mass as expected, deviating from suggestions of a universal halo profile for dwarf and late-type galaxies.

We introduce a new fundamental constraint on galaxy formation by finding that the central DM fraction decreases with stellar age. This result is only partially explained by the size–age dependencies, and the residual trend is in the opposite direction to basic DM halo expectations. Therefore, we suggest that there may be a connection between age and halo contraction and that galaxies forming earlier had stronger baryonic feedback, which expanded their haloes, or lumpier baryonic accretion, which avoided halo contraction. An alternative explanation is a lighter initial mass function for older stellar populations.

Key words: galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: fundamental parameters – dark matter.

1 INTRODUCTION

The formational history of early-type galaxies (ETGs; ellipticals and lenticulars) remains an outstanding question. While these dynamically hot systems may be basically understood as end-products of galaxy mergers, the details of these mergers and their cosmological context are unclear. High-redshift (z) observations have made initially surprising discoveries that many ETGs were already present at early times with mature stellar populations and that these galaxies were much more compact than those in the present day (e.g. Glazebrook et al. 2004; Daddi et al. 2005; Trujillo et al. 2006).

The evolution in ETG sizes is still controversial in both observation and interpretation (e.g. van Dokkum, Kriek & Franx 2009; Mancini et al. 2010; Valentinuzzi et al. 2010; van Dokkum et al. 2010). However, the most likely scenario is for a combination of effects where individual galaxies grow in size by accretion of smaller, gas-poor galaxies (an ‘inside-out’ picture of galaxy formation) and where younger ETGs are formed with larger sizes because of their decreased cold gas content and the lower background densities of

dark matter (DM; e.g. Bezanson et al. 2009; Naab, Johansson & Ostriker 2009; Saracco, Longhetti & Andreon 2009; van der Wel et al. 2009; Hopkins et al. 2010; Shankar et al. 2010).

The role of DM is considered fundamental to the formation of galaxies, and DM halo properties have been extensively studied in cases such as gas-rich spirals and nearby dwarfs which have suitable observational tracers (e.g. McGaugh et al. 2007; Walker et al. 2009; Kalirai et al. 2010). Studying DM in the general population of ETGs is in many ways more difficult, with ongoing surveys of the available large-radius halo tracers attempting to remedy our ignorance in this area (e.g. Coccato et al. 2009; Proctor et al. 2009; Romanowsky et al. 2009; Thomas et al. 2009, hereafter T+09; Woodley et al. 2010).

DM can alternatively be studied with less precision but in more extensive ETG samples by considering the well-studied central regions – inside the ‘effective radius’ (R_{eff}) enclosing half the stellar light, where DM is generally thought to be a minor yet potentially detectable contributor to the mass. Here one of the classic approaches is to analyse the ‘Fundamental Plane’ (FP) relating ETG sizes, luminosities and central velocity dispersions (σ_0). The FP shows a ‘tilt’ or systematic deviation from simple expectations based on galaxies with constant dynamical mass-to-light ratios M/L ,

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probably implying systematic differences in the stellar populations or in DM content.

After many years of debate, there is still no consensus on what is driving the FP tilt (e.g. Trujillo, Burkert & Bell 2004; Ferreras, Saha & Williams 2005; Boylan-Kolchin, Ma & Quataert 2006; Cappellari et al. 2006, hereafter C+06; Dekel & Cox 2006; Bolton et al. 2007, 2008; Jun & Im 2008; Allanson et al. 2009; Graves 2009; Tortora et al. 2009, hereafter Paper I; Grillo & Gobat 2010; Humphrey & Buote 2010; La Barbera et al. 2010a). Some work suggests stellar population variations or non-homologies in the luminosity profiles as the tilt drivers. However, it is a fairly generic expectation from the standard cosmological framework for galaxy formation that the central DM content of ETGs will systematically increase with luminosity – a point we discussed in Paper I and develop further in this paper. If the FP tilt is *not* caused in large part by DM, there could be problems implied for galaxy formation theory. Here one could pursue two different philosophies: to empirically and robustly determine the reasons for the FP tilt without recourse to theory (e.g. making no assumptions about the underlying DM profiles) or to adopt the theoretical framework as broadly correct and consider the detailed implications for ETG composition and formation.

Following the second approach, it is now time to begin moving on from phenomenological questions about the FP tilt and to establish more direct connections between DM in ETGs and their formational histories. In particular, the central DM content could prove crucial to solving the size-evolution puzzles mentioned above and more fundamentally to understanding the assembly of ETGs. To this end, we now extend the analysis of Paper I, which combined models of stellar dynamics and population synthesis to infer *total* and *stellar* masses in a large sample in nearby galaxies and thereby to analyse the FP tilt. There have been previous suggestions that age and star formation time-scales are important fourth parameters in the FP (Terlevich & Forbes 2002; Gargiulo et al. 2009; Graves 2009). We will now explore this possibility systematically, using the results from Paper I to consider additional correlations involving DM content and star formation histories (SFHs).

It should be noted at the outset that the data and the analysis techniques that we use for deriving mass and SFH parameters may not be the most state-of-the-art. However, our aim is to pioneer a framework for interpreting any data of this kind in a broad cosmological context (where, in particular, emerging high- z data sets provide only crude observational constraints on ETGs), and we are so far able to tentatively identify some basic and intriguing trends.

This paper is organized as follows. In Section 2, we present our basic observational results. We go on to analyse some implications of the trends of DM with mass in Section 3 and with age in Section 4. In Section 5, we summarize our conclusions. Appendices A and B include analyses of statistical and systematic uncertainties.

2 OBSERVATIONAL RESULTS

Here we present our basic observational results. Section 2.1 provides an overview of our galaxy sample and mass inferences and summarizes the trends versus mass. Section 2.2 presents results related to galaxy age.

2.1 Summary of sample and initial results

Our starting point is a collection of 335 local ETGs from Prugniel & Simien (1996) that we recently re-analysed in Paper I. The sample was selected to have measurements of σ_0 and at least two colours, with a subsample of ~ 220 systems that also include the maximum rotation velocity (V_{\max}). There are 218 elliptical galaxies and 117

lenticular/S0 systems: as shown in Paper I, the two subsamples show similar stellar and DM properties and we will consider them jointly in this work.

We next summarize the main steps of the Paper I analysis, starting with the stellar population models. We used a set of simple stellar population (SSP) synthetic spectra from the prescription of Bruzual & Charlot (2003, hereafter BC03) to fit the observed galaxy colours. We assumed a Salpeter (1955) or Chabrier (2001, 2002, 2003) initial mass function (IMF) alternatively, with initial masses m in the range of $0.1\text{--}100 M_{\odot}$. For this paper, we instead use an intermediate Kroupa (2001) IMF as our default model.¹

These single-burst models were convolved with an exponential star formation rate (SFR) $\propto e^{-t/\tau}$ to generate more general SFHs, where τ is a characteristic time-scale. The age, metallicity (Z) and τ were free parameters of the model while stellar M/L (Υ_{\star}) was inferred from the best-fitting solution for each galaxy. The reliability of the modelling technique and the intrinsic parameter scatter, as well as the presence of spurious correlations among the stellar parameters induced by the stellar modelling procedure, has been checked through Monte Carlo simulations (see Paper I for details, as well as Appendix B3 of this paper). In addition, a recent novel, completely independent technique for estimating Υ_{\star} using globular cluster systems has yielded results in perfect agreement with ours (fig. 6 of Forte, Vega & Faifer 2009).

We derived the dynamical M/L (Υ_{dyn}) within R_{eff} by means of Jeans analysis assuming spherical symmetry, isotropy of the velocity dispersion tensor and introducing a rotation velocity correction in the spherically averaged rms velocity $v_{\text{rms}} = \sqrt{v^2 + \sigma^2}$, which is a measure of the total kinetic energy of the system within R_{eff} . In the following, we will use Υ_{dyn} obtained with the multicomponent model in Paper I [i.e. the Sérsic (1968) profile for the stars and singular isothermal sphere for the total mass]; we have verified that the particular mass model choice does not affect the conclusions of this paper.

Central DM fractions are inferred by stellar and dynamical M/L estimates, being by definition

$$f_{\text{DM}} = \frac{M_{\text{tot}} - M_{\star}}{M_{\text{tot}}} = 1 - \frac{M_{\star}}{M_{\text{tot}}} = 1 - \frac{\Upsilon_{\star}}{\Upsilon_{\text{dyn}}}, \quad (1)$$

where Υ_{\star} and Υ_{dyn} are obtained in Paper I;² the DM mass includes any mass in diffuse gas. Note that as seen in Fig. 1, a fraction of the galaxies (under any of our IMF choices) appear to have $f_{\text{DM}} < 0$, which would be an unphysical situation. This is not necessarily a worry, since errors in the M/L estimates can scatter some data to $f_{\text{DM}} < 0$. We have investigated this issue quantitatively in Appendix A, finding a strong but not definitive suggestion that some ETGs are not compatible with having a Salpeter IMF. Similar points were made in other studies (Renzini 2005; C+06; Ferreras, Saha & Burles 2008; Cappellari et al. 2009; Barnabè et al. 2010) using conceptually similar techniques. Treu et al. (2010, hereafter T+10) on the other hand claimed from an analysis of gravitational lenses that a Salpeter IMF is preferred, with a possible systematic IMF variation. We will

¹ For old stellar populations ($\gtrsim 1$ Gyr), changing the IMF generally impacts the colours at less than the ~ 0.01 mag level. The IMF change can then be fairly represented as a simple overall stellar M/L renormalization, which is reduced by a factor of 1.6 in changing from Salpeter to Kroupa (see e.g. fig. 4 of BC03).

² While Υ_{dyn} is a three-dimensional quantity, Υ_{\star} is *projected* and susceptible to additional variations from large-radius material along the line of sight. However, we have estimated using some simple models that Υ_{\star} would typically be affected at only the ~ 1 per cent level.

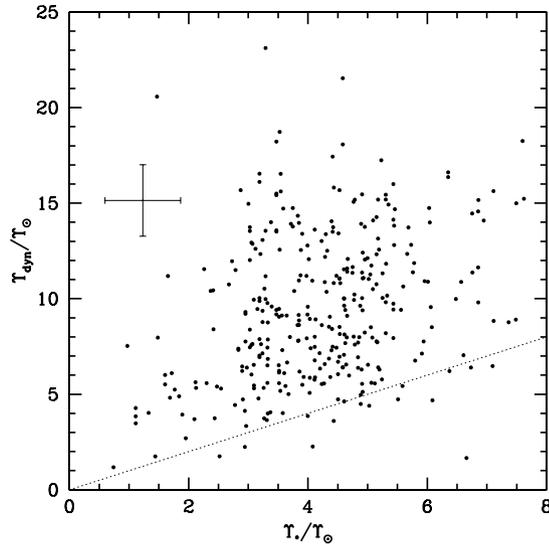


Figure 1. Stellar versus dynamical M/L within R_{eff} for the ETG sample as derived in Paper I, using a Kroupa IMF. The dotted line is the one-to-one relation and typical statistical uncertainties are illustrated by the error bars at left. For most of the galaxies, the presence of a significant mass excess is evident, implying a non-zero DM fraction in the central R_{eff} .

return to these issues in later sections, and for now note that an assumed universal Kroupa IMF is a reasonable starting point.

We found in Paper I that the f_{DM} value correlates with luminosity and stellar mass (Fig. 2), a trend that could be the *phenomenological* cause of the FP tilt.³ We also examined the *formational* causes of the tilt in Paper I (Section 6), using toy models of galaxies in a cosmological context to verify that the expected scaling relations of DM haloes might naturally explain the f_{DM} trends. We will develop this theme further in the following sections and also begin studying the implications of the SFHs.

One basic SFH result is illustrated by Fig. 3: the stellar ages and masses of the galaxies are correlated, providing another example of the well-known ‘archaeological downsizing’ phenomenon wherein more massive systems form earliest. We plan to consider the physical reasons for this downsizing in a subsequent paper, comparing cosmologically based models for SFHs and DM halo assemblies. Here we will simply treat the SFHs as given which we will attempt to correlate with DM properties, envisioning these as fundamental ingredients for constraining future theoretical models.

2.2 New trends with age

Here we move beyond the FP analysis of Paper I and examine in more detail the DM trends, including the connections with SFHs. We plotted in Fig. 2 the empirical values for f_{DM} versus stellar mass; the two quantities are correlated albeit with substantial scatter and suggestions of non-monotonicity. We have also colour-coded the data points in Fig. 2 by stellar *age*,⁴ and can see immediately

³ In principle, the tilt could instead be caused by a systematic variation of the IMF with galaxy luminosity (see Paper I). We will reconsider IMF variations later in this paper. The effects of ‘non-homology’ of the stellar profiles were implicitly corrected for in our modelling.

⁴ Hereafter, where applied, the smoothed colour coding is obtained by linearly interpolating the average values of the third (colour-coded) quantity over a regular grid of the other two plotted quantities. This procedure allows us to readily see qualitative trends in the data, with the caveat that *quantifying* these trends can be unreliable after smoothing.

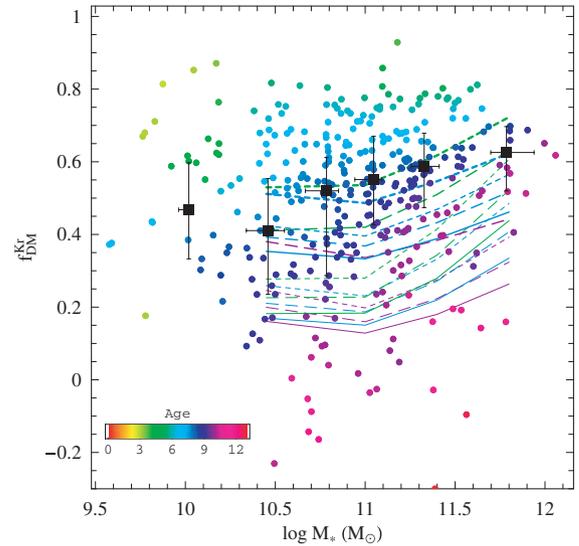


Figure 2. Relation between the central DM fraction f_{DM} and stellar mass M_* for ETGs, for an assumed Kroupa IMF. Individual galaxy data points are shown, coloured according to their ages as illustrated by the colour bar at lower left; the local age variations in the plots are smoothed in order to see the general trends. Except for the lowest masses (which are dominated by S0s), f_{DM} increases on average with M_* (illustrated by the black points with error bars showing the binned averages with scatter). The curves show various predictions of Λ CDM toy models (see Section 3.1). The bottom set of nine curves shows model predictions without adiabatic contraction (AC), for three age bins (3, 7, 11 Gyr, again coloured by age), and three ϵ_{SF} values (0.7, 0.3, 0.1: solid, long-dashed, short-dashed, respectively). The top set of curves show predictions with AC, including a smaller variation of model parameters for the sake of clarity. Long-dashed curves show the three age bins again, with $\epsilon_{\text{SF}} = 0.3$ and the G+04 AC recipe. Solid and short-dashed curves show the $\epsilon_{\text{SF}} = 0.7$ and 0.1 cases for G+04 AC and 7 Gyr. The heavy long-dashed curves show $\epsilon_{\text{SF}} = 0.1$ for B+86 AC and 7 and 3 Gyr.

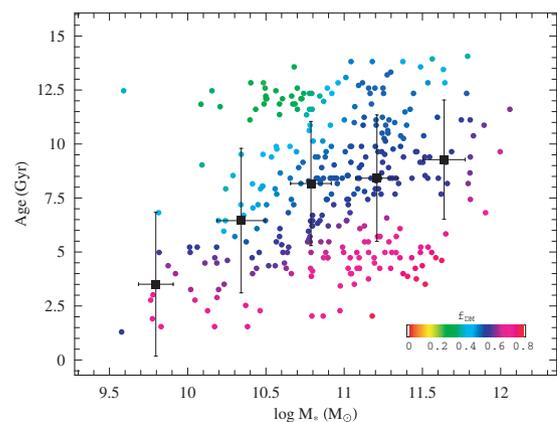


Figure 3. Galaxy stellar age versus stellar mass. The points are colour-coded according to the DM fraction (after smoothing; see the inset colour bar). The overall average trend in mass bins is shown by points with error bars representing the 1σ scatter.

that much of the scatter in f_{DM} might be explained by systematic correlations with SFHs. In particular, *at a fixed mass the galaxies with the youngest stars are found to have on average the highest DM fractions*. This is a central result of our paper which we consider in more detail below. Trends involving τ will be studied in a subsequent paper.

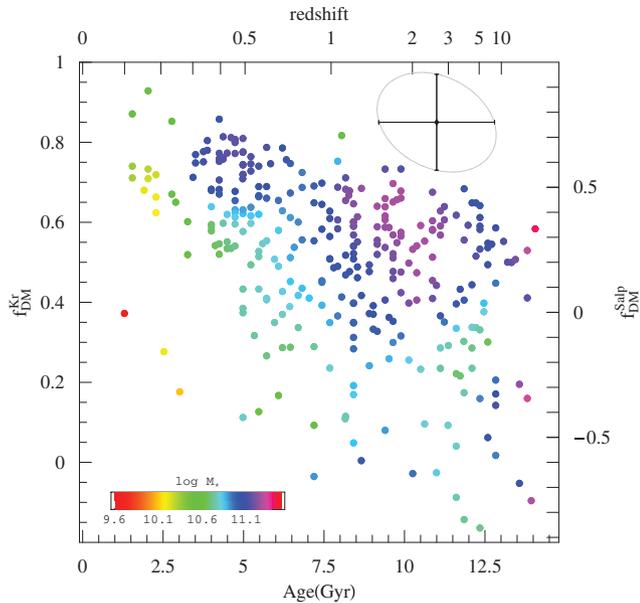


Figure 4. Central DM fraction of ETGs versus stellar age. The results for a Kroupa IMF are indicated by the left-hand axis and for Salpeter by the right-hand axis. The data points are colour-coded by the locally averaged stellar mass, with colour coding as in the inset. A typical uncertainty ellipse for an individual galaxy (at the centre of the range) is shown at upper right.

We show two other projections of the three-dimensional space of age–mass– f_{DM} in Figs 3 and 4. The combination of f_{DM} –mass and age–mass correlations (Figs 2 and 3) suggests that there could be an overall f_{DM} –age correlation. However, Fig. 4 shows that there is a net *anticorrelation* between f_{DM} and age, which implies that f_{DM} couples more strongly to age than to mass (we will revisit this issue in Section 4.1).

Before analysing these trends in more detail, we consider that there is an important *fourth* parameter involved, the effective radius R_{eff} . This is because f_{DM} is evaluated at R_{eff} , a variable parameter that has its own systematic dependencies. As we found in Paper I and will revisit below, the well-known ETG size–mass relation ($R_{\text{eff}}-M_*$) can explain much of the $f_{\text{DM}}-M_*$ trends as an ‘aperture effect’, without invoking any intrinsic variation in the DM properties. This is because the stellar R_{eff} varies more quickly with mass than do the inner DM halo properties [in Λ cold dark matter (ΛCDM)], meaning that for a more massive galaxy, the larger R_{eff} encloses a higher fraction of the total DM.

The next question is whether there is a link between R_{eff} and age that could be driving the f_{DM} –age trends. As mentioned in Section 1, the size–mass relation is thought to evolve strongly with time, such that ETGs at a fixed mass are more compact at earlier times. This trend can be considered generically to include contributions both from size growth of existing galaxies and from the birth of new galaxies with systematically larger sizes. As a consequence of the latter effect, we would expect the population of ETGs at a fixed time (e.g. $z = 0$) to show a correlation between size and age at fixed mass, such that the older galaxies are smaller.

This qualitative prediction has recently been confirmed by low- z observations (SB09; van der Wel et al. 2009), and we now investigate it for our own ETG sample. We plot R_{eff} against age in Fig. 5, where galaxies are colour-coded according to four fiducial stellar mass bins that we will use throughout the paper: $\log_{10}(M_*/M_{\odot}) \sim 10.3, 10.8, 11.2, 11.6$ for a Kroupa IMF. The overall galaxy distribution does not show any significant dependence of R_{eff} on age,

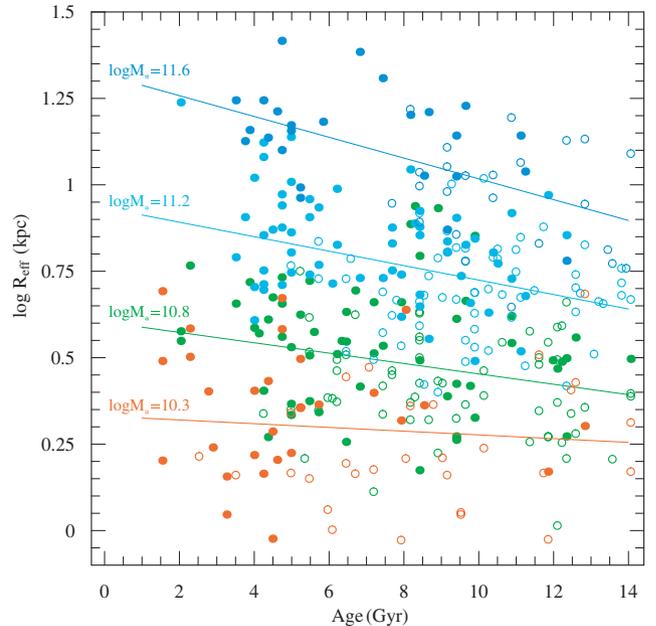


Figure 5. Size versus age relation for ETGs at $z = 0$, with galaxies colour-coded by their stellar mass. Solid lines show the average trends in these mass bins and are labelled with the $\log_{10}(M_*/M_{\odot})$ values (Kroupa IMF). Open and filled symbols show objects that have lower and higher f_{DM} , respectively, than the average for other galaxies of similar mass and age.

but for any given mass there is a clear anticorrelation (particularly for the most massive galaxies), i.e. older galaxies have on average smaller effective radii. We quantify the size–age dependency by fitting a log–linear relation in each of the mass bins. The slopes range from -0.005 to -0.030 , where the normalization increases with mass in accordance with the typical $R_{\text{eff}}-M_*$ relation (see e.g. Paper I). These trends appear to be independent of the galaxy subtype (ellipticals versus SOs), which is a crucial point because one might otherwise speculate that the younger objects are SOs with systematically different properties than ellipticals.

These size–age trends might be driven by higher densities of DM and gas at earlier epochs of galaxy formation (see references in Section 1). However, as with the age–mass correlation, we will for now set aside a physical interpretation of the size–age correlation and simply consider it as a nuisance factor that must be treated appropriately in our later toy models of DM content. These size–age results are not meant to be universally applicable to ETGs but are a characterization of this particular data set as needed for our self-consistent galaxy toy models.

The size–age trends clearly go in the right direction to explain the f_{DM} –age trends: older galaxies could exhibit lower f_{DM} simply because their R_{eff} values are smaller, so smaller volumes of their DM haloes are probed. In principle, one might be able to see directly from the data whether f_{DM} is more directly coupled to age or to R_{eff} , and to this end we have highlighted the high- and low- f_{DM} galaxies in Fig. 5. However, it is not immediately obvious whether R_{eff} or age is the stronger factor, probably because of the scatter in the data. Instead, in the next section we will adopt a model-dependent approach, attempting to make sense of the observed trends in the context of standard DM+galaxy models.

Before continuing, in Appendix B we carry out a number of cross-checks on the robustness of our basic f_{DM} –age result. First, comparing independent constraints from the literature we find some support for an f_{DM} –age anticorrelation. However, there may be

disagreement in the *residual* f_{DM} -age trend after accounting for the R_{eff} -age relations (which we examine further in the following sections). Such differences illustrate the need for a further in-depth study of this topic. We also check in Appendix B the effects of our stellar population modelling assumptions, finding that our main conclusions are fairly insensitive to these.

Besides the systematic uncertainties, there is the issue of the statistical errors being correlated, in the sense that an error in age correlates with an error in Υ_* and produces an artificial f_{DM} -age anticorrelation (see the error ellipse in Fig. 4). To gauge this effect, we have simulated mock data sets with random errors applied to the ages, propagated these to errors on Υ_* and then regenerated the f_{DM} -age plot. The artificial trend generated by the errors turns out to change the slope by only $\sim -0.005 \text{ Gyr}^{-1}$ (compared to the observed slope of $\sim -0.04 \text{ Gyr}^{-1}$). In the rest of this paper, we will continue to assume that our empirical results are correct and go on to examine the implications for galaxy structure and formation.

3 DARK MATTER IMPLICATIONS

Given the preceding inferred DM trends, we begin here with interpretations of the mass dependencies, constructing Λ CDM-based toy models in Section 3.1. We consider DM density profiles in Section 3.2 and comparisons to related literature results in Section 3.3.

3.1 Λ CDM halo models and trends with mass

We now construct a series of toy models of galaxies including DM haloes based on Λ CDM theory, as previously done in Napolitano et al. (2005) and in Paper I. With this approach, we are not modelling individual galaxies but generating typical representatives in the space of mass and age. Each model is spherically symmetric and is comprised of a stellar spheroid and a DM halo. The spheroid has a Sérsic density distribution with an n -index as described in Paper I and an $R_{\text{eff}}-M_*$ -age dependence taken from the log-linear fits in Fig. 5.

The DM halo has an NFW density profile (Navarro, Frenk & White 1997) following a typical concentration-virial mass ($c_{\text{vir}}-M_{\text{vir}}$) relation.⁵ The DM halo is optionally ‘adiabatically contracted’, which is an approximate way to model the expected drag of dissipatively infalling baryons on the surrounding DM, producing a halo with a higher central DM density than in collisionless N -body simulations.

The two adiabatic contraction (AC) recipes we use are the most common ones, from Blumenthal et al. (1986, hereafter B+86) and Gnedin et al. (2004, hereafter G+04), with the latter providing a weaker (and probably more realistic) effect. Other recipes are also available (e.g. Sellwood & McGaugh 2005; Gustafsson, Fairbairn & Sommer-Larsen 2006), but a more important caveat is that the entire slow, smooth-infall AC scenario may not be correct for forming cosmological structures – particularly the ETGs which may have experienced relatively violent, clumpy assembly histories. Current high-resolution simulations do suggest that the standard AC models are not accurate in detail and that the baryonic effects on DM haloes might be highly variable and stochastic (Abadi et al. 2009; Pedrosa, Tissera & Scannapieco 2009, 2010; Duffy et al. 2010; Tissera et al.

⁵ As in our previous papers, we have adopted a theoretical relation based on a DM power spectrum normalization of $\sigma_8 = 0.9$ (Bullock et al. 2001). We have also derived models based on a more recent estimate of $\sigma_8 = 0.8$ (Macciò, Dutton & van den Bosch 2008), which predict lower central DM content, but only at the level of changing f_{DM} by ~ -0.05 .

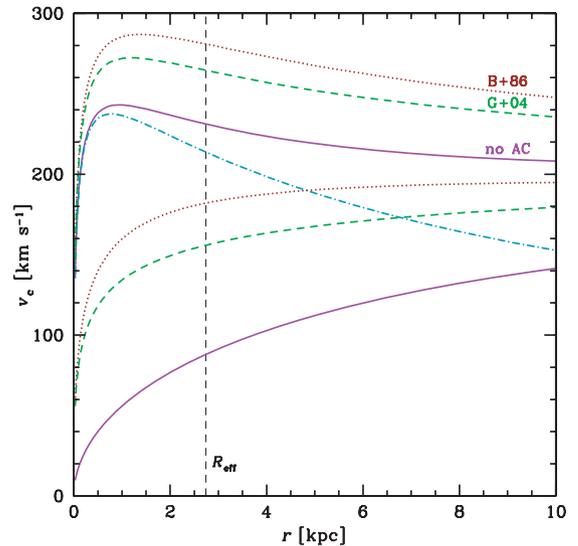


Figure 6. Circular velocity profiles with radius, $v_c(r) \equiv GM(r)/r$, for cosmologically motivated galaxy toy models. The stellar and halo masses are fixed to $\log_{10}(M_*/M_\odot) = 10.85$ and $\log_{10}(M_{\text{vir}}/M_\odot) = 12.14$, respectively, corresponding to $\epsilon_{\text{SF}} = 0.3$. Three cases are considered with different AC recipes and are indicated by curves of different colours and line styles, labelled by each recipe. For each case, a mass decomposition is shown, with the DM contribution as the bottom three curves, the stellar contribution as the intermediate steeply declining blue dot-dashed curve (the same for all cases) and the total as the top three curves. A vertical dashed line shows the value of the effective radius adopted.

2010). For simplicity, we will consider our AC and no-AC models as representing two extreme models of galaxy formation, where the baryons have alternatively strong or weak net impact on the DM halo (cf. Lackner & Ostriker 2010).

In order to match the galaxies to their haloes, we must relate their stellar masses to their total (or virial) halo masses. We can parametrize this connection by assuming that the cosmological baryonic fraction, $f_{\text{bar}} = \Omega_{\text{bar}}/\Omega_{\text{m}} = 0.17$ (Spergel et al. 2007), is conserved within each halo. The stellar mass is then given by an efficiency of star formation ϵ_{SF} , so that $M_* = 0.17\epsilon_{\text{SF}}M_{\text{vir}}$. We use a series of plausible values $\epsilon_{\text{SF}} = (0.1, 0.3, 0.7)$ (e.g. Zheng, Coil & Zehavi 2007; Brown et al. 2008; Conroy & Wechsler 2009; More et al. 2009; Guo et al. 2010; Lagattuta et al. 2010; Moster et al. 2010; Schulz, Mandelbaum & Padmanabhan 2010, hereafter S+10).

We also adopt the Kroupa IMF as our fiducial choice but consider the impact of Salpeter or Chabrier IMFs. Our model has effectively three free ‘parameters’: ϵ_{SF} , AC recipe and IMF choice. Fig. 6 shows an example of a series of galaxy toy models for fixed stellar and halo masses and varying AC assumptions. It can be seen that the DM content inside $1 R_{\text{eff}}$ varies most by changing from a no-AC model to an AC model. The choice between the B+86 and G+04 recipes makes only a small difference on these scales.⁶ This model may seem simple, but as far as we know *it has never been applied to the FP before*, except for the related study of S+10 as we discuss later.

We now begin by deriving the predicted DM fraction within R_{eff} as a function of mass, in bins of constant age (i.e. fixing the R_{eff}

⁶ The AC recipe differences become more important at smaller radii, where the detailed treatment of the stellar mass profile also becomes more important, e.g. Hernquist vis-a-vis Sérsic models.

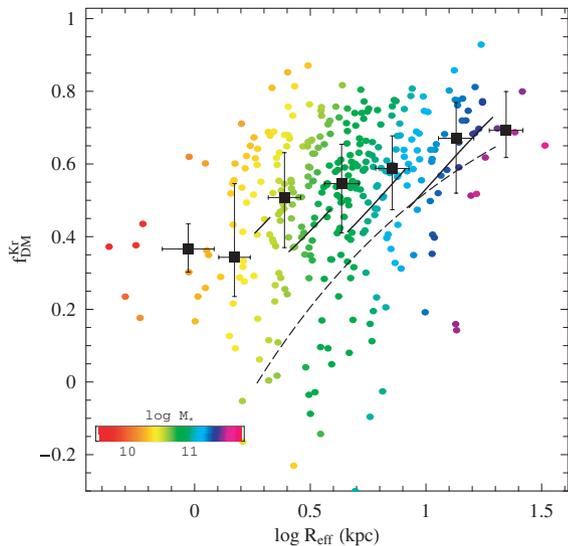


Figure 7. Central DM fraction versus effective radius. Black boxes with error bars show the average trend in bins, while smaller points show individual galaxies, coloured by the (smoothed) stellar mass according to the inset colour bar. f_{DM} correlates with both M_* and R_{eff} which are themselves correlated quantities, but the tighter trend in this figure compared with Fig. 2 indicates that R_{eff} is the more important parameter driving the FP tilt. Sample Λ CDM toy models in bins of constant M_* (assuming $\epsilon_{\text{SF}} = 0.1$ and G+04 AC) are shown as solid curves: an increase of f_{DM} with R_{eff} is naturally expected in these models. The dashed curve shows the trend found by Humphrey & Buote (2010).

values for a given mass), and with a number of different assumptions on ϵ_{SF} and AC recipe. The results are shown in Fig. 2 where it can be seen that without AC, the models generally underpredict f_{DM} . The AC models on the other hand match the data much better *on average*, including reasonable reproductions of the trend with M_* ; there is no strong preference between the B+86 and G+04 recipes. (The data clearly show a large variation around the models, part of which is just observational noise and part of which is an important systematic variation that we will examine in Section 4.) These conclusions are IMF dependent: with a Salpeter IMF, the f_{DM} data points all shift downwards to lower values, and the no-AC model is preferred on average⁷ (although there are suspicions of Salpeter not being a valid alternative; see Section 2.1).

For any fixed IMF, we see that it is a fairly generic expectation in a Λ CDM context for f_{DM} to increase systematically with M_* and to produce a DM-driven tilt in the FP. The reason is the aperture effect mentioned earlier: the scalelengths of galaxies’ stellar parts change more rapidly than for their haloes, causing the R_{eff} region to encompass increasingly large amounts of DM. This interpretation is supported by the $f_{\text{DM}}-R_{\text{eff}}$ trends in Fig. 7 and by various FP analyses in the literature based on simulations of galaxy formation (Boylan-Kolchin, Ma & Quataert 2005; Oñorbe et al. 2005, 2006; Boylan-Kolchin et al. 2006; Robertson et al. 2006; Covington 2008; Hopkins, Cox & Hernquist 2008). A similar point from an observa-

⁷ In Paper I we focused on the Salpeter IMF and stated erroneously that the consistency of the data with our initial no-AC toy models was independent of the IMF, when examining the relation between the mass and mean DM density within $2R_{\text{eff}}$. This relation for the *data* is indeed unaffected by IMF changes since both the implied masses and the densities vary with the IMF. However, we neglected considering the effects on the *models*, where only the predicted densities change (because the R_{eff} values at a fixed mass are changed; see further discussion in Section 3.2).

tional basis is made by fig. 5 of Humphrey & Buote (2010), and we include their results for comparison in Fig. 7; here the absolute values of f_{DM} cannot be readily compared since their R_{eff} values come from the K band and should be systematically different from ours in the B band, but their general trend with size seems comparable to ours. The role of R_{eff} in more model-dependent DM scalings is considered further in Section 3.2.

Systematic trends in ϵ_{SF} can affect the slope of $f_{\text{DM}}-M_*$ in either direction, but the effect is fairly weak. Even for a factor of 7 change in assumed halo mass, f_{DM} changes only by ~ 0.1 ; i.e. large changes in overall halo mass do not translate to large changes in the central DM content. To cancel out the DM-induced tilt, ϵ_{SF} would have to systematically increase with mass, which goes in the opposite direction to conventional findings, wherein the most massive systems have the highest virial M/L . If the FP tilt turns out to *not* be driven by f_{DM} trends (as found in Trujillo et al. 2004; Jun & Im 2008; Allanson et al. 2009), then this would be inconsistent with vanilla Λ CDM expectations (i.e. where central DM densities scale slowly relative to the observed stellar densities).

3.2 Dark matter profiles

We next consider some interesting implications for DM density profiles with radius. It is notoriously difficult to determine the DM content in the centres of ETGs, and a scant few studies to date have been able to compare the empirical DM properties to Λ CDM theory. T+09 and Paper I (Fig. 14) have found mean central densities for DM haloes that scale with mass roughly as expected for NFW haloes. Here we add the finding that the DM content also scales (to first order) with galaxy size as expected.

We will see in Section 4.1 that much of the observed $f_{\text{DM}}-age$ anticorrelation can be traced to the $R_{\text{eff}}-age$ anticorrelation, implying that our Λ CDM toy models are at least roughly correct. This point is illustrated more clearly in Fig. 8, where for one example mass bin, the stellar mass has been subtracted from the total mass, and the residual DM content is quantified as the mean value within R_{eff} , $\langle \rho_{\text{DM}} \rangle$.⁸

By measuring the DM content of individual galaxies with similar masses at different radii, we are effectively able to map out a mean DM density profile over a factor of 3 in radius. This trick of studying composite profiles is formally valid only if the binned galaxies all have approximately the same DM profile – as might be the case for a no-AC NFW model. The condition is clearly violated for models that include AC since the contraction varies among the differently sized galaxies. However, the impact of this non-homology is reasonably small as discussed further below, such that we can get an approximate idea of the DM slopes as well as more firmly test the null hypothesis of a universal NFW profile.⁹

⁸ Note that this is different from the analysis of Graham et al. (2006b) who looked at $\langle \rho_{\text{DM}} \rangle$ inside the effective radius of the *DM* rather than of the *stars*. Their derived density slopes reflect the steeply declining outer regions of the DM haloes, while ours involve the central regions.

⁹ One might suspect that at some level we are getting out what we are putting in, since our default dynamical models used in deriving f_{DM} assume an $\alpha = 2$ total density profile in order to extrapolate the σ_0 measurements to R_{eff} . However, we have confirmed that using the alternative constant- M/L profile yields similar results (α unchanged for the Salpeter IMF and *steeper* by $\Delta\alpha \sim 0.1-0.2$ for the Kroupa IMF). A galaxy sample with direct mass constraints nearer to R_{eff} (e.g. C+06; Auger et al. 2009; Graves 2009) would be ideal for pursuing the density profile analysis more robustly. We have also checked that these results are not sensitive to the stellar population models (Appendix B3).

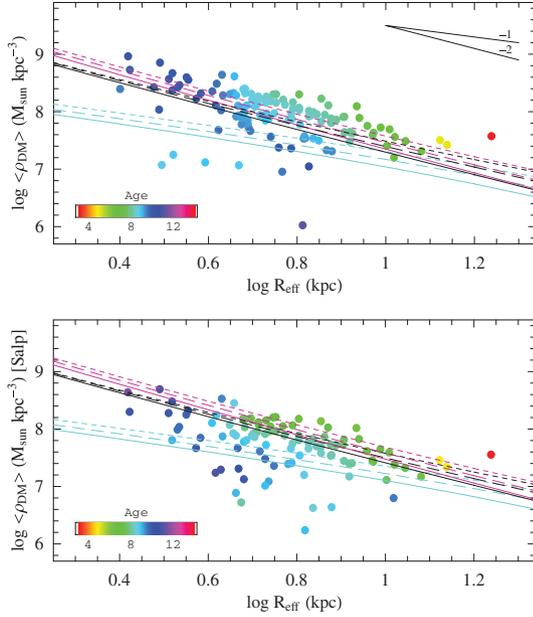


Figure 8. Relation between DM density within $1 R_{\text{eff}}$ and the R_{eff} value for ETGs in one mass bin. The top panel shows the case of a Kroupa IMF, where the mass corresponds to $\log_{10}(M_*/M_{\odot}) \sim 11.25$, and the bottom panel shows a Salpeter IMF for the same galaxies [$\log_{10}(M_*/M_{\odot}) \sim 11.4$]. Data for individual galaxies are shown as points coloured by their ages (after smoothing; see the inset colour bars). Λ CDM-based models are shown as curves: the lower blue curves are for no-AC, the middle black ones for G+04 AC and the top red ones for B+86 AC. For each set of AC models, fixed values of $\epsilon_{\text{SF}} = 0.7, 0.3, 0.1$ are shown as solid, long-dashed and short-dashed curves, respectively. Lines in the upper right of the top panel illustrate power-law slopes of $\alpha = -1$ and -2 . The scatter in points (including the apparent plume to very low densities) could be entirely due to observational errors (see Fig. A1).

The first point to notice from Fig. 8 is that the data for $\langle \rho_{\text{DM}} \rangle$ and f_{DM} (see Fig. 7) have opposite trends versus R_{eff} , which turns out to be crucial evidence for cuspy DM haloes on scales of ~ 2 – 15 kpc. As R_{eff} increases for a fixed galaxy mass and halo profile, the amounts of both enclosed DM and f_{DM} increase. On the other hand, the local DM density of a cuspy halo decreases with radius, so a larger R_{eff} means a smaller $\langle \rho_{\text{DM}} \rangle$.

More quantitatively, we may consider a general power-law density scaling for the central DM halo, $\rho_{\text{DM}}(r) \sim r^{-\alpha}$. For $\alpha < 3$, $M_{\text{DM}}(r) \sim r^{3-\alpha}$ and therefore $\langle \rho_{\text{DM}} \rangle \sim R_{\text{eff}}^{-3} M_{\text{DM}}(r = R_{\text{eff}}) \sim R_{\text{eff}}^{-\alpha}$. For an uncontracted NFW halo, $\alpha \sim 1$ at the smallest radii, but near R_{eff} in our toy models, $\alpha \sim 1.1$ – 1.3 (steeper for higher masses because of their larger radii; see also Graham et al. 2006a). AC effectively makes the DM even cuspiest, with $\alpha \sim 1.6$ – 1.8 for the G+04 recipe and $\alpha \sim 1.8$ – 1.9 for B+86 (see also Fig. 6). The effect of constructing composite profiles from non-universal haloes is for the apparent slopes to be steeper by $\Delta\alpha \sim 0.3$. This offset can be subtracted from the observational results to estimate the true slopes (see below), while more rigorous tests are made by constructing composite toy models for comparison to the data.

Our observational results for a Kroupa IMF are illustrated by Fig. 8 (top). The vast majority of the data points in this fixed mass bin do appear to clump around a common DM mass profile, with a composite slope of $\alpha \sim 1.9$ (so we may infer a true slope of $\alpha_0 \sim 1.6$). A small minority of galaxies may follow a shallower trend, with a residual trend of DM fraction versus age that is barely visible by the colours of the data points (to be discussed in Section 4.1).

The results for the other mass bins are similar, with $\alpha \sim 1.7$ – 2.1 , and with suggestions of a shallower slope with increasing mass (see Fig. 9), although it is hard to tell for sure since the exact slopes are sensitive to the details of the fitting procedure.¹⁰

We also construct NFW-based model predictions as in Section 3.1 and show them in Fig. 8. The steep slope for the majority of the points, along with the shallower slope with increasing mass, turns out to correspond nicely to a range of NFW+AC models,¹¹ with a minority of points consistent with uncontracted NFW haloes. The model interpretations are somewhat degenerate to systematic variations of ϵ_{SF} with R_{eff} (e.g. S+10 suggest an anticorrelation, which would imply that the intrinsic slopes are steeper than they appear). However, as seen in Fig. 8, the ϵ_{SF} effects are fairly small and overshadowed by our fitting uncertainties anyway.

The $\langle \rho_{\text{DM}} \rangle$ – R_{eff} relation is not a knock-down confirmation of Λ CDM but intriguingly it does seem to weigh against alternative models. If there is no DM, then of course we should have $f_{\text{DM}} = \langle \rho_{\text{DM}} \rangle = 0$, although there might be a systematic error in the mass analysis, e.g. with the IMF. In this more general case we expect $\alpha \sim 3$ – and the apparent f_{DM} to be constant with age and R_{eff} – in strong contradiction to the data.

The same argument may apply to, and pose a challenge for, alternative gravity theories that seek to explain observational mass discrepancies at large galactic radii, usually in spiral galaxies (e.g. Brownstein & Moffat 2006; Frigerio Martins & Salucci 2007). It is beyond the scope of this paper to consider any of these theories in detail, but we will briefly comment on the most well-studied case of modified Newtonian dynamics (MOND; e.g. Tiret et al. 2007; Sanders & Land 2008; Ferreras et al. 2009). In this formalism, there is a characteristic acceleration scale $a_0 \sim 1.2 \times 10^{-8} \text{ cm s}^{-2}$, above which mass discrepancies should be weak or non-existent. In the central regions of the ETGs studied here, the accelerations are large ($\sim 4a_0$) and the dynamical mass should agree well with the inferred stellar mass.

Our observed mass discrepancies, along with the striking R_{eff} dependencies, would at first glance appear to rule out MOND without DM in ETGs, which might be related to similar results in galaxy groups and clusters (e.g. Sanders 2003; Pointecouteau & Silk 2005; Angus, Famaey & Buote 2008; Richtler et al. 2008). However, we cannot make a definitive statement without carrying out a careful MONDian analysis (see also fig. 19 in Gerhard et al. 2001, hereafter G+01; preliminary findings in Romanowsky 2006; and forthcoming results in Cardone et al. 2010). In particular, the systematic uncertainties in our Υ_{dyn} and Υ_* estimates must be taken into account, along with possible fine-tuning of the ‘interpolation function’ between the Newtonian and MONDian regimes (cf. McGaugh 2008).

If there are DM haloes with cored profiles (e.g. Burkert 1995; Saxton & Ferreras 2010), then f_{DM} would still increase with R_{eff} , but $\langle \rho_{\text{DM}} \rangle$ should be constant ($\alpha \sim 0$). Again, using the wrong IMF could contribute an $\alpha \sim 3$ effect, which in combination with a constant DM core could produce a shallower than $\alpha \sim 3$ trend – but

¹⁰ A fair amount of attention in the literature has been focused on the slopes of *total* mass profiles rather than the DM slopes discussed here (e.g. Jiang & Kochanek 2007; Koopmans et al. 2009; Nipoti, Treu & Bolton 2009; Humphrey & Buote 2010). We have not made comparisons in detail, but our toy models imply total slopes at R_{eff} of $\alpha \sim 1.8$ – 2.2 , becoming shallower with increasing mass. This trend is driven mostly by the stellar density becoming shallower and is affected little by AC.

¹¹ The overall DM density shows a tendency to increase with mass, relative to fixed- ϵ_{SF} models (see Paper I) so the overall connection between AC and mass is not yet clear.

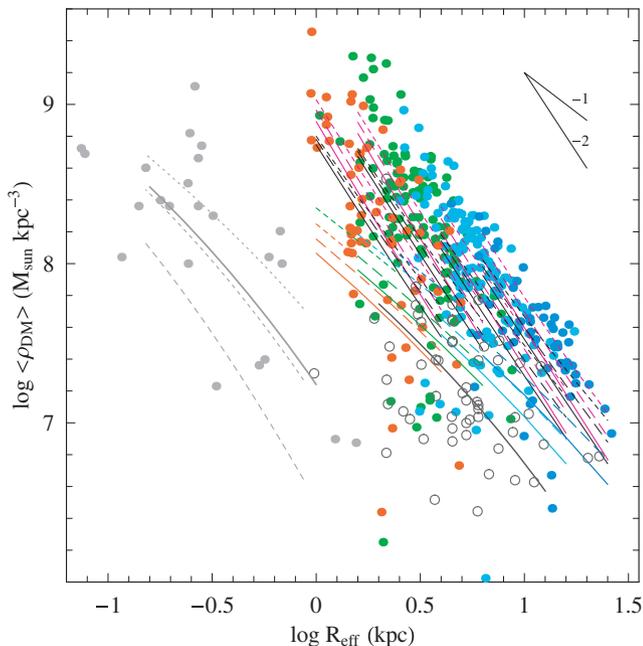


Figure 9. As Fig. 8, but with all ETG mass bins plotted together (using a Kroupa IMF; see Fig. 5 for colour key). In addition, data are shown for dwarf spheroidal galaxies (grey filled circles on the left; Walker et al. 2009, 2010a) and for late-type disc galaxies (black open circles; McGaugh et al. 2007). The Λ CDM toy model curves are as in Fig. 8, with the addition of sample grey and black model curves for the dwarf and disc galaxies. The dwarf models include a no-AC halo sequence with $\epsilon_{\text{SF}} = 0.04, 0.004, 0.0004$, along with a case of $\epsilon_{\text{SF}} = 0.004$ with AC (G+04 recipe; solid curve). The spiral model uses a no-AC halo with $\epsilon_{\text{SF}} = 0.3$.

experimenting with IMF adjustments, we cannot recover a constant-core trend.

Any of these scenarios could still be salvaged with epicycles, e.g. with systematic errors that correlate with galaxy size, but Ockham’s razor suggests that the first conclusion should be to prefer a standard cuspy DM model. Note that detailed dynamical models of individual galaxies have difficulty uniquely decomposing the stellar and DM mass profiles and thereby distinguishing between cored and cusped DM haloes (cf. T+09; Napolitano et al. 2009, hereafter N+09). Our approach that includes the modelling of stellar populations in a large galaxy sample provides the first clear (albeit indirect) constraint on the central DM profiles of ordinary ETGs.

Before moving on, we consider the Salpeter IMF alternative, with data and models for the sample mass bin shown in Fig. 8 (bottom). The data are seen to typically lie in between the AC and no-AC models, with a slope of $n \sim 1.7$. Unlike the Kroupa case where both the f_{DM} amplitude and the DM density slope match the AC models nicely, the Salpeter case shows less overall consistency with a simple Λ CDM model, although a caveat is that the frequent cases with apparent negative DM densities can make the results difficult to interpret.

3.3 Context and comparisons

There are several recent observational analyses of DM in galaxies that are relevant to our current study: we will make some comparisons first to studies of ETGs and then to other galaxy types. We begin with two papers that are very similar in spirit to ours, as they combine estimates of total masses and stellar-population-based

masses to derive central DM constraints in large samples of elliptical galaxies and to compare these to Λ CDM-based toy models.

T+10 studied a sample of strong gravitational lenses, using a combination of lensing and dynamics to derive the total masses. They found that a Salpeter IMF is generally consistent with a no-AC model, as did we. They did not consider AC models explicitly, and it seems very likely that (as we have found in this paper) lower mass IMFs would be favoured if AC were included.

These authors also found indications of systematic changes with galaxy mass that could be interpreted as a change either in IMF or in the DM density slope such that it *steepens* with mass (α increasing). The latter effect as they mention could be due to AC becoming stronger with mass and appears to go in the opposite direction of our results. It is beyond the scope of this paper to track down the reasons for this difference (their study did probe somewhat different regimes of radius and mass).

S+10 derived their total central masses from stellar dynamics and used weak gravitational lensing to derive the total halo masses and concentrations (which we have had to *assume* in our models¹²). Adopting a Kroupa IMF, they found relatively high central DM masses which they showed to be nicely consistent with G+04 AC models when breaking the sample down into several bins in mass and in size (the latter test being implicitly equivalent to our plots of $\langle\rho_{\text{DM}}\rangle$ versus R_{eff} as an alternative test of the DM profiles). Their large-radius constraints also allowed them to determine a correlation between the size and virial mass at fixed M_* , which in the context of our formalism implies an anticorrelation between R_{eff} and ϵ_{SF} . If their sample has an R_{eff} –age anticorrelation as in our sample, then this would imply an ϵ_{SF} –age correlation. We will come back to this point in Section 4.1.

In addition to the qualitative consistency of the S+10 results with ours, their quantitative DM fractions of typically $f_{\text{DM}} \sim 0.6$ – 0.7 can be compared to ours over the same mass range [$\log_{10}(M_*/M_{\odot}) \sim 10.8$ – 11.6] where we find $f_{\text{DM}} \sim 0.4$ – 0.6 . This small difference might be accounted for by their sample selection of only round ‘central’ galaxies, which are plausibly more DM-rich than flattened, satellite galaxies. These authors do note that their results seem to be somewhat high compared to other literature studies with $f_{\text{DM}} \sim 0.3$ – 0.5 (C+06; Gavazzi et al. 2007; see also G+01; T+09).

Jiang & Kochanek (2007) analysed a sample of galaxies using strong lensing and stellar dynamics, and Λ CDM mass models similar to ours. Using priors on either Υ_* (equivalent to a Kroupa IMF) or total masses from weak lensing (Gavazzi et al. 2007), they find that an AC model with $\epsilon_{\text{SF}} \sim 0.3$ is preferred. The strong+weak lensing analysis of Gavazzi et al. (2007), on the other hand, found consistency with a no-AC model.

T+09 constructed detailed dynamical models of Coma cluster ETGs, where Υ_* and Υ_{dyn} were determined simultaneously from the dynamics, assuming a parametrized DM model. They found relatively high DM densities (similar to ours), which they compared to cosmological semi-analytic models (using no-AC models for spirals as an additional constraint and adopting cored DM halo models for the ETGs which in this context tend to *under*-estimate the central DM densities). They concluded that B+86 AC haloes were preferred over no-AC.

Humphrey et al. (2009) used X-ray emission from hot gas surrounding a small sample of massive ETGs to constrain their mass

¹² The direct mass–concentration results of S+10 follow well the theoretical expectations for a *Wilkinson Microwave Anisotropy Probe 5* cosmology (Macciò et al. 2008), although one might expect ETGs to be systematically biased towards higher concentration haloes.

profiles using Λ CDM-motivated models. With a Kroupa IMF, they found that no-AC models were preferred (although with somewhat high concentrations). One of their sample galaxies was also the subject of a detailed dynamical study using stars and globular clusters (Shen & Gebhardt 2010), whose mass results were systematically higher and might very well be consistent with an AC model. This appears to be part of a significant unresolved tension between dynamical and X-ray mass results. On the other hand, other X-ray studies find surprisingly high DM concentrations for galaxy groups (Buote et al. 2007; Duffy et al. 2008), which should generally overlap with the massive end of our ETG sample and might support our finding that AC models are preferred.

Lackner & Ostriker (2010) used toy models for ‘dissipationless’ and ‘dissipational’ ETG formation which should be roughly equivalent to our no-AC and AC models. Adopting the C+06 Υ_* values assuming a Kroupa IMF, they found that both types of models predicted too much DM in the galaxy centres. This disagrees with our conclusion that AC models predict about the correct amount of DM, and no-AC models do not predict enough. Although there are some differences between our observational results and those of C+06 (see Appendix B3), it may also be that Lackner & Ostriker did not allow for enough freedom in their halo masses and concentrations to fit the data. Our own preliminary analysis of the C+06 data suggested that a G+04 AC model would work well on average (Romanowsky 2006).

We next consider our ETG results in relation to other galaxy types, showing density–size relations for different ETG mass bins, along with data from other galaxy types in the literature, plotted together in Fig. 9. We first of all consider 56 late-type galaxies from McGaugh et al. (2007), where dynamical masses were measured by classical rotation curves and stellar disc masses by various methods including the modelling of stellar populations. We take results from the latter (with a Kroupa IMF) and then infer the DM densities within R_{eff} , which we define as $1.68 \times R_d$, where the disc scalelength R_d is taken from various literature sources such as McGaugh (2005).

We plot these late-type galaxies in Fig. 9 and see that their DM haloes are approximately five times less dense than those of early types at the same radii. Qualitatively similar conclusions were reached by G+01 and T+09, while we add the observation that the DM density *slopes* are different: the late types have shallower slopes than the early types ($\alpha \sim -1$ versus $\alpha \sim -2$). The ETGs could be brought into closer consistency with the late types if a Salpeter IMF were adopted for the former, but if anything a reverse IMF trend might be expected a priori (as we will discuss further in Section 4.4).

We also show in Fig. 9 a sample Λ CDM model curve for the late types, for the case of $\log_{10}(M_*/M_\odot) = 10$, $\epsilon_{\text{SF}} = 0.3$ and no AC. This curve is very similar to the equivalent one for the lowest mass bin of the ETGs, so for clarity we do not show additional model curves for the late types (whose AC model curves should differ slightly from the ETG curves). The no-AC model appears to be preferred for the late types,¹³ echoing the conclusions of many other studies (Kassin, de Jong & Weiner 2006; Dutton et al. 2007; Gnedin et al. 2007; McGaugh et al. 2007; Xue et al. 2008) and contrasting with the inference of AC for the early types when the same IMF is adopted.

¹³ However, McGaugh et al. (2007) among others found that Λ CDM no-AC models do not match the data in detail, with cored DM haloes being preferred. Note that if these haloes do have constant-density cores, then there is less concern about selecting the fairest radius for comparing their densities to ETG haloes.

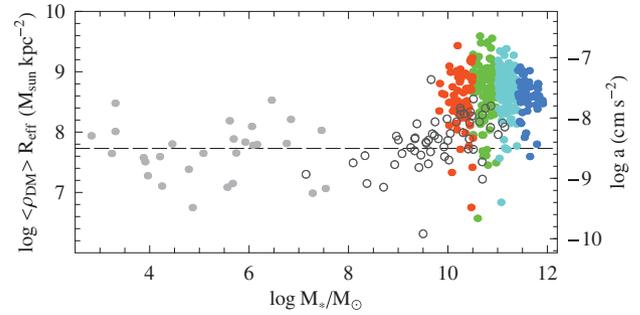


Figure 10. Projected density of DM in the central regions of galaxies versus stellar mass. The data and symbols are as in Fig. 9. For comparison, the projected DM density result of Donato et al. (2009) is shown as a dashed line (see also Gentile et al. 2009; our density calculation is not exactly equivalent). The equivalent acceleration scale is provided on the right axis; the characteristic scale in MOND is $\log a_0/(\text{cm s}^{-2}) \sim -7.9$. Note that these quantities are all expressing the projection of the 3D density within R_{eff} rather than a true total surface density, which would require model-dependent extrapolations outside the regions probed by the data.

We next consider dwarf spheroidals belonging to the Milky Way, from the compilation of Walker et al. (2009) (cf. the alternative compilation of Wolf et al. 2010). This study extended upon earlier work that found a common DM mass within a fixed physical radius (Strigari et al. 2008) and suggested that these systems share a common ‘universal’ DM profile with $\alpha = 1.6 \pm 0.4$ (see also Peñarrubia et al. 2010). Plotting their data in Fig. 9,¹⁴ we first note that the dwarf haloes do not appear to join up naturally with those of more massive galaxies (cf. Walker et al. 2010b), and we infer that sampling a limited range of galaxy types, masses and radii can readily produce the impression of a universal profile. This conclusion is bolstered by our overplotting of a Λ CDM toy model for the dwarfs [using a fiducial stellar mass of $\log_{10}(M_*/M_\odot) = 5.5$, Sersic index $n = 0.25$ and $\epsilon_{\text{SF}} = 0.004$; a model with AC is very similar because of the DM dominance]. For a factor of 10 variation in halo mass, the central DM densities change by only a factor of ~ 3 , which is less than the scatter in the data.

Finally, motivated by the work of Donato et al. (2009) who claimed a universal DM *projected* density for all galaxy types and masses (see also Kormendy & Freeman 2004), we calculate our own approximate DM surface density (Σ_{DM}) from the data points in Fig. 9 by multiplying (ρ_{DM}) by R_{eff} and showing the results versus stellar mass in Fig. 10. Although our results are calculated within galaxy R_{eff} rather than the DM core radius as in Donato et al., we reproduce their results reasonably well, not only finding a fairly constant surface density on average across a large mass range of dwarfs and late types, but also suggesting a systematic increase with mass as found by Boyarsky et al. (2009).

It is beyond the scope of this paper to consider Λ CDM theoretical predictions for the surface density in detail (cf. Zhao, Xu & Dobbs 2008; Boyarsky et al. 2009; Koposov et al. 2009; Li et al. 2009; Macciò, Kang & Moore 2009; Okamoto & Frenk 2009; Cooper et al. 2010; Stringer et al. 2010). However, the broad consistency of the uncontracted NFW models with the dwarf and spiral data in Fig. 9 suggests that the surface density result is not particularly surprising within the standard paradigm and does not necessarily imply cored DM haloes.

¹⁴ The (ρ) values in Walker et al. (2009) appear to be total mass including baryonic contributions, but this should be only an ~ 10 per cent effect.

We also find from Fig. 10 that the ETG galaxies violate the constant density scenario for the other galaxies by a factor of ~ 10 on average and a factor of ~ 5 in the same mass regime. This disagrees with the conclusions of Donato et al., and we note that their ETG results were based on a weak lensing analysis rather than central dynamical results equivalent to those used for the other galaxies. Other ETG dynamical results from the literature support ours (G+01; T+09; Boyarsky et al. 2009; see also the strong lensing analysis of Cardone & Tortora 2010).

In summary, we can synthesize the constraints on the central DM content of all types of galaxies assuming a universal Kroupa IMF, based on our study and on the literature. We conclude that the early-type and late-type galaxies are broadly consistent with simple Λ CDM predictions with and without AC, respectively.

The main exceptions to this emerging consensus are from X-ray studies (see above) and from the only compilation of large-radius dynamical results (N+09), which suggests very strong systematic AC differences among the ETGs. However, as discussed in Appendix B2, selection effects are suspected for the latter study. We postpone further discussion of the implications of our DM results for galaxy formation to Sections 4.3 and 4.4.

4 INTERPRETING TRENDS WITH AGE

Finally, we arrive at an analysis of the age trends. We provide an overview of the observations compared with basic Λ CDM toy models in Section 4.1. We then attempt to explain the age trends by the systematics of DM halo concentrations (Section 4.2), AC variations (Section 4.3) and non-universal IMFs (Section 4.4).

4.1 Overview of trends

As previously mentioned, the variations in f_{DM} at fixed M_* appear to correlate with age, and comparison with the models in Fig. 2 suggests that plausible systematic variations in ϵ_{SF} with age would not be enough to account for the trends in the data. To see this more clearly, we construct models of f_{DM} versus age in fixed mass bins, comparing these with the data in Fig. 11 (keeping in mind that we expect some scatter in the data points from observational errors and from the necessary inclusion of a range of masses in each mass bin).

Both data and models agree that f_{DM} decreases with age at fixed mass, which can be seen as a natural consequence of the R_{eff} -age trends.¹⁵ However, the data in every mass bin show even steeper age trends than predicted by the models – particularly if we adopt models with AC, a process that couples the DM to the baryons and partially counteracts the trend for small stellar sizes to produce small f_{DM} (i.e. relative to uncontracted haloes, it leads to a *shallower* trend of f_{DM} with age). The difference in slope between the data ($df_{\text{DM}}/d\text{age} \sim -0.04 \text{ Gyr}^{-1}$) and the simple models ($\sim -0.01 \text{ Gyr}^{-1}$ for AC and $\sim -0.02 \text{ Gyr}^{-1}$ for no-AC) suggests a systematic variation with the age of the DM itself. After correcting for the effects of R_{eff} (in an inevitably model-dependent way), we support our initial conclusion from Section 2.2 that the central DM content is connected more closely to age than to mass, since f_{DM} varies with age by up to ~ 0.3 relative to the model curves and by only ~ 0.1 with mass.

¹⁵ Shankar & Bernardi (2009, hereafter SB09) also found a similar trend of excess dynamical mass with age, which they stated as driven mainly by the R_{eff} variations. However, they did not demonstrate this conclusion explicitly nor did they connect to specific DM models.

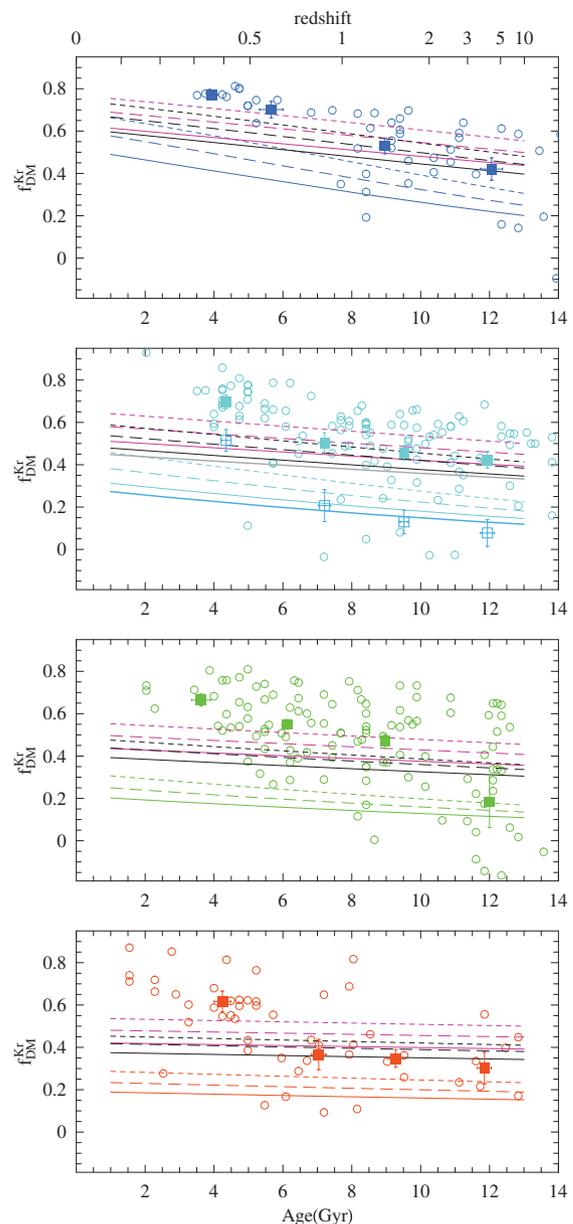


Figure 11. Relation between the central DM fraction and stellar age for ETGs in four mass bins: $\log_{10}(M_*/M_{\odot}) = 11.6, 11.2, 10.8, 10.3$ from top to bottom. Data for four age bins are plotted as points with error bars as the mean and standard deviation for f_{DM} . Λ CDM-based model predictions are shown as curves (see the Fig. 8 caption for details; the model trends with age are driven by R_{eff} variations). The default Kroupa IMF is used here, but for comparison the results for a Salpeter IMF are included in the $\log_{10}(M_*/M_{\odot}) = 11.2$ panel (second from top): the large open boxes show the binned data values and the lowest solid curve shows the $\epsilon_{\text{SF}} = 0.7$ no-AC model, which differs little from the equivalent Kroupa curve. Changes in the IMF affect the *data* much more than the *models* (where the IMF indirectly links the luminosity and halo concentration).

We illustrate the effect of age in a different way in Fig. 12 by plotting the mean central DM density versus age, for both data and models in one mass bin. Because of the anticorrelation of R_{eff} with age, the models predict a slight *increase* of $\langle \rho_{\text{DM}} \rangle$ with age. However, the data show a *decrease* (also apparent from the colours in Fig. 8), consistent with the results based on f_{DM} .

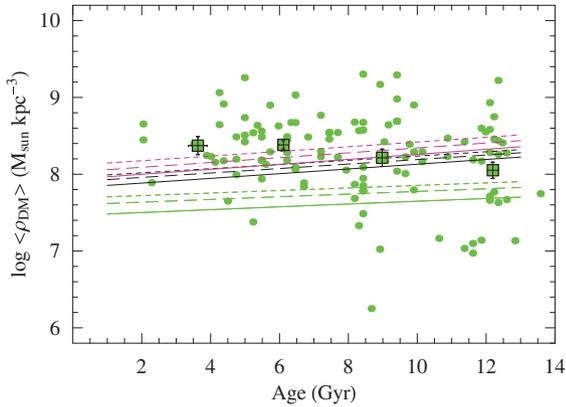


Figure 12. Mean DM density within $1 R_{\text{eff}}$ as a function of stellar age, in the $\log_{10}(M_*/M_{\odot}) = 10.8$ mass bin. Curves show model predictions as in Fig. 11, while small points show empirical results for individual galaxies. Large points with error bars show the average trend in age bins.

One obvious candidate to explain the DM–age trends would be ϵ_{SF} , such that the younger galaxies simply have more massive DM haloes overall. This scenario might be expected given cold-mode accretion at high- z (Dekel et al. 2009a), where star formation is thought to proceed very efficiently; subsequent halo growth would have to include infall material with high ϵ_{SF} in order to not overly dilute the net ϵ_{SF} at low- z . However, although it is plausible that an ϵ_{SF} –age correlation contributes to the f_{DM} trends,¹⁶ we suspect that it is probably not the dominant factor. By examining Figs 11 and 12, systematic variations of a factor of 10 or more in ϵ_{SF} at fixed M_* would be required. While a great deal of attention has been paid to associations of galaxies with halo masses *on average*, there has been much less work on the *scatter* in the trends. It is outside the scope of this paper to investigate the plausibility of order-of-magnitude variations, but the new work of More et al. (2010) suggests a scatter in virial M/L of factors of only ~ 2 – 5 for non-satellite ETGs (increasing with M_*).

4.2 Concentration variations

We next consider possible connections between DM profiles and age, such that the amount of DM within R_{eff} is not directly coupled to the total halo mass. First, our NFW toy models assumed the exact same mass–concentration relation for all galaxy haloes. However, this should really be a mean trend whose intrinsic scatter correlates with the redshift of halo collapse z_c , since halo density basically reflects the background mass density of the universe at the time of collapse.

This concentration–age systematic can immediately be seen to go in the wrong direction to explain the f_{DM} –age data, if we make the reasonable assumption that *stellar* age correlates at least weakly with the assembly age of the DM halo. Older galaxies should therefore have *higher* f_{DM} and $\langle \rho_{\text{DM}} \rangle$ because of their denser, earlier collapsing haloes.¹⁷

¹⁶ S+10 found for a sample of low- z ETGs that halo mass correlates with R_{eff} at fixed stellar mass, by a factor of ~ 1.5 – 2 in mass. If R_{eff} anticorrelates with age in their sample as in ours, then ϵ_{SF} correlates with age.

¹⁷ The stellar component could very well also be denser in older galaxies, but we are already accounting for this effect by including the R_{eff} –age relations in the toy models.

To illustrate this issue more quantitatively, we construct a modified version of our Λ CDM toy models. We begin with the simplest assumption that the age of a galaxy’s stars is directly associated with its halo’s z_c . This is clearly not correct in general since in the modern picture of ‘downsizing’, the star formation in massive galaxies largely precedes the DM assembly and vice versa in less massive systems (e.g. fig. 5 of Conroy & Wechsler 2009). However, it is beyond the scope of this paper to consider differential assembly histories at fixed final mass, so for now we will adopt our simple prescription as stated above.

To a good approximation, haloes of all masses and at all times ‘collapse’ with a fixed concentration of $c = K \sim 4$; the subsequent concentration evolution is basically mediated by the expansion of the universe. If a halo collapsed at $z = z_c$, then at $z = 0$ its concentration will be $c \simeq K(1 + z_c)$ with a scatter of $\Delta \log c \sim 0.05$ – 0.06 (see Wechsler et al. 2002; Macciò et al. 2008).

This line of argument implies that *all galaxies with the same stellar age have the same halo concentrations*, independent of mass. Certainly this is an oversimplified picture and might be undermined by late, minor gas-rich mergers biasing the inferred stellar ages. For now we will adopt this model as an ansatz and see where it leads us, with predictions illustrated by the model concentration–age bands in Fig. 13.

To derive estimates from the *data* for concentrations, we proceed as follows. First, we collect the individual galaxy results into bins of age and stellar mass, calculating the average f_{DM} values (similar to Fig. 11). We then fit our Λ CDM toy models (Section 3.1) to the f_{DM} data. Given a fixed IMF and DM profile (NFW with or without AC), the remaining free parameters are now ϵ_{SF} and c . We fix ϵ_{SF} to be constant in each mass bin and set it to a value that forces the average c (including all ages) to approximately agree with the

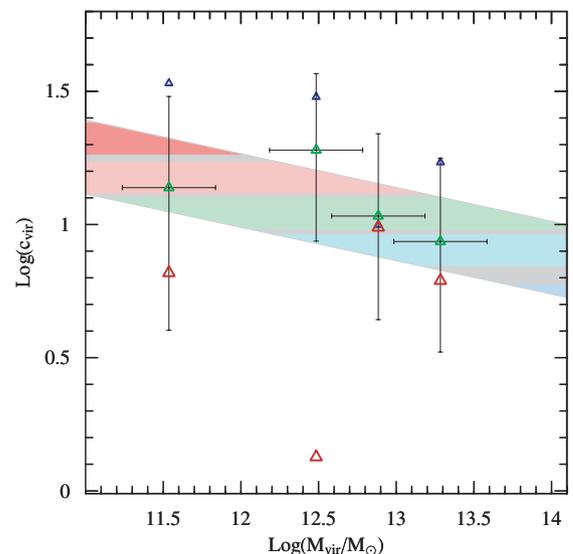


Figure 13. The concentration of NFW haloes, as a function of halo mass. The underlying grey diagonal band shows the expected average trend for 68 per cent of the haloes from Λ CDM simulations. The horizontal colour bands show halo ‘ages’ corresponding to collapse redshifts: 12, 11, 10, 8, 5 Gyr or $z_c \sim 4, 3, 2, 1, 0.5$, for red, pink, green, cyan, blue bands (top to bottom), respectively. The triangles with error bars show the data (assuming the Kroupa IMF) in four mass bins, with age sub-bins of 12, 8, 4 Gyr (red, green, blue colours, respectively). Sample error bars are included for the intermediate-age bins, showing the scatter in masses and the 1σ uncertainties in the best-fitting concentrations. These uncertainties mean that the one point with a very low concentration should not be taken too literally.

theoretically predicted mean value. Note that for models with AC, the c value denotes the ‘original’ concentration after correcting the observations for AC.

The results of this fitting exercise for a Kroupa IMF and G+04 AC prescription are overplotted as data points in Fig. 13. To reproduce the mean $c-M_{\text{vir}}$ relation, we find $\epsilon_{\text{SF}} = 0.25$ for the lowest mass bin and $\epsilon_{\text{SF}} = 0.1$ for the others; it is noteworthy that such a simple, plausible Λ CDM-motivated model can be constructed to agree with both central and global DM constraints for ETGs.

Now considering the age dependencies, we see that there is a clear trend for older galaxies to have less concentrated haloes – an effect that is diametrically opposed to our assumed trend for star formation to mimic halo collapse. If taken at face value, these concentration results would imply that ETGs on average formed their stars not long after their haloes collapsed ($z_{\star} \sim 1$ versus $z_c \sim 2$), but the late-collapsing haloes tended to form their stars very early ($z_{\star} \sim 4$ versus $z_c \sim 0.4$) while the early collapsing haloes formed their stars very late ($z_{\star} \sim 0.4$ versus $z_c \sim 5$), although these latter trends are more apparent for the low-mass galaxies. We are not aware of any physical mechanism that would produce such counter-intuitive variations *at a fixed mass*, and we suspect that concentration effects do not explain the observed DM–age trends. These conclusions are similar when adopting a Salpeter IMF with no AC.

4.3 Adiabatic contraction

The next alternative for systematic differences in the DM profiles is that the radial profiles have been systematically affected by baryons in different ways. The simplest picture in this context is for AC strength to vary with age. We do not explore this effect in detail, but refer to Fig. 11 for a qualitative understanding. If the IMF has a slightly higher mass normalization than Kroupa (or if the ϵ_{SF} values are somewhat lower than the range adopted so far), then the youngest galaxies are consistent with the B+86 AC recipe, with AC strength decreasing with age until the oldest galaxies have had no AC. Variations in ϵ_{SF} would then play a minor role in the central DM trends. In principle, we could confirm this interpretation by examining the $\langle \rho_{\text{DM}} \rangle$ radial profiles for different age bins, but the data are currently too noisy for this purpose.

Is there any theoretical reason to expect AC to correlate with stellar age in this way? Or more generally, since we do not directly confirm the AC models, why should the DM be preferentially shifted away from the centres of ‘older’ galaxies, and towards the centres of ‘younger’ galaxies, by baryonic effects? One possibility concerns the smoothness of baryonic infall: the standard AC model is motivated by smooth gaseous dissipation, while clumpier collapse – both in the baryons and in the DM – is expected to diminish or even reverse the contraction because of ‘feedback’ from dynamical friction (e.g. Debattista et al. 2008; Romano-Díaz et al. 2008; Jardel & Sellwood 2009; Johansson, Naab & Ostriker 2009).

Alternatively, more direct baryonic feedback could occur via outflows from active galactic nuclei (AGNs) and supernovae, potentially puffing up the DM halo (e.g. Binney, Gerhard & Silk 2001; Mo & Mao 2004; Mashchenko, Wadsley & Couchman 2008; Peirani, Kay & Silk 2008; Duffy et al. 2010; Governato et al. 2010). The notion of feedback effects on DM haloes is becoming commonplace, but the puzzle here is how to explain a systematic trend with time. Realistic simulations of galaxy formation including baryonic processes are only now becoming powerful enough to investigate such questions in detail.

One specific scenario worth special mention is the emerging paradigm of high- z massive galaxy formation by cold streams of

DM and baryons (Agertz, Teyssier & Moore 2009; Dekel et al. 2009a; Dekel, Sari & Ceverino 2009b; Kereš et al. 2009a,b). The clumpiness of the infall in this picture could effectively bypass the contraction process (as proposed by Elmegreen, Bournaud & Elmegreen 2008). These ‘wild disc’ galaxies are thought to initially form hot stellar discs and bulges *in situ* and then evolve into low- z ellipticals, S0s and early-type spirals, whether by passive fading or by eventual merging (Conroy et al. 2008; Genel et al. 2008).

We would like to connect the DM–age trends for ETGs with the impacts of physical processes in the cosmological context of galaxy formation. To begin doing so, let us first consider even broader questions relating the different *types* of massive galaxies. As discussed in Section 3.3, our initial assessment of low- z galaxies is that the early types have haloes with AC and the late types without AC. This type difference is not simply an age difference since among the ETGs, younger galaxies appear to have *stronger* AC. These observations raise an interesting cladistical question: if ETGs are generally formed in the mergers of spirals per the conventional wisdom, do they have roughly the same central DM densities? Also, if the most massive ETGs are assembled by dry mergers of smaller ETGs, are their respective DM densities consistent with this picture?

A related problem was raised long ago for the *stellar* densities of galaxies. It is well known that in the merger of collisionless systems, the phase-space density f cannot be increased, but observationally the central stellar f values of ETGs are much higher than those of spirals (Carlberg 1986; Hernquist, Spergel & Heyl 1993; Mao & Mo 1998). This puzzle was solved by taking into account the dissipational infall of gas during a merger, forming new stars in a high-density central starburst (e.g. Kormendy & Sanders 1992; Mihos & Hernquist 1994; Bekki & Shioya 1998; Springel 2000).

Returning to the DM issues, the 6D density f is of course not the same as the 3D density ρ , or 2D Σ , but we will not carry out a rigorous analysis of f here.¹⁸ Instead, noting that discrepancies involving f tend to be associated with ρ and Σ problems, we will take a very simplified approach to conserving mass: the amount of central DM in a merger remnant should be roughly equal to the sum of the progenitors’ central DM (e.g. Boylan-Kolchin & Ma 2004; Kazantzidis, Zentner & Kravtsov 2006).

The observed DM density differences between early and late types appear from Figs 9 and 10 to be factors of ~ 5 . However, in this context the figures are difficult to interpret because the two galaxy types occupy an almost disjoint region of $R_{\text{eff}}-M_{\star}$ space. Comparison of the small samples of overlapping objects in this space indicates that the real density differences may only be ~ 2 on average.

We illustrate the trends in a different way in Fig. 14, where the mean densities are plotted against stellar masses, in bins of similar physical radius. The simplest expectation here for galaxy mergers is that the DM density within the same radius grows linearly with mass, which the figure shows is borne out remarkably well if considering the spirals as the progenitors of the ETGs. The full story is more complicated, as the true ETG progenitors are thought to have been high- z gas-rich systems with dense DM haloes (e.g. discussions in Section 4.2 and in G+01; T+09). Also, there are potential effects in the mergers that could raise or lower the densities relative to the

¹⁸ We have attempted rough estimates of the coarse-grained phase-space density via $\langle \rho_{\text{DM}} \rangle \sigma_{\text{DM}}^{-3}$ (e.g. Dalcanton & Hogan 2001), but the results are very sensitive to how the DM dispersion σ_{DM} is handled and will require more careful analysis.

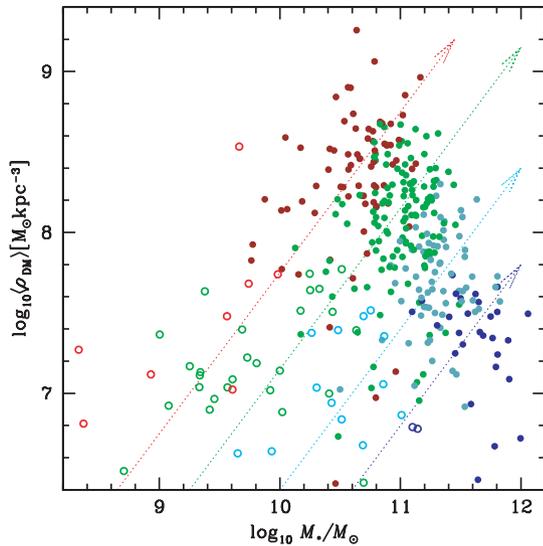


Figure 14. Mean DM density within R_{eff} versus galaxy stellar mass. Filled dark circles are ETGs, while open bright circles are spirals, with colours indicating bins of constant size ($R_{\text{eff}} \sim 2, 4, 9, 15$ kpc, from top to bottom). Arrows indicate linear density growth with mass within the same radius.

linear relation. A detailed theoretical exploration of DM density changes in mergers would be a useful contribution at this point, using also the constraint that the ETG haloes appear cuspier than those of the spirals. However, we can conclude that so far there does not appear to be an obvious problem in generating ETG halo densities from late-type progenitors.

Now returning to the DM trends with *age* in the ETGs, progenitor bias seems an unlikely explanation. Higher DM densities in the early progenitors would produce the *opposite* low- z correlations with what we observe, and furthermore current observations of candidate high- z progenitors indicate relatively *low* central DM content, similar to low- z spirals (Burkert et al. 2010).

Instead, the possibility then remains that halo contraction or expansion is a strong function of z . The scenario would be for halo contraction to become more important for ETGs that formed at later times, perhaps in a transition from the cold stream phenomenon discussed above to a more merger-dominated evolution that produces stronger contraction. Early, strong feedback might also play a role.

An additional effect to consider is dissipationless merging of ETGs: with repeated dry mergers, the initially segregated stellar and DM profiles should eventually converge to a well-mixed NFW-type profile for the *total* mass, implying a net migration of DM towards the galaxy’s centre. One might surmise on this basis that older galaxies would have increased the central DM content because of their longer post-star-formation period of merging. The simulations of Ruszkowski & Springel (2009) highlighted the effect of f_{DM} augmentation through dry merging. However, the main f_{DM} driver here may be the aperture effect due to the merger-induced R_{eff} growth (as the results of Boylan-Kolchin et al. 2005, 2006 suggest), rather than a significant evolution of the DM profile itself.

4.4 Initial mass functions

After exhausting a litany of possibilities for DM properties to vary with age, we consider the alternative that our DM inferences are wrong because we have assumed a universal IMF when estimating the stellar M/L values. The implication would then be that the IMF varies systematically with stellar age, in the sense that younger

galaxies have a higher mass IMF, i.e. they have more stellar mass for a given luminosity, so the high *dynamical* M/L observations are due not to excess DM but to higher M_* . In this context, when we couch the IMF in terms of ‘Salpeter’ and ‘Kroupa’, we will not literally mean the same detailed IMF shape as those functions, but rather an equivalent overall Υ_* normalization.

The subject of IMF variations is highly controversial and unresolved: for a review, see Bastian, Covey & Meyer (2010). There are theoretical reasons to expect lower mass IMFs for stellar populations that formed in the early universe (Larson 2005; Klessen, Spaans & Jappsen 2007) and observational suggestions for IMFs to be lower mass at higher z (e.g. Davé 2008; van Dokkum 2008; Holden et al. 2010). These ideas would fit in well with the IMF–age trend we suggest above. All these results might further be consistent with suggestions that the IMF becomes more top-heavy at higher SFRs (Weidner & Kroupa 2006; Calura & Menci 2009; Haas & Anders 2010), since SFRs should have generally been higher at earlier times. Similarly, the more compact nature of the older ETGs might imply higher gas densities in the star-forming epoch, which has also been suggested to produce more top-heavy IMFs (Lee et al. 2004; Meurer et al. 2009; Krumholz et al. 2010).

To be more quantitative about the implied IMF variations, we construct a simple ad hoc model as follows. First we derive an average $f_{\text{DM}}-M_*$ trend based on the data (Fig. 2), and for each individual galaxy, solve for the Υ_* value that would agree with this trend (given its individual Υ_{dyn} value). This value can be expressed relative to the observed Salpeter-based value by a fraction $f_{\text{IMF}} \equiv \Upsilon_*/\Upsilon_{*,\text{Salp}}$ (cf. T+10). The result of this exercise is shown in Fig. 15: we see a clear trend for the data to require a decreasing IMF mass with age, which we roughly parametrize by a linear function of f_{IMF} versus age.

Given this new IMF assumption, we then re-derive some of our DM inferences. First, we show the revised $f_{\text{DM}}-age$ trends in Fig. 16. As intended by construction, the data show roughly constant DM fractions with age in fixed mass bins (and the galaxies with Salpeter IMFs are not in danger of being unphysical since these were the ones with apparently large f_{DM} to begin with). As we have seen earlier, the $f_{\text{DM}}-age$ trends cannot be immediately interpreted without folding

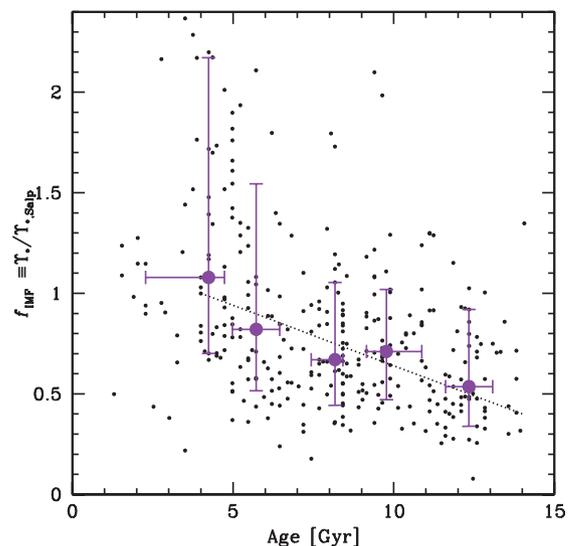


Figure 15. Stellar mass-to-light ratio produced by a variable IMF (relative to Salpeter) versus stellar age. The small points show the results calculated for individual galaxies, the large points with error bars show median values in bins and the dotted line shows the overall trend adopted.

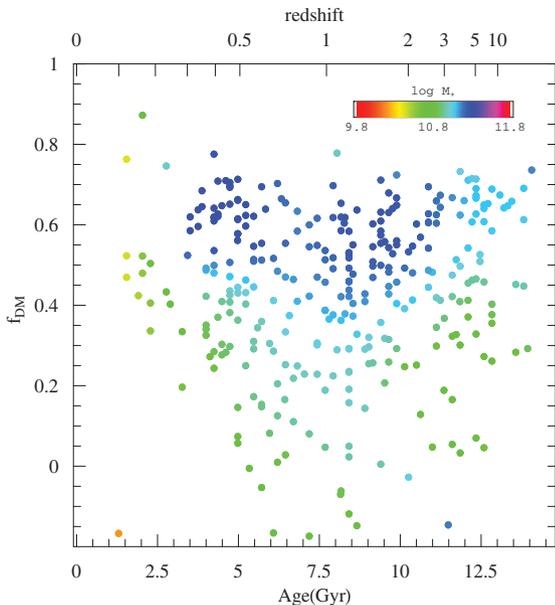


Figure 16. As Fig. 4, with the age-dependent IMF trend taken from Fig. 15.

in the R_{eff} –age dependencies. Upon investigation, we find that with the varying IMF assumption, the strong R_{eff} –age anticorrelations (Fig. 5) have disappeared so that there is on average no correlation.

We find it a remarkable coincidence that the same IMF model removes the age trends in both f_{DM} and R_{eff} . This finding suggests that a systematic IMF variation could be real and that it might account in part for the apparent trends with age and redshift of galaxy sizes (as suggested by Muzzin et al. 2009). Note that with our (somewhat arbitrarily chosen) variable IMF, the f_{DM} observations are now generally consistent with AC halo models.

Given the opposing correlations of age with f_{IMF} and M_* (Fig. 3), we might expect an anticorrelation between f_{IMF} and M_* . We do see a weak suggestion of this effect, which is a bit clearer when considering f_{IMF} versus σ_0 (Fig. 17). This result appears inconsistent with the strong *positive* f_{IMF} – σ_0 correlation found by T+10. However, these authors did not allow for AC in their models, which may impact their conclusions.

As mentioned above, there are reasons to expect an IMF–age trend for ETGs qualitatively similar to our toy model interpretation. (Further analysis would be required to check for problems involving FP twisting with redshift; cf. Renzini 2006.) Now by considering also relatively young, low-SFR, low-density systems as late types, we would expect their IMFs to be more bottom-heavy (high f_{IMF}). Mergers of these systems to form the younger present-day ETGs would then also be expected to result in bottom-heavy IMFs, which is qualitatively consistent with the trend from Fig. 15, but might be in conflict with IMF constraints in low- z galaxy discs.

5 CONCLUSIONS

We have continued an analysis that began in Paper I of a large data set of nearby ETGs, combining dynamics and stellar populations to constrain the central DM content. After having identified variations in DM as the main cause of the tilt of the FP, we have moved on to consider various scaling relations of the DM haloes and connections to the SFHs of the galaxies.

Our basic observational findings are that the central DM fraction f_{DM} within an effective radius R_{eff} has a strong anticorrelation with

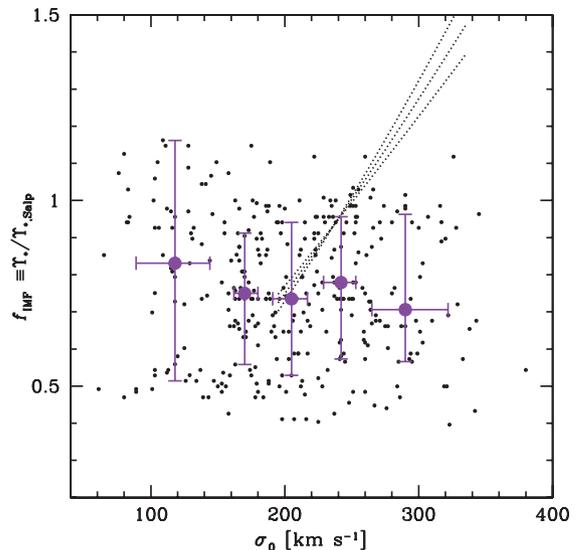


Figure 17. Stellar mass-to-light ratio produced by variable IMF (relative to Salpeter) versus central velocity dispersion. The large points with error bars show median values in bins and the dotted curves show the trend found by T+10, including their range of slopes. The overall normalization of our IMF is somewhat arbitrary, so the key point of comparison is in the slope versus σ_0 .

stellar age and that the galaxy sizes also have an age anticorrelation. We have constructed composite profiles of DM density with radius, finding that they are on average cuspy, with inferred density exponents of ~ -1.6 near R_{eff} . These profiles are steeper than literature findings for spiral galaxies, and the central DM densities of the early types are denser overall, suggesting that gas-rich mergers would need to produce a net halo contraction.

To further interpret the data, we have generated a series of Λ CDM toy models, including variable contributions from AC.

The results from comparisons of models to data are as follows.

- (i) Models with AC fit well overall with a Kroupa IMF, while models without AC prefer a Salpeter IMF.
- (ii) The size–age trends can explain *part* of the f_{DM} –age trends: older galaxies show less evidence for DM because their more compact stellar centres probe less volume of the DM halo.
- (iii) The remaining f_{DM} –age trends are not easily explained by variations in halo mass or concentration and suggest differences in baryonic effects on the DM in the sense that younger galaxies have undergone AC while older galaxies have not.
- (iv) An alternative scenario is for the IMF to be less massive for older stellar populations.

There is ample scope for future insights and improvements. We plan to further investigate the galaxies’ SFHs in the context of theoretical mass assembly histories. Environmental trends can be investigated, as these are expected to be important (e.g. Thomas et al. 2005; Wechsler et al. 2006; Niemi et al. 2010; Rogers et al. 2010; La Barbera et al. 2010b). Forthcoming high-quality, homogeneous, multiwavelength large surveys of low-redshift ETGs should also be able to refute, confirm or extend the trends presented here (Graves 2009; Cappellari et al. 2010). Finally, the gold standard for probing galaxy mass profiles is extended kinematics data along with detailed dynamical modelling – both to provide more leverage on the DM independently of the stellar mass and to sift individual galaxies for the presence of a DM core or cusp (e.g. Thomas et al. 2007; de Lorenzi et al. 2009; Forestell 2009; Weijmans et al. 2009).

Even if some of our current conclusions turn out to be completely wrong, we hope to have introduced a useful framework for interpreting mass results for large data sets of ETGs over cosmic time. The DM constraints are part of a spectrum of clues that can ultimately be combined to pin down the modes of ETG formation (e.g. Almeida, Baugh & Lacey 2007; Covington et al. 2008). Other avenues with considerable promise include central rotation (Jesseit et al. 2009), extra light (Hopkins & Hernquist 2010) and orbital structure (Burkert et al. 2008), as well as halo rotation (Hoffman et al. 2010) and metallicity gradients (Foster et al. 2009; Weijmans et al. 2009), and globular cluster constraints (Rhode et al. 2007; Shin & Kawata 2009; Bekki 2010).

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REFERENCES

- Abadi M. G., Navarro J. F., Fardal M., Babul A., Steinmetz M., 2009, preprint (arXiv:0902.2477)
- Agertz O., Teyssier R., Moore B., 2009, *MNRAS*, 397, L64
- Allanson S. P., Hudson M. J., Smith R. J., Lucey J. R., 2009, *ApJ*, 702, 1275
- Almeida C., Baugh C. M., Lacey C. G., 2007, *MNRAS*, 376, 1711
- Angus G. W., Famaey B., Buote D. A., 2008, *MNRAS*, 387, 1470
- Auger M. W., Treu T., Bolton A. S., Gavazzi R., Koopmans L. V. E., Marshall P. J., Bundy K., Moustakas L. A., 2009, *ApJ*, 705, 1099
- Barnabè M., Auger M. W., Treu T., Koopmans L., Bolton A. S., Czoske O., Gavazzi R., 2010, *MNRAS*, submitted (arXiv:1002.1083)
- Bastian N., Covey K. R., Meyer M. R., 2010, *ARA&A*, in press (arXiv:1001.2965)
- Bekki K., 2010, *MNRAS*, 401, L58
- Bekki K., Shioya Y., 1998, *ApJ*, 497, 108
- Bernardi M., 2009, *MNRAS*, 395, 1491
- Bezanson R., van Dokkum P. G., Tal T., Marchesini D., Kriek M., Franx M., Coppi P., 2009, *ApJ*, 697, 1290
- Binney J., Gerhard O., Silk J., 2001, *MNRAS*, 321, 471
- Blumenthal G. R., Faber S. M., Flores R., Primack J. R., 1986, *ApJ*, 301, 27 (B+86)
- Bolton A. S., Burles S., Treu T., Koopmans L. V. E., Moustakas L. A., 2007, *ApJ*, 665, L105
- Bolton A. S., Treu T., Koopmans L. V. E., Gavazzi R., Moustakas L. A., Burles S., Schlegel D. J., Wayth R., 2008, *ApJ*, 684, 248
- Boyarisky A., Neronov A., Ruchayskiy O., Tkachev I., 2009, preprint (arXiv:0911.3396)
- Boylan-Kolchin M., Ma C.-P., 2004, *MNRAS*, 349, 1117
- Boylan-Kolchin M., Ma C.-P., Quataert E., 2005, *MNRAS*, 362, 184
- Boylan-Kolchin M., Ma C.-P., Quataert E., 2006, *MNRAS*, 369, 1081
- Brown M. J. I. et al., 2008, *ApJ*, 682, 937
- Brownstein J. R., Moffat J. W., 2006, *ApJ*, 636, 721
- Bruzual A. G., Charlot S., 2003, *MNRAS*, 344, 1000 (BC03)
- Bullock J. S., Kolatt T. S., Sigad Y., Somerville R. S., Kravtsov A. V., Klypin A. A., Primack J. R., Dekel A., 2001, *MNRAS*, 321, 559
- Buote D. A., Gastaldello F., Humphrey P. J., Zappacosta L., Bullock J. S., Brighenti F., Mathews W. G., 2007, *ApJ*, 664, 123
- Burkert A., 1995, *ApJ*, 447, L25
- Burkert A., Naab T., Johansson P. H., Jesseit R., 2008, *ApJ*, 685, 897
- Burkert A. et al., 2010, *ApJ*, submitted (arXiv:0907.4777)
- Calura F., Menci N., 2009, *MNRAS*, 400, 1347
- Cappellari M. et al., 2006, *MNRAS*, 366, 1126 (C+06)
- Cappellari M. et al., 2009, *ApJ*, 704, L34
- Cappellari M. et al., 2010, in Koopmans L. V. E., Treu T., eds, *Dark Matter in Early-Type Galaxies*, in press (arXiv:0908.1904)
- Cardone V. F., Tortora C., 2010, *MNRAS*, submitted
- Cardone V. F., Angus G., Diaferio A., Tortora C., Molinaro R., 2010, *MNRAS*, submitted
- Carlberg R. G., 1986, *ApJ*, 310, 593
- Chabrier G., 2001, *ApJ*, 554, 1274
- Chabrier G., 2002, *ApJ*, 567, 304
- Chabrier G., 2003, *PASP*, 115, 763
- Coccatto L. et al., 2009, *MNRAS*, 394, 1249
- Conroy C., Wechsler R. H., 2009, *ApJ*, 696, 620
- Conroy C., Shapley A. E., Tinker J. L., Santos M. R., Lemson G., 2008, *ApJ*, 679, 1192
- Conroy C., Gunn J. E., White M., 2009, *ApJ*, 699, 486
- Cooper A. P. et al., 2010, *MNRAS*, in press (arXiv:0910.3211)
- Covington M. D., 2008, PhD thesis, Univ. California, Santa Cruz
- Covington M., Dekel A., Cox T. J., Jonsson P., Primack J. R., 2008, *MNRAS*, 384, 94
- Daddi E. et al., 2005, *ApJ*, 626, 680
- Dalcanton J. J., Hogan C. J., 2001, *ApJ*, 561, 35
- Davé R., 2008, *MNRAS*, 385, 147
- de Lorenzi F. et al., 2009, *MNRAS*, 395, 76
- Debatista V. P., Moore B., Quinn T., Kazantzidis S., Maas R., Mayer L., Read J., Stadel J., 2008, *ApJ*, 681, 1076
- Dekel A., Cox T. J., 2006, *MNRAS*, 370, 1445
- Dekel A. et al., 2009a, *Nat*, 457, 451
- Dekel A., Sari R., Ceverino D., 2009b, *ApJ*, 703, 785
- Donato F. et al., 2009, *MNRAS*, 397, 1169
- Duffy A. R., Schaye J., Kay S. T., Dalla Vecchia C., 2008, *MNRAS*, 390, L64
- Duffy A. R., Schaye J., Kay S. T., Dalla Vecchia C., Battye R. A., Booth C. M., 2010, *MNRAS*, in press (arXiv:1001.3447)
- Dutton A. A., van den Bosch F. C., Dekel A., Courteau S., 2007, *ApJ*, 654, 27
- Elmegreen B. G., Bournaud F., Elmegreen D. M., 2008, *ApJ*, 688, 67
- Ferreras I., Saha P., Williams L. L. R., 2005, *ApJ*, 623, L5
- Ferreras I., Saha P., Burles S., 2008, *MNRAS*, 383, 857
- Ferreras I., Mavromatos N. E., Sakellariadou M., Yusaf M. F., 2009, *Phys. Rev. D*, 80, 103506
- Forestell A. D., 2009, PhD thesis, Univ. Texas
- Forte J. C., Vega E. I., Faifer F., 2009, *MNRAS*, 397, 1003
- Foster C., Proctor R. N., Forbes D. A., Spolaor M., Hopkins P. F., Brodie J. P., 2009, *MNRAS*, 400, 2135
- Frigerio Martins C., Salucci P., 2007, *MNRAS*, 381, 1103
- Gargiulo A. et al., 2009, *MNRAS*, 397, 75
- Gavazzi R., Treu T., Rhodes J. D., Koopmans L. V. E., Bolton A. S., Burles S., Massey R. J., Moustakas L. A., 2007, *ApJ*, 667, 176
- Genel S. et al., 2008, *ApJ*, 688, 789
- Gentile G., Famaey B., Zhao H., Salucci P., 2009, *Nat*, 461, 627
- Gerhard O., Kronawitter A., Saglia R. P., Bender R., 2001, *AJ*, 121, 1936 (G+01)
- Glazebrook K. et al., 2004, *Nat*, 430, 181
- Gnedin O. Y., Kravtsov A. V., Klypin A. A., Nagai D., 2004, *ApJ*, 616, 16 (G+04)
- Gnedin O. Y., Weinberg D. H., Pizagno J., Prada F., Rix H.-W., 2007, *ApJ*, 671, 1115
- Governato F. et al., 2010, *Nat*, 463, 203
- Graham A. W., Merritt D., Moore B., Diemand J., Terzić B., 2006a, *AJ*, 132, 2701

- Graham A. W., Merritt D., Moore B., Diemand J., Terzić B., 2006b, *AJ*, 132, 2711
- Graves G. J., 2009, PhD thesis, Univ. California
- Graves G. J., Faber S. M., Schiavon R. P., 2009a, *ApJ*, 693, 486
- Graves G. J., Faber S. M., Schiavon R. P., 2009b, *ApJ*, 698, 1590
- Grillo C., Gobat R., 2010, *MNRAS*, 402, L67
- Guo Q., White S., Li C., Boylan-Kolchin M., 2010, *MNRAS*, in press (arXiv:0909.4305)
- Gustafsson M., Fairbairn M., Sommer-Larsen J., 2006, *Phys. Rev. D*, 74, 123522
- Haas M. R., Anders P., 2010, *A&A*, 512, 79
- Hernquist L., Spergel D. N., Heyl J. S., 1993, *ApJ*, 416, 415
- Hoffman L., Cox T. J., Dutta S., Hernquist L., 2010, *ApJ*, submitted (arXiv:1001.0799)
- Holden B. et al., 2010, *ApJ*, submitted
- Hopkins P. F., Hernquist L., 2010, *MNRAS*, 402, 985
- Hopkins P. F., Cox T. J., Hernquist L., 2008, *ApJ*, 689, 17
- Hopkins P. F., Bundy K., Hernquist L., Wuyts S., Cox T. J., 2010, *MNRAS*, 401, 1099
- Humphrey P. J., Buote D. A., 2010, *MNRAS*, 403, 2143
- Humphrey P. J., Buote D. A., Brighenti F., Gebhardt K., Mathews W. G., 2009, *ApJ*, 703, 1257
- Jardel J. R., Sellwood J. A., 2009, *ApJ*, 691, 1300
- Jesseit R., Cappellari M., Naab T., Emsellem E., Burkert A., 2009, *MNRAS*, 397, 1202
- Jiang G., Kochanek C. S., 2007, *ApJ*, 671, 1568
- Johansson P. H., Naab T., Ostriker J. P., 2009, *ApJ*, 697, L38
- Jun H. D., Im M., 2008, *ApJ*, 678, L97
- Kalirai J. S. et al., 2010, *ApJ*, 711, 671
- Kassin S. A., de Jong R. S., Weiner B. J., 2006, *ApJ*, 643, 804
- Kauffmann G. et al., 2003, *MNRAS*, 346, 1055
- Kazantzidis S., Zentner A. R., Kravtsov A. V., 2006, *ApJ*, 641, 647
- Kereš D., Katz N., Fardal M., Davé R., Weinberg D. H., 2009a, *MNRAS*, 395, 160
- Kereš D., Katz N., Davé R., Fardal M., Weinberg D. H., 2009b, *MNRAS*, 396, 2332
- Klessen R. S., Spaans M., Jappsen A.-K., 2007, *MNRAS*, 374, L29
- Koopmans L. V. E. et al., 2009, *ApJ*, 703, L51
- Koposov S. E., Yoo J., Rix H.-W., Weinberg D. H., Macciò A. V., Escudé J. M., 2009, *ApJ*, 696, 2179
- Kormendy J., Freeman K. C., 2004, in Ryder S. D., Pisano D. J., Walker M. A., eds, *Proc. IAU Symp. 220, Dark Matter in Galaxies*. Astron. Soc. Pac., San Francisco, p. 377
- Kormendy J., Sanders D. B., 1992, *ApJ*, 390, L53
- Kroupa P., 2001, *MNRAS*, 322, 231
- Krumholz M. R., Cunningham A. J., Klein R. I., McKee C. F., 2010, *ApJ*, 713, 1120
- La Barbera F., de Carvalho R. R., de la Rosa I. G., Lopes P. A. A., 2010a, *MNRAS*, submitted (arXiv:0912.4558)
- La Barbera F., Lopes P. A. A., de Carvalho R. R., de la Rosa I. G., Berlind A. A., 2010b, *MNRAS*, submitted (arXiv:1003.1119)
- Lackner C. N., Ostriker J. P., 2010, *ApJ*, 712, 88
- Lagattuta D. J. et al., 2010, *ApJ*, submitted (arXiv:0911.2236)
- Larson R. B., 2005, *MNRAS*, 359, 211
- Lee H.-C., Gibson B. K., Flynn C., Kawata D., Beasley M. A., 2004, *MNRAS*, 353, 113
- Li Y.-S., Helmi A., De Lucia G., Stoehr F., 2009, *MNRAS*, 397, L87
- Macciò A. V., Dutton A. A., van den Bosch F. C., 2008, *MNRAS*, 391, 1940
- Macciò A. V., Kang X., Moore B., 2009, *ApJ*, 692, L109
- McGaugh S. S., 2005, *ApJ*, 632, 859
- McGaugh S. S., 2008, *ApJ*, 683, 137
- McGaugh S. S., de Blok W. J. G., Schombert J. M., Kuzio de Naray R., Kim J. H., 2007, *ApJ*, 659, 149
- Mancini C. et al., 2010, *MNRAS*, 401, 933
- Mashchenko S., Wadsley J., Couchman H. M. P., 2008, *Sci*, 319, 174
- Meurer G. R. et al., 2009, *ApJ*, 695, 765
- Mao S., Mo H. J., 1998, *MNRAS*, 296, 847
- Mihos J. C., Hernquist L., 1994, *ApJ*, 437, L47
- Mo H. J., Mao S., 2004, *MNRAS*, 353, 829
- More S., van den Bosch F. C., Cacciato M., Mo H. J., Yang X., Li R., 2009, *MNRAS*, 392, 801
- More S., van den Bosch F. C., Cacciato M., Skibba R., Mo H. J., Yang X., 2010, *MNRAS*, preprint (arXiv:1003.3203)
- Moster B. P., Somerville R. S., Maulbetsch C., van den Bosch F. C., Macciò A. V., Naab T., Oser L., 2010, *ApJ*, 710, 90
- Muzzin A., van Dokkum P., Franx M., Marchesini D., Kriek M., Labbé I., 2009, *ApJ*, 706, L188
- Naab T., Johansson P. H., Ostriker J. P., 2009, *ApJ*, 699, L178
- Napolitano N. R. et al., 2005, *MNRAS*, 357, 691
- Napolitano N. R. et al., 2009, *MNRAS*, 393, 329 (N+09)
- Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
- Niemi S.-M., Heinämäki P., Nurmi P., Saar E., 2010, *MNRAS*, in press (arXiv:1002.0847)
- Nipoti C., Treu T., Bolton A. S., 2009, *ApJ*, 703, 1531
- Okamoto T., Frenk C. S., 2009, *MNRAS*, 399, L174
- Oñorbe J., Domínguez-Tenreiro R., Sáiz A., Serna A., Artal H., 2005, *ApJ*, 632, L57
- Oñorbe J., Domínguez-Tenreiro R., Sáiz A., Artal H., Serna A., 2006, *MNRAS*, 373, 503
- Pedrosa S., Tissera P. B., Scannapieco C., 2009, *MNRAS*, 395, L57
- Pedrosa S., Tissera P. B., Scannapieco C., 2010, *MNRAS*, 402, 776
- Peirani S., Kay S., Silk J., 2008, *A&A*, 479, 123
- Peñarrubia J., Benson A. J., Walker M. G., Gilmore G., McConnachie A., Mayer L., 2010, *MNRAS*, submitted (arXiv:1002.3376)
- Pointecouteau E., Silk J., 2005, *MNRAS*, 364, 654
- Proctor R. N., Forbes D. A., Romanowsky A. J., Brodie J. P., Strader J., Spolaor M., Mendel J. T., Spitler L., 2009, *MNRAS*, 398, 91
- Prugniel Ph., Simien F., 1996, *A&A*, 309, 749
- Renzini A., 2005, in Corbelli E., Palle F., eds, *Astrophysics and Space Science Library Vol. 327, The Initial Mass Function 50 Years Later*. Springer, Dordrecht, p. 221
- Renzini A., 2006, *ARA&A*, 44, 141
- Rhode K. L., Zepf S. E., Kundu A., Larner A. N., 2007, *AJ*, 134, 1403
- Richtler T., Schuberth Y., Hilker M., Dirsch B., Bassino L., Romanowsky A. J., 2008, *A&A*, 478, L23
- Robertson B., Cox T. J., Hernquist L., Franx M., Hopkins P. F., Martini P., Springel V., 2006, *ApJ*, 641, 21
- Rogers B., Ferreras I., Pasquali A., Bernardi M., Lahav O., Kaviraj S., 2010, *MNRAS*, in press (arXiv:1002.0835)
- Romanowsky A. J., 2006, in Mamon G. A., Combes F., Deffayet C., Fort B., eds, *EAS Publications Series Vol. 20, Mass Profiles and Shapes of Cosmological Structures*. EDP Sciences, Les Ulis Cedex, p. 119
- Romanowsky A. J., Strader J., Spitler L. R., Johnson R., Brodie J. P., Forbes D. A., Ponman T., 2009, *AJ*, 137, 4956
- Romano-Díaz E., Shlosman I., Hoffman Y., Heller C., 2008, *ApJ*, 685, L105
- Ruszkowski M., Springel V., 2009, *ApJ*, 696, 1094
- Salpeter E. E., 1955, *ApJ*, 121, 161
- Sánchez-Blázquez P., Forbes D. A., Strader J., Brodie J., Proctor R., 2007, *MNRAS*, 377, 759
- Sanders R. H., 2003, *MNRAS*, 342, 901
- Sanders R. H., Land D. D., 2008, *MNRAS*, 389, 701
- Saracco P., Longhetti M., Andreon S., 2009, *MNRAS*, 392, 718
- Saxton C. J., Ferreras I., 2010, *MNRAS*, in press (arXiv:1002.0845)
- Schmitt H. R., Storch-Bergmann T., Cid Fernandes R., 1999, *MNRAS*, 303, 173
- Schulz A. E., Mandelbaum R., Padmanabhan N., 2010, *MNRAS*, submitted (arXiv:0911.2260) (S+10)
- Sellwood J. A., McGaugh S. S., 2005, *ApJ*, 634, 70
- Sérsic J. L., 1968, *Atlas de Galaxies Australes*. Observatorio Astronómico de Córdoba, Argentina
- Shankar F., Bernardi M., 2009, *MNRAS*, 396, L76 (SB09)
- Shankar F., Marulli F., Bernardi M., Dai X., Hyde J. B., Sheth R. K., 2010, *MNRAS*, 403, 117
- Shen J., Gebhardt K., 2010, *ApJ*, 711, 484

- Shin M.-S., Kawata D., 2009, *ApJ*, 691, 83
 Sommer-Larsen J., Toft S., 2010, *ApJ*, submitted (arXiv:0909.0943)
 Spergel D. N. et al., 2007, *ApJS*, 170, 377
 Spolaor M., Forbes D. A., Proctor R. N., Hau G. K. T., Brough S., 2008, *MNRAS*, 385, 675
 Springel V., 2000, *MNRAS*, 312, 859
 Strigari L. E., Bullock J. S., Kaplinghat M., Simon J. D., Geha M., Willman B., Walker M. G., 2008, *Nat*, 454, 1096
 Stringer M., Cole S., Frenk C., 2010, *MNRAS*, in press (arXiv:0911.1888)
 Terlevich A. I., Forbes D. A., 2002, *MNRAS*, 330, 547
 Thomas D., Maraston C., Bender R., Mendes de Oliveira C., 2005, *ApJ*, 621, 673
 Thomas J., Saglia R. P., Bender R., Thomas D., Gebhardt K., Magorrian J., Corsini E. M., Wegner G., 2007, *MNRAS*, 382, 657
 Thomas J., Saglia R. P., Bender R., Thomas D., Gebhardt K., Magorrian J., Corsini E. M., Wegner G., 2009, *ApJ*, 691, 770 (T+09)
 Tirit O., Combes F., Angus G. W., Famaey B., Zhao H. S., 2007, *A&A*, 476, L1
 Tissera P. B., White S. D. M., Pedrosa S., Scannapieco C., 2010, *MNRAS*, in press (arXiv:0911.2316)
 Tortora C., Napolitano N. R., Romanowsky A. J., Capaccioli M., Covone G., 2009, *MNRAS*, 396, 1132 (Paper I)
 Treu T., Auger M. W., Koopmans L. V. E., Gavazzi R., Marshall P. J., Bolton A. S., 2010, *ApJ*, 709, 1195 (T+10)
 Trujillo I., Burkert A., Bell E. F., 2004, *ApJ*, 600, L39
 Trujillo I. et al., 2006, *ApJ*, 650, 18
 Valentinuzzi T. et al., 2010, *ApJ*, 712, 226
 van der Wel A., Bell E. F., van den Bosch F. C., Gallazzi A., Rix H.-W., 2009, *ApJ*, 698, 1232
 van Dokkum P. G., 2008, *ApJ*, 674, 29
 van Dokkum P. G., Kriek M., Franx M., 2009, *Nat*, 460, 717
 van Dokkum P. G. et al., 2010, *ApJ*, 709, 1018
 Walker M. G., Mateo M., Olszewski E. W., Peñarrubia J., Wyn Evans N., Gilmore G., 2009, *ApJ*, 704, 1274
 Walker M. G., Mateo M., Olszewski E. W., Peñarrubia J., Wyn Evans N., Gilmore G., 2010a, *ApJ*, 710, 886
 Walker M. G., McGaugh S. S., Mateo M., Olszewski E., Kuzio de Naray R., 2010b, *ApJ*, submitted
 Wechsler R. H., Bullock J. S., Primack J. R., Kravtsov A. V., Dekel A., 2002, *ApJ*, 568, 52
 Wechsler R. H., Zentner A. R., Bullock J. S., Kravtsov A. V., Allgood B., 2006, *ApJ*, 652, 71
 Weidner C., Kroupa P., 2006, *MNRAS*, 365, 1333
 Weijmans A.-M. et al., 2009, *MNRAS*, 398, 561
 Wolf J., Martinez G. D., Bullock J. S., Kaplinghat M., Geha M., Munoz R. R., Simon J. D., Avedo F. F., 2010, *MNRAS*, submitted (arXiv:0908.2995)
 Woodley K. A., Gomez M., Harris W. E., Geisler D., Harris G. L. H., 2010, *AJ*, 139, 1871
 Xue X. X. et al., 2008, *ApJ*, 684, 1143
 Zhao H., Xu B.-X., Dobbs C., 2008, *ApJ*, 686, 1019
 Zheng Z., Coil A. L., Zehavi I., 2007, *ApJ*, 667, 760

APPENDIX A: INVESTIGATING NEGATIVE DARK MATTER FRACTIONS

Our derivation of f_{DM} from estimates of Υ_{\star} and Υ_{dyn} yields a number of cases where galaxies have an unphysical $f_{\text{DM}} < 0$. These cases comprise 2, 7 and 25 per cent of the sample for the Chabrier, Kroupa and Salpeter IMFs, respectively. Here we investigate whether these fractions could be compatible with simple observational scatter in Υ_{\star} and Υ_{dyn} .

We have estimated the uncertainties in Υ_{\star} to be ~ 15 per cent using Monte Carlo simulations of spectral energy distribution fitting. For Υ_{dyn} , we may generically consider it as derived via a virial

relation:

$$\Upsilon_{\text{dyn}} = \frac{K \sigma_{\text{eff}}^2 R_{\text{eff}}}{GL}, \quad (\text{A1})$$

where K is a virial coefficient, G is the gravitational constant and L is the luminosity (see equation 19 of C+06). In this context, the value of R_{eff} is arbitrary, and the uncertainty in the physical units R_{eff}/L is dominated by the ~ 10 per cent distance uncertainty. To estimate σ_{eff} , we first consider the central σ_0 with a measurement error of ~ 5 per cent and extrapolate this to the value averaged over R_{eff} – an exercise with an ~ 5 per cent uncertainty (C+06; we consider systematics separately). The value of K was found by C+06 to have an ~ 15 per cent scatter. The net uncertainty in Υ_{dyn} is then ~ 23 per cent and the uncertainty on f_{DM} varies from ~ 0.1 to ~ 0.2 for high to low f_{DM} values.

We begin by assuming that some fraction of the galaxy sample have true $f_{\text{DM}} \sim 0.1$, which we find is the lowest plausible value in a Λ CDM context (Section 3.1). Using the uncertainties discussed above, we generate a random sample of f_{DM} measurements and scale a histogram of their frequency to reproduce the number of $f_{\text{DM}} < 0$ observations while not exceeding any of the $f_{\text{DM}} > 0$ frequencies (Fig. A1). For the Chabrier and Kroupa IMFs, this scaling process implies that ~ 15 per cent of the sample have $f_{\text{DM}} \sim 0.1$. For the Salpeter IMF on the other hand, it is difficult to find a scaling that reproduces the negative- f_{DM} tail without violating the observed distribution to higher values; the best fit has ~ 55 per cent with $f_{\text{DM}} \sim 0.1$.

This difficulty with Salpeter would be eased if a true $f_{\text{DM}} = 0$ is assumed for half the galaxies, if there were systematic errors at the level of ~ 15 per cent, or if the random Υ_{\star} measurement errors were actually at the level of ~ 25 per cent or were non-Gaussian. We would not at present rule out any of these possibilities and so cannot categorically exclude a universal Salpeter IMF. Recalibrating our Υ_{dyn} and Υ_{\star} estimates to external results (as discussed in appendix A of Paper I) would, on the other hand, make the situation more problematic for Salpeter by increasing the number of objects with $f_{\text{DM}} < 0$.

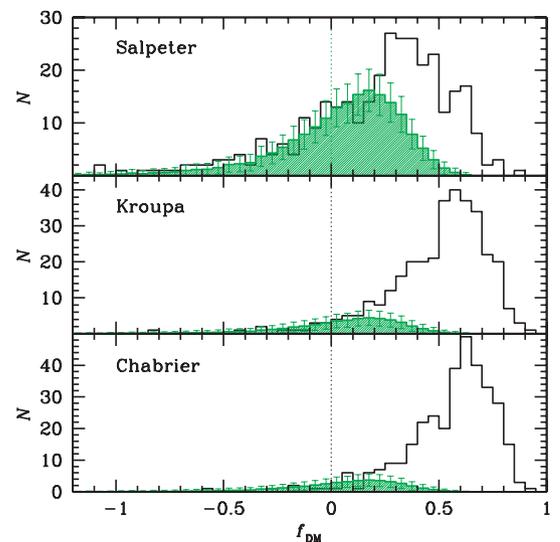


Figure A1. Distributions of DM fraction for ETGs. The open histograms are the observational results, with a different IMF for each panel (as labelled). The shaded green histograms are random realizations of galaxy subsamples having intrinsic $f_{\text{DM}} = 0.1$, with error bars illustrating the approximate Poissonian uncertainties. See the text for further details.

APPENDIX B: CROSS-CHECKS ON AGE DEPENDENCIES

Here we carry out various tests on the robustness of the DM–age trends found in Section 2.2. In Section B1 we compare our f_{DM} results to other literature results, and in Section B2 we examine the implications of results at larger radii. We explore systematic effects in our stellar population models in Section B3.

B1 Central DM content

We check here whether our results on central f_{DM} and age are consistent with other results in the literature. First, we consider the Sloan Digital Sky Survey (SDSS)-based analysis of SB09. They estimated central dynamical and stellar masses for a large sample of ETGs at $z \sim 0.1\text{--}0.2$ using somewhat different techniques to ours in Paper I. After correcting the galaxy luminosities to a common age, they found total M/L trends that depend systematically on age (their fig. 2). Using the information in Bernardi (2009) to convert these corrected luminosities to stellar masses, we convert their M/L values to M_{dyn}/M_* and then to f_{DM} . It also appears that they assumed a Chabrier IMF, so we further convert their results to a Kroupa IMF for comparison with our default models and results.

The results of this exercise are plotted in Fig. B1: there are indications of an anticorrelation between f_{DM} and age, particularly for the higher mass bins. For comparison, our results are also plotted; other than a 30 per cent overall offset in M/L whose origin is unclear, the results are generally consistent. The SB09 data do suggest somewhat shallower correlations that might be explained by a pure R_{eff} –age effect (as stated by these authors), but they do not cover a large enough range in age to be sure.

Graves, Faber & Schiavon (2009b) carried out a different analysis of quiescent ETGs in the SDSS. They again analysed the central dynamical and stellar masses separately and mapped various stellar

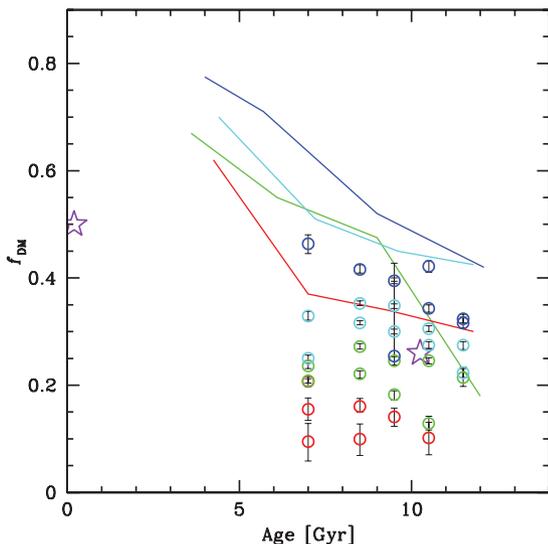


Figure B1. DM fraction with an age of $z \sim 0.1\text{--}0.2$ ETGs, in bins of stellar mass, from SB09; the data have been converted from using a Chabrier IMF to a Kroupa IMF. The colours show the same mass bins as in Fig. 11. For comparison, the lines show our median results in each mass bin. Also shown as large stars are typical results from low- and high- z ETG samples from C+06 and Cappellari et al. (2009). The ‘ages’ plotted in this case actually correspond to the redshifts of the observations, but we do not have enough information to correct these to $z = 0$ ages (which will qualitatively be higher than the plotted points).

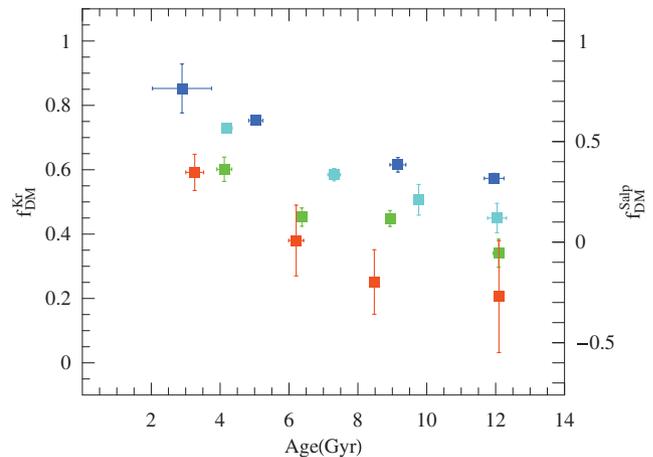


Figure B2. Our results for central DM fraction versus age, in bins of constant σ_0 ($\sim 120, 180, 250, 300 \text{ km s}^{-1}$).

populations parameters on approximate slices of the FP defined by σ_0 and R_{eff} . Although these results do imply a systematic correlation between size and age at fixed mass, if one considers a single σ_0 – R_{eff} grid-point, then one can control for size and mass dependencies and look for any residual correlations in the perpendicular direction of *surface brightness*.

These authors found a negative correlation between age and surface brightness, which because dynamical mass is approximately constant implies that there is a positive correlation between age and overall M/L . They mentioned that in a forthcoming paper they will find that stellar population effects are not enough to explain this age– M/L correlation, which therefore implies a net positive age– f_{DM} correlation. This result is *opposite* to the trend that we find for a residual negative age– f_{DM} correlation after the age– R_{eff} correlations are accounted for. It remains to be seen exactly how and why our results differ.

Fig. 4 of Bastian et al. (2010) provides an intriguing compilation of comparisons between Υ_{dyn} and Υ_* at $z = 0$ and $z \sim 2$ for ETGs from C+06 and Cappellari et al. (2009). While there are many issues in making such a direct comparison between the two galaxy samples the apparent trend does agree qualitatively with our results, although we cannot evaluate the potential residuals from a size-driven trend (Fig. B1).

Since Graves et al. (2009b) among others emphasize the primary importance of σ_0 (rather than M_*) in correlating with stellar population parameters, we show in Fig. B2 our results for f_{DM} –age in bins of constant M_* . The trends are similar to those found in bins of constant σ_0 . However, we do not use σ_0 in general in this paper because it is less straightforward to incorporate in the toy models and because f_{DM} depends on a dynamical mass determination that is intrinsically correlated with σ_0 .

B2 Large-radius DM content

Next we consider DM conclusions from dynamical studies extending to large galactocentric radii, quantified as the M/L gradient parameter introduced in Napolitano et al. (2005). The data set of 25 galaxies is from Paper I (fig. C2, with an update on NGC 4374 from Napolitano et al., in preparation), where we confirmed that the gradient correlated well with f_{DM} from the central regions. We show some of our toy model predictions from this paper in Fig. B3 (top

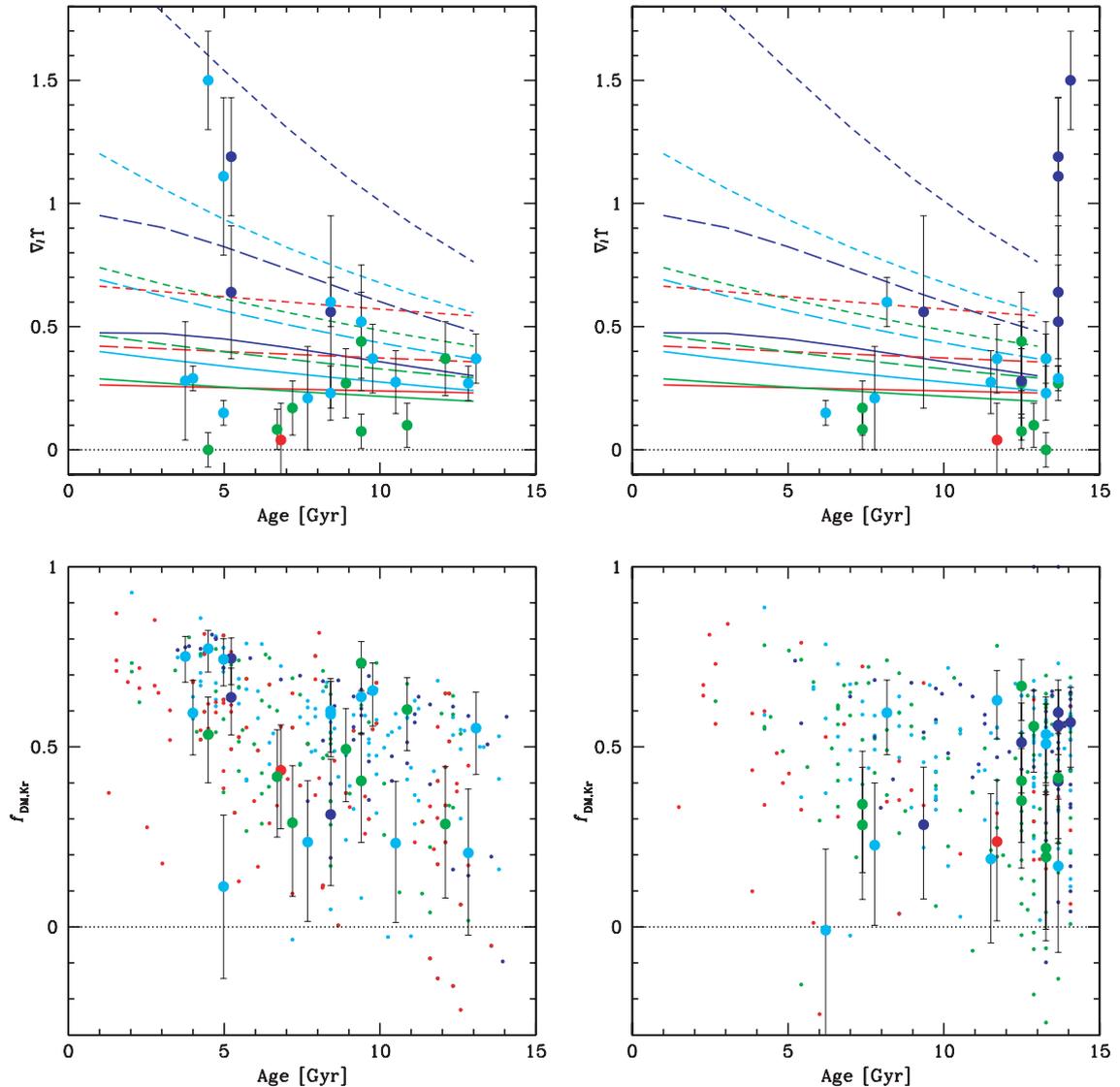


Figure B3. DM–age trends in a subsample of galaxies with large-radius dynamical tracers. The top row shows the ‘mass-to-light ratio radial gradient’, which as an observational parameter is based on dynamics only and is independent of the IMF. Coloured model curves and mass bins are as in Fig. 11. The models are for the case of a Kroupa IMF using a G+04 AC recipe, with line styles showing different ϵ_{SF} values (note that AC has only a small impact on these model predictions since it is a process that acts strongly at small radii). The bottom row shows the DM fraction within R_{eff} , with the small symbols for our full galaxy sample (compare Fig. 4) and large symbols for the subsample with large-radius results. The left-hand panels show results based on our general stellar populations models with metallicity left as a free parameter and the right-hand panels show the results with metallicity fixed to solar ($Z = 0.02$).

left-hand panel). In the absence of any residual DM–age trends, the gradient is expected to decrease with age, since it is defined relative to the galaxy R_{eff} , which in turn decreases with age.

At first glance, the data appear to support this expectation, with the residuals unclear because of the scatter. However, closer examination of the data sets plotted raises a red flag: the four ‘young’, DM-dominated galaxies are all well-known nearby massive group- and cluster-central ellipticals that are normally thought to have very old stellar populations (averaged over R_{eff}), *not* with ages of ~ 4 –5 Gyr (NGC 1407, M49, M87, NGC 5846; e.g. C+06; Sánchez-Blázquez et al. 2007; Spolaor et al. 2008).

It turns out that these galaxies were all fitted with an unrealistic super-solar metallicity ($Z = 0.05$) which is probably a reflection of the fundamental age–metallicity degeneracy in stellar population analyses. We re-ran our analysis with all the metallicities fixed to $Z = 0.02$, which for the four problematic galaxies yielded more

credible ages of 14 Gyr.¹⁹ The revised M/L -gradient results are shown in the top right-hand panel of Fig. B3, where it now appears that the DM content may *increase* with age. The first thing to keep in mind when considering this apparent inconsistency with our main f_{DM} –age result is that the DM content within R_{eff} and within $\sim 5R_{\text{eff}}$ may very well *not* be tightly correlated. The central DM content may be less a reflection of the overall DM content and more closely related to the details of the baryonic–DM interplay at the centres of haloes.

The main effect, however, appears to be small number statistics coupled with selection effects. In the bottom panels of Fig. B3, we

¹⁹ For galaxies with a previous best fit of $Z = 0.02$, the age and Υ_* results with fixed Z are *not* exactly the same as before. This is because our procedure does not involve a simple best fit, but rather a Monte Carlo approach using the median of a distribution of best-fitting values (see Paper I).

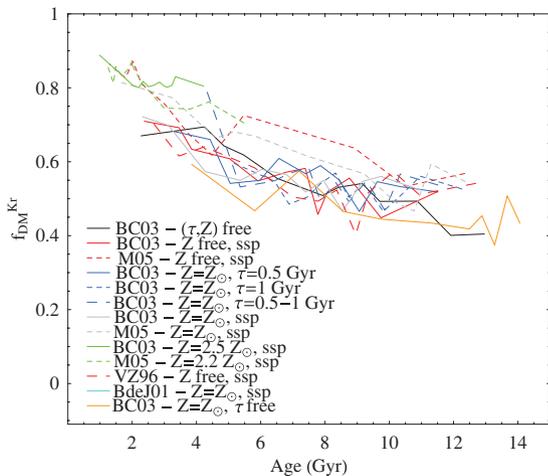


Figure B4. Trend of DM fraction with age, for different stellar population assumptions. The solid black line is our standard model; see fig. A2 of Paper I for explanations of the other line styles. An additional model not included here is from Conroy, Gunn & White (2009), but for the wavelengths and ages used here, this would be equivalent to BC03 (see their fig. 10). We have also tested delayed-exponential and truncated (step-function) SFH models, which are not shown here but again do not substantially change the trends above.

show the f_{DM} results for both metallicity assumptions. The stellar M/L is not very sensitive to the age–metallicity degeneracy and in fact is affected in a way that roughly parallels the overall f_{DM} –age trend: higher age and lower Z yield higher M/L_* and lower f_{DM} . Thus, the general trend for our full sample is not qualitatively affected by changing Z , but quantitatively shifts to larger ages overall (see also Fig. B4).

The subsample of galaxies that have large-radius data is not so fortunate. When fixing $Z = 0.02$, their f_{DM} –age results buck the overall trend, just as seen from the large-radius results. This sample has only one ‘young’ object ($\lesssim 7$ Gyr), and the old objects are dominated by systems such as M87 which are known to reside at the centres of massive groups and clusters and appear to comprise the high- f_{DM} tail of the overall distribution. Such galaxies are some of the first targets of large-radius dynamical studies because of their rich supply of mass tracers such as globular clusters and planetary nebulae. The price to pay from selecting such systems is that they could provide a very biased view of the Universe, as now appears to be the case when considering the central DM content.

This selection effect may explain the curious DM halo mass–concentration trends found by N+09 and demonstrates the importance of constructing an unbiased galaxy sample. For now, the DM–age implications from large-radius tracers are totally inconclusive and will require completion of a large systematic survey to make any progress. In the meantime, valuable spot-checks could be provided by studying the large-radius DM content of a few galaxies that help drive the apparent f_{DM} –age relation, e.g. young DM-rich systems such as NGC 3626 or old DM-poor ones such as NGC 7454.

B3 Modelling systematics

We next consider whether systematic uncertainties in our modelling of stellar populations could be affecting our DM inferences. We have

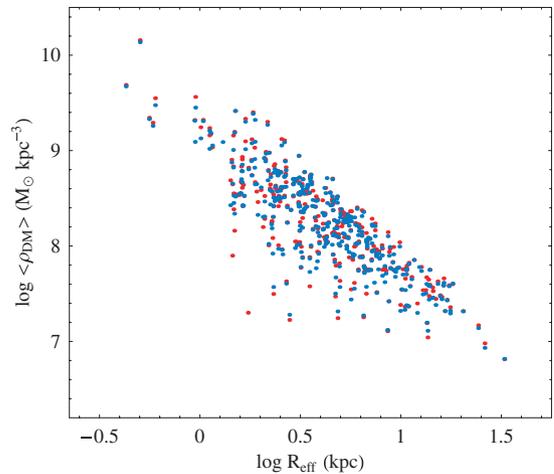


Figure B5. Central DM density versus effective radius for our full ETG galaxy sample. Red points show results using our standard assumptions in the stellar populations models and blue shows the case where metallicity is fixed to solar.

already explored this issue for the f_{DM} –mass relations in Paper I; here we consider f_{DM} –age. Age and stellar M/L are both derived from the same models, and a positive error in age would correlate with a positive error in stellar M/L which would produce a *negative* error in f_{DM} and thus mimic an anticorrelated f_{DM} –age trend solely because of the correlated errors.

To see if such a systematic could account for our observational result, we experiment with different stellar population assumptions, showing the results in Fig. B4. The only cases where the f_{DM} –age trend is appreciably different from our standard estimate is when we unrealistically fix $Z = 0.05$ (green lines).

Given the problem identified in Section B2 where allowing the metallicity to vary freely can result in too many ‘young’ galaxies, we consider in particular our results when we fix $Z = Z_{\odot}$. The implications for f_{DM} were shown in Figs B3 (right-hand panels) and B4 (solid orange line), and we also show the overall $\langle \rho_{\text{DM}} \rangle$ – R_{eff} trend in Fig. B5. The results are similar to our standard model.

As a reminder, in Paper I (appendix A) we compared independent estimates of Υ_{dyn} and Υ_* to C+06 for galaxies in common. Our Υ_* values for the same IMF were ~ 20 per cent lower and our Υ_{dyn} values were higher (by ~ 10 per cent for the brighter galaxies and ~ 30 per cent for the fainter galaxies). Thus our f_{DM} values are systematically higher, becoming more discrepant for fainter galaxies (up to $\Delta f_{\text{DM}} \sim 0.25$). The C+06 results on their own may imply consistency with no-AC models for the fainter galaxies and strong AC for the brighter ones (Romanowsky 2006), which would be in better agreement with the results of N+09. We have also checked the f_{DM} –age trends using a recalibration as discussed in Paper I, and found that for the fainter galaxies, the anticorrelation becomes somewhat *steeper*, so our basic f_{DM} –age result does not go away.

An issue not discussed in Paper I is the potential effect of AGN emission on the observed galaxy colours and thus on the Υ_* and age inferences. The AGN colours would generally mimic a stellar population of an age of ~ 3 Gyr and skew the inferences towards younger ages and higher f_{DM} . Matching the SDSS DR4 ETG catalogue with the AGN catalogue of Kauffmann et al. (2003), we find

that ~ 15 per cent of the galaxies in our sample's mass range have strong AGNs (Type II with strong [O III] emission; see also Graves, Faber & Schiavon 2009a). However, based on the AGN study of Schmitt, Storchi-Bergmann & Cid Fernandes (1999), we estimate that the contaminant light would account for only ~ 1 – 2 per cent

of the total light within R_{eff} and would bias our results by at most 5 per cent.

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