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VIMOS mosaic integral-field spectroscopy of the bulge and disc of the early-type galaxy NGC 4697

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ABSTRACT

We present an integral-field study of the internal structure, kinematics and stellar population of the almost edge-on, intermediate-luminosity (L^*) elliptical galaxy NGC 4697. We build extended two-dimensional (2D) maps of the stellar kinematics and line strengths of the galaxy up to ~ 0.7 effective radii (R_{eff}) using a mosaic of eight VIMOS (VISible Multi-Objects Spectrograph, on the Very Large Telescope) integral-field unit pointings. We find clear evidence for a rotation-supported structure along the major axis from the 2D kinematical maps, confirming the previous classification of this system as a ‘fast rotator’. We study the correlations between the third and fourth Gauss–Hermite moments of the line-of-sight velocity distribution (LOSVD) h_3 and h_4 with the rotation parameter (V/σ), and compare our findings to hydrodynamical simulations. We find remarkable similarities to predictions from gas-rich mergers. Based on photometry, we perform a bulge/disc decomposition and study the stellar population properties of the two components. The bulge and the disc show different stellar populations, with the stars in the bulge being older ($\text{age}_{\text{bulge}} = 13.5^{+1.4}_{-1.4}$ Gyr, $\text{age}_{\text{disc}} = 10.5^{+1.6}_{-2.0}$ Gyr) and more metal poor ($[M/H]_{\text{bulge}} = -0.17^{+0.12}_{-0.1}$, $[M/H]_{\text{disc}} = -0.03^{+0.02}_{-0.1}$). The evidence of a later-formed, more metal-rich disc embedded in an older, more metal poor bulge, together with the LOSVD structure, supports a mass assembly scenario dominated by gas-rich minor mergers and possibly with a late gas-rich major merger that left a previously rapidly rotating system unchanged. The bulge and the disc do not show signs of different stellar initial mass function (IMF) slopes, and both match well with a Milky Way-like IMF.

Key words: galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: formation.

1 INTRODUCTION

Integral-field spectroscopy (IFS) makes it possible to derive spatially resolved and continuously sampled line-of-sight velocity distribution (LOSVD) two-dimensional (2D) maps, providing a unique view of the internal structure of galaxies, as well as detailed morphological and dynamical information, which are essential to fully understand the physical mechanisms that regulate galaxy formation and evolution.

An increasing number of studies have recently derived the integrated and spatially resolved properties of a statistical sample of nearby galaxies using IFS (e.g. Bacon et al. 2001; Thatte et al. 2006;

Czoske et al. 2008; Rawle et al. 2008; Scott et al. 2012; Fogarty et al. 2014; González Delgado et al. 2014; Sánchez-Blázquez et al. 2014; Walcher et al. 2014). Among the most successful, the ATLAS^{3D} (Cappellari et al. 2011) data set provides a complete inventory of the baryon budget and a detailed 2D analysis of stellar and gaseous kinematics, together with resolved stellar populations of a complete and statistically significant sample of early-type galaxies (ETGs) in the local Universe.

One of the most interesting results of the ATLAS^{3D} team is that it is possible to connect the 2D features of the LOSVD to galaxy formation scenarios by comparing these to high-resolution (HR) simulations (e.g. Naab et al. 2014, hereafter N+14). For the majority of ETGs (~ 86 per cent), the formation processes result in disc-like objects that maintain the axisymmetric shape of the progenitors, while only a small fraction of systems (the most massive

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ones, 12 per cent) rotate slowly with no indication of discs. Possible processes that can form embedded disc components include minor mergers, major merger with disc re-growth, gas accretion events and secular evolution (Krajinović et al. 2011).

It is possible, in principle, to discriminate between the different formation scenarios by accurately analysing the higher moments of the LOSVD, the Gauss–Hermite coefficients h_3 and h_4 (representing skewness and kurtosis, respectively) and comparing these to HR simulations. Gas has a strong impact on the non-Gaussian shapes of the LOSVDs of the merger remnants. In particular, when the values of h_3 and V_{rot} have opposite signs, the leading (pro-grade) wing is steeper than the trailing (retrograde) wing. This indicates the presence of an embedded disc (see e.g. Bender, Saglia & Gerhard 1994). When an anti-correlation between the h_3 and the V/σ values does not exist but the galaxy shows a fast rotation profile, the rotation could be caused by a recent gas-poor major merger that typically does not permit re-growth of a significant disc component (Barnes & Hernquist 1996; Naab, Khochfar & Burkert 2006b; Hoffman et al. 2009, hereafter H+09). The sign of the h_4 coefficient distinguishes a peaked shape ($h_4 > 0$) from a flat-topped one ($h_4 < 0$), where the peak is broad and the wings are narrow.

The age and metal content of galaxies can provide additional important information about the formation mechanism and its time-scale. To obtain an unbiased and detailed understanding of the star formation history of ETGs from their integrated light, it is necessary to study carefully their unresolved stellar populations, and to break the degeneracies existing between age, metallicity and element abundances (e.g. Worthey 1994). The best way to achieve this is to look at as many stellar features as possible (Spiniello et al. 2014, hereafter S+14). Clearly, in this context, the very narrow wavelength coverage of (SAURON) (Spectroscopic Areal Unit for Research on Optical Nebulae) is not ideal, since stellar population studies would rely on only a few well-known blue–optical absorption lines ($H\beta$, Mgb , $Fe5015$ and $Fe5270$).

The Visible Multi-Object Spectrograph (VIMOS) integral-field unit (IFU) on the VLT covers a wider wavelength range (5250–7400 Å using the blue grism and 6450–8600 Å using the red grism, in its HR mode), with comparable spectral and spatial resolution, and is clearly a better choice. This point is particularly relevant after the recent finding that the low-mass end of the initial mass function (IMF) is not universal and might become steeper with increasing stellar mass and stellar velocity dispersion (Treu et al. 2010; van Dokkum & Conroy 2010; Cappellari et al. 2012, 2013; Spiniello et al. 2012; Dutton et al. 2013; La Barbera et al. 2013; Tortora, Romanowsky & Napolitano 2013; S+14; Tortora et al. 2014a,b, but see also Smith, Lucey & Conroy 2015) or with changes in other galaxy structural parameters (e.g. compactness, stellar density, size; see Spiniello, Trager & Koopmans 2015b). Indeed, to constrain the low-mass end of the IMF, it is crucial to look at the red part of the optical spectra, where most of the gravity-sensitive absorption features are present. Some of the most promising indices, like the optical TiO molecular lines (bTiO, aTiO, TiO1; S+14), can be captured with VIMOS but are outside the SAURON coverage. Moreover, the SWELLS survey (Treu et al. 2011) has shown that there could be one universal normalization of the IMF for disc-like stellar populations (presumably Chabrier-like) and a heavier (Salpeter-like) one for the older and possibly more metal rich stellar populations found in massive bulges and spheroids (Dutton et al. 2013).

Furthermore, 2D IFS is crucial for constraining the low-mass end of the IMF slope, given the very recent observational evidence that seems to indicate that the IMF normalization varies within in-

dividual galaxies (Martín-Navarro et al. 2015; Sarzi, Spiniello & Krajinović, in preparation) following the velocity dispersion profile (but see also Smith et al. 2015; Spiniello et al. 2015a). In this context, a strong observational limitation is the small field of view (FOV) of IFS instruments, typically of the order of tens of arcsec, preventing good coverage of the target galaxies on the sky and a proper spatially resolved study of the IMF slope. Only very recently has a new class of integral-field spectrographs with larger FOV been built to achieve the challenging task of studying the internal structure of galaxies. Unfortunately, most of them still cover a small range of wavelength and therefore are not suitable for constraining the low-mass end of the IMF normalization (e.g. VIRUS-P, 3600–6800 Å, Hill et al. 2008, or CALIFA, 3750–7000 Å, Sánchez et al. 2012 in its red configuration). The VLT Multi Unit Spectroscopic Explorer (MUSE) represents the perfect solution, combining a large FOV (1×1 arcmin) and a wide wavelength range (4650–9300 Å). Finally, in the near future the MaNGA (Mapping Nearby Galaxies at APO; Bundy et al. 2015) project will further improve the situation, covering the wavelength range 3600–10 000 Å up to $\sim 1.5R_{\text{eff}}$ with a spatial resolution of ~ 2.5 arcsec (and a resolution of $R \sim 2000$). We refer to table 1 of Brodie et al. (2014) for a detailed comparison between current and future integral-field surveys.

Thus, at the present day, VIMOS remains among the most competitive instruments for galaxy structure and stellar population studies, provided that the small FOV is compensated for by mosaic observations, i.e. contiguous IFU pointings astrometrized and stacked together to create a larger final FOV.

This is indeed the approach we decided to follow in order to attempt to systematically study the stellar population and the kinematics of the whole surface of nearby ETGs up to almost the effective radius over a large wavelength range and with medium to high spectral resolution. Given the challenges in the data reduction, calibration and visualization, in this paper we present the analysis of one of the galaxies observed in a pilot programme focusing on the kinematics and the stellar population analysis of two flattened galaxies (NGC 821 and NGC 4697).¹ In particular, we focus on NGC 4697, for which we have eight VIMOS pointings covering ~ 80 arcsec \times 60 arcsec effective area, which allows us to investigate regions up to ~ 45 arcsec from the galaxy centre, i.e. $\sim 0.7R_{\text{eff}}$ ($R_{\text{eff}} = 66$ arcsec from de Lorenzi et al. 2008).

1.1 Target selection and outline of the paper

NGC 4697 is an E6, almost edge-on galaxy located along the Virgo southern extension, with evidence of substructures [from *Hubble Space Telescope* (HST)] ACS images; Lauer et al. 1995 and the presence of a strong streaming motion component (known since the pioneering work of Bertola & Capaccioli 1975). The galaxy is representative of the class of discy, fast-rotating elliptical L^* galaxies (Bender et al. 1994), which are candidates for having less dark matter (DM; Covone & Napolitano 2004; Napolitano et al. 2005; Cappellari et al. 2006) than the boxy, slow rotators of comparable mass. As such, it is among the sample of elliptical galaxies whose planetary nebula (PN) kinematics show surprisingly little sign of a massive DM halo (Méndez et al. 2001, 2009; Romanowsky et al. 2003; Napolitano et al. 2005; Douglas et al. 2007). Finally, the

¹ Programme IDs P079.B-0622A and P085.B-0949 (PI: N.R. Napolitano).

spheroidal part of NGC 4697 may well be a large bulge (Dejonghe et al. 1996) and therefore this system is an interesting candidate for testing the SWELLS (Treu et al. 2011) claim about the different normalization of the low-mass end of the IMF slope (universal, Milky Way-like IMF for the discs and more bottom-heavy IMF, possibly steepening with velocity dispersion for bulges; Dutton et al. 2013).

Based on the combined information of the 2D kinematical characterization of the system, and stellar population analysis, we are able to fully characterize the stellar population of the disc and bulge components of NGC 4697. This is important *per se* in order to investigate the variation of the stellar populations and IMF slope in different structural components of the system and possibly to infer a formation scenario for it, but it will also allow us to correctly normalize the stellar mass of the galaxy in forthcoming dynamical analysis.

The paper is organized as follows: in Section 2 we outline the observations and data reduction, in Section 3 we introduce our kinematical technique and we show 2D kinematical maps as well as one-dimensional (1D) kinematics profiles along the major axis of the galaxy. We compare the latter with kinematics profiles in the literature. In particular, in this section, we highlight correlations between the h_3 and h_4 parameters and V/σ and try to connect these to predictions from galaxy formation simulations. In Section 4, we focus on the stellar population analysis performed via line-index strength measurements on the 1D spectra of the disc and of the bulge separately. We also attempt to constrain the IMF normalization of the two components. Finally, discussion and general findings are presented in Section 5 together with conclusions.

2 OBSERVATION AND DATA REDUCTION

The IFS observations have been carried out with VIMOS, the multi-mode wide-field optical instrument mounted on the European Southern Observatory (ESO) VLT, as part of a proposal carried on in two different runs: May 2007 and March 2010, respectively.

The IFU head of VIMOS consists of a square array of 80×80 micro-lenses with a sampling of 0.67 arcsec per spatial element (spaxel) at the Nasmyth focus of the telescope. In its HR mode only the central 1600 lenses are used, resulting in an FOV of 27 arcsec \times 27 arcsec. The seeing was generally below 0.8 arcsec during the observations, so that each fibre of the IFU gives a spectrum that is essentially independent of its neighbours. All data sets have been taken with the HR-blue grism, covering a wavelength range between 4100 and 6300 Å with a dispersion of 0.51 Å pixel⁻¹. The spectral resolution of the HR-blue grism is $\lambda/\Delta\lambda = 2550$, where $\Delta\lambda$ is the full width at half-maximum (FWHM) of the instrument profile.

In order to cover the effective radius and to have a high signal-to-noise ratio (S/N) among the whole FOV, we use a mosaic approach. During the first observation run, we have obtained six different pointings, each made by six exposures of ~ 800 s, while for the second run we have obtained two pointings along the major axis of the system, with four exposures of 825 s each. The position angle of all pointings has been fixed to PA = -20° . For accurate measurements of the stellar kinematics, one must properly evaluate the sky background contamination. Since VIMOS does not have sky-dedicated fibres, and the target galaxy covers the entire FOV, it has been necessary to make separate telescope pointings (of ~ 300 s) to measure the sky background for each observing night.

2.1 Data reduction

Data reduction has been carried out using the VIMOS pipeline (version 2.1.1) by the ESO ‘Data Flow System Group’,² via the Gasgano data file organizer developed by ESO. A schematic representation of the data reduction processes performed within Gasgano is provided by Zanichelli et al. (2005). The reduction involves standard steps for fibre-based IFS. (i) Subtraction of a master bias frame created by averaging all the bias frames; (ii) fibre identification (correct association of a fibre position on the IFU mask to a corresponding position on the CCD) and correction of the differences in the fibre-to-fibre transmission; (iii) Wavelength calibration; and (iv) scientific image combination.

We create a final data cube for each pointing by averaging over the final single-observing block flux-calibrated science cubes produced by the pipeline using the IFU tables, with the correspondence between each row (i.e. each spectrum) of the extracted spectra and the fibre’s position on the IFU head. The four IFU quadrants are treated independently up to the combination of all exposures into a single-pointing data cube.

We subtract the median sky spectrum (extracted directly from the ‘sky frame’) from the ‘science frame’. Median combining was needed to ensure that any residual contamination from faint objects was removed. We are confident that this sky-subtraction procedure is accurate enough but, if prominent sky emission lines are still present in the spectra after sky subtraction, we mask them out during the stellar kinematics recovery process. We also investigate possible effects of the differential atmospheric refraction by checking whether or not the photometric centre of the galaxy changes its position as a function of wavelength in the data cube. We find that the position does not change by more than 2 pixels (~ 1.3 arcsec) between 4000 and 6200 Å. Each single-pointing data cube is astrometrically calibrated independently by matching the position of the photometric centre or of a star (when the centre is not visible in that particular pointing) to an *HST* ACS image of the central region of NGC 4697. Subsequently, the different pointings are linearly interpolated on to a common 3D grid uniformly sampled, after normalizing them for exposure times, using a simple grid interpolation³ to obtain the final mosaic data cube. As a further check of the robustness of this mosaic reconstruction, we re-check the WCS (world coordinate system) astrometry on our final combined image, using as reference the same *HST*/ACS image.

The results of this comparison are visible in Fig. 1 which shows the final reconstructed mosaic image, obtained by averaging the data cube along the wavelength direction, as well as the ACS image. 20 blue isophotes, derived from the *HST* image, are shown on both images to check the accuracy of our astrometrical calibration. Although most of the point sources visible in the *HST* image are too faint to be detected in our VIMOS observations, the positions of the centre and of a star (East-North-East, ENE) coincide between the two instruments. On the right side of Fig. 1, we also plot a 1D spectrum extracted at the photometric centre of the galaxy and another 1D spectrum extracted at ~ 40 arcsec, in a position which is highlighted in the reconstructed FOV by a small yellow star.

² N. Devillard, Y. Jung, R. Palsa, C. Izzo, P. Ballester, C. Sabet, K. Banse, M. Kiesgen, L. Lundin, A. Modigliani, D. J. McKay, N. P. F. McKay, J. Moeller-Larsen and L. de Bilbao.

³ We interpolate, shift and rotate each single-pointing data cube by matching the position of the centre and the rotational angle into this data cube container.

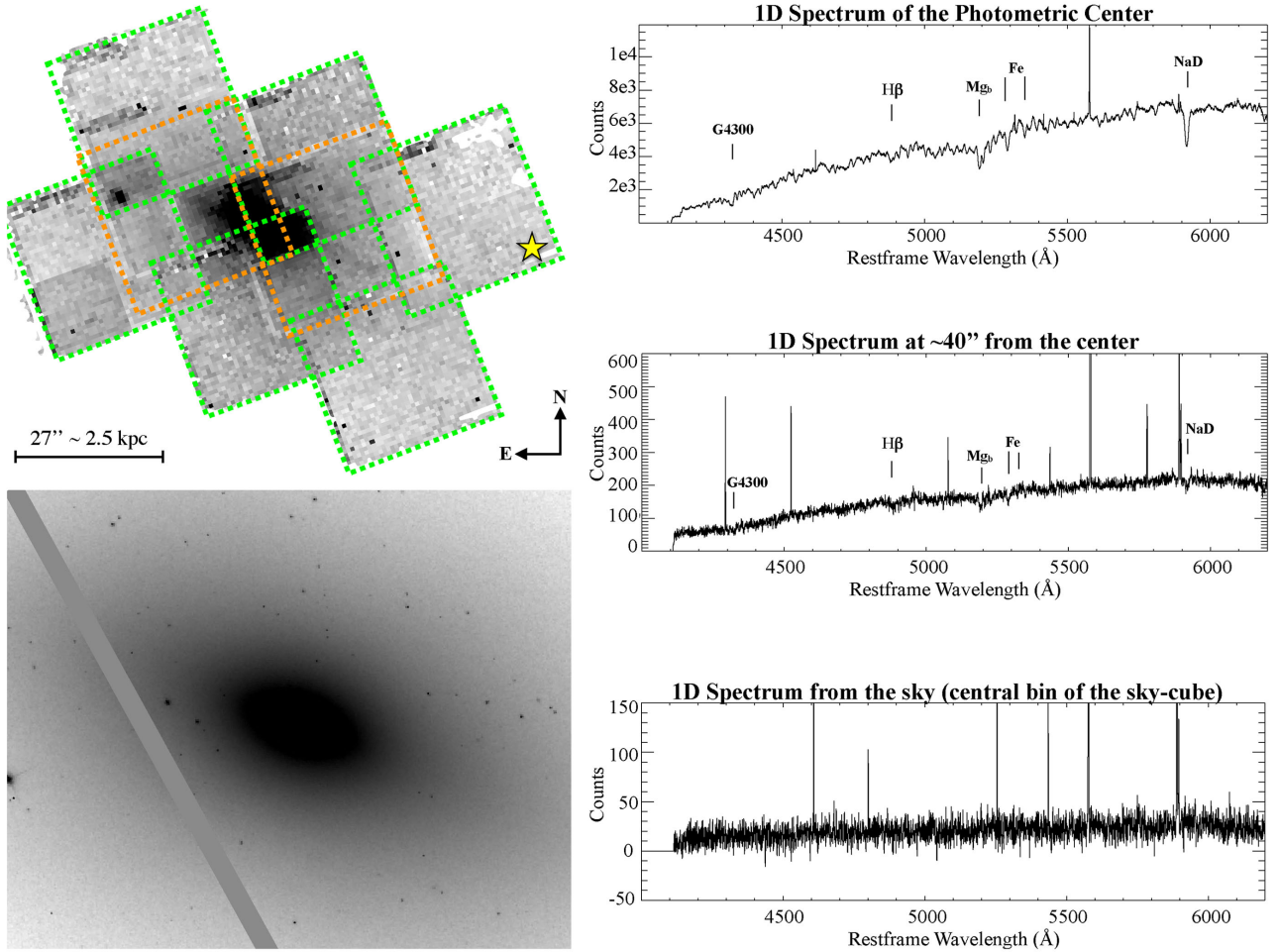


Figure 1. Upper-left panel: NGC 4697 final reconstructed FOV generated with the mosaic of all the on-target pointings, summing the fluxes in the data cube over the whole wavelength range. Each dotted square represents a single pointing of $27'' \times 27''$. Green cubes are data from the first run (P79), whereas orange cubes are the two central pointings from the second one (P85). Bottom-left panel: *HST* ACS image of the central region of NGC 4697 shown at the same spatial scale as the reconstructed VIMOS image. Blue isophotes (derived from the *HST* image) are shown on both images to check the validity and precision of the astrometrical calibration (see the text for more details). Right-hand panels: 1D galaxy spectra extracted in the central spatial bin (top panel) and in a spatial bin at $\sim 40''$ from the galaxy centre (middle panel), and 1D sky spectrum (bottom panel) extracted from the central bin of the sky-cube. The spaxel from which the middle-panel 1D spectrum has been extracted is highlighted on the reconstructed FOV by a yellow star. The spectra are plotted versus rest-frame wavelength and several spectral features of the galaxy such as G4300, H β , Mgb and Fe lines are marked. Sky emission lines have not been removed from the spectra but have been properly masked out during the fitting procedures.

3 KINEMATICS MEASUREMENTS

To recover the LOSVD first, we apply a spatial binning method in order to increase the S/N over the single-fibre spatially resolved spectra of the final data cube (Section 3.1). Then we measure the LOSVD for the galaxy using a maximum penalized likelihood approach (Section 3.2). In this way, we obtain spatially resolved 2D profiles of rotation velocity, V , and velocity dispersion, σ , as well as the higher order Gauss–Hermite moments h_3 and h_4 , which quantify the asymmetric and symmetric departures of the LOSVD from a pure Gaussian (related to the skewness and kurtosis, respectively). In the following sections, we describe in more detail the spatial binning procedure and the LOSVD recovery.

3.1 Voronoi 2D spatial binning

We rebin the data with a fixed final S/N by locally averaging neighbouring fibres (using the Voronoi adaptive technique `IDL` code of Cappellari & Copin 2003). We use an adaptive binning scheme

where the size of the bin is adapted to the local S/N. Bigger spatial bins are created in the low-S/N regions, with higher resolution and smaller bins retained in the high-S/N regions (towards the galaxy centre).⁴

We tested several values for the final Voronoi fixed S/N, choosing a final threshold of $S/N = 70$. In this way, we are able to obtain a HR spectrum for each spaxel, which is fundamental for performing a proper stellar population analysis based on line-index measurements, while also preserving spatial information that is crucial for recovering detailed spatially resolved kinematics information up to ~ 0.7 effective radius ($R_{\text{eff}} = 66''$).

3.2 LOSVD fitting

We measure rotation velocity, velocity dispersion and higher order moments (V , σ , h_3 , h_4) for each Voronoi-rebinned spaxel spectrum

⁴ Spaxels are arranged in a regular grid, so the size is a well-defined quantity.

using the `PPXF` code of Cappellari & Emsellem (2004). This software, working in pixel space, finds the combination of stellar templates which, convolved with an appropriate LOSVD, best reproduces the galaxy spectrum. The best-fitting parameters of the LOSVD are determined by minimizing χ^2 , which measures the agreement between the model and the observed galaxy spectrum, over the set of N good pixels. We use the MILES library of stellar spectra (Sánchez-Blázquez et al. 2006; Falcón-Barroso et al. 2011): 985 well-calibrated stars covering the region from 3525 to 7500 Å at a spectral resolution of 2.54 Å (FWHM), $\sigma \sim 57 \text{ km s}^{-1}$ at the central wavelength, obtained at the 2.5 m Isaac Newton Telescope. The `PPXF` code also allows the user to mask out noisy regions of the galaxy spectrum. We mask out sky emission lines, focusing on absorption lines from 4300 to 5400 Å (including in this way $H\beta$, O lines, Mgb doublet and Fe). Uncertainties on individual points are obtained via a Monte Carlo simulation.

We follow the approach described in Cappellari & Emsellem (2004), to recover rotation, velocity dispersion and the higher order terms of the Gauss–Hermite series simultaneously. The method makes use of a penalization factor to the solution, which biases the results to suppress the noise. Obviously, one would like the penalty to leave the LOSVD virtually unaffected. Thus, we first perform the kinematics fit without including any penalty (i.e. setting the keyword `BIAS = 0` in the `PPXF` code) to make sure that our results have not been biased and also to properly estimate the uncertainties on the velocity distribution. This produces a noisy solution (especially for σ) but allows us to set a range of values for h_3 and h_4 which can be used to estimate the right value for the penalization factor.

In particular, this has been done with the ‘`ppxf simulation`’ procedure (included in the publicly available code). The routine performs a Monte Carlo simulation of the spectra, which repeats the full fitting process for a large number of different realizations of the data obtained by randomly adding noise to the original spectra with a resulting S/N equal to the chosen threshold value ($S/N = 70$ in our case), and using as input for $[h_3, h_4]$ the maximum representative values measured in the non-penalized `PPXF` fit of the previous step ($[-0.3, 0.3]$). At this point, we repeat the fitting procedure while using as penalty the largest value such that the mean difference between the output and the input $[h_3, h_4]$ is well within the rms scatter of the simulated values ($BIAS = 1.5$).

Finally, we inspect by eye zones of the FOV where we expect to have contamination (e.g. near the brighter stars) and spectra for which the reduced χ^2 of the fit was larger than 3. In particular, we identify ~ 30 ‘problematic’ Voronoi cells, for which we go back to the single-fibre spectra and re-create the Voronoi cell while excluding contaminated or noisy spectra. Then we re-run the `PPXF` procedure, in some cases also setting by hand the pixel regions where the fit is performed.

We take into account all the main systematic uncertainties, mostly caused by template mismatch, and wavelength coverage. In order to minimize template-mismatch errors, we fit a broad set of stellar templates together with the kinematics, after convolving both to a common resolution. We also test the robustness of the kinematic results by changing the spectral region of the fit and we find results always consistent within 2σ (not including systematics). The uncertainties on the final inferred kinematics are estimated by adding in quadrature the formal uncertainty given by the Monte Carlo simulation and the scatter in the results for different templates and spectral regions. Typically, we find uncertainties of the order of 5 per cent for rotation and velocity dispersion and at most of the order of 10 per cent for the higher order LOSVD moments. For the central bin, we obtain a velocity dispersion of $\sigma_* = 161 \pm$

8 km s^{-1} perfectly consistent with previously published results (i.e. Binney, Davies & Illingworth 1990; de Lorenzi et al. 2008).

3.3 2D kinematical maps

The reconstructed mosaic ($\sim 80 \text{ arcsec} \times 60 \text{ arcsec}$ effective area) is among the largest 2D maps ever measured for an elliptical galaxy and provides a powerful tool to fully characterize the internal kinematics of the system and the spatially resolved properties of its underlying stellar population, which are necessary ingredients to understand the dynamics and its connection with formation mechanisms. Here below, we characterize the 2D behaviour of all the relevant kinematical quantities, while in the next sections we discuss the properties of the stellar populations and try to establish their connection with the kinematical structure. The full dynamical analysis as well as the derivation of the mass and orbital profile of the galaxy will be presented in a forthcoming paper.

Fig. 2 shows 2D maps of rotation velocity, velocity dispersion and the Gauss–Hermite moments h_3 and h_4 of the LOSVD. On the left column, these quantities are shown as derived from the

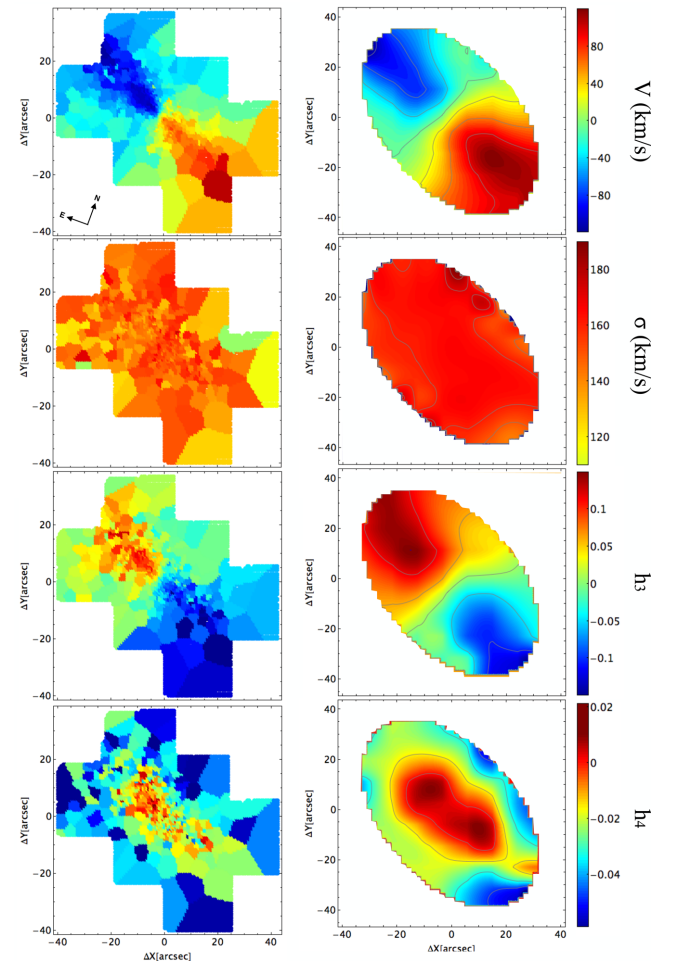


Figure 2. NGC 4697 2D LOSVD maps in an $80 \text{ arcsec} \times 80 \text{ arcsec}$ box. Panels on the left column show the unsmoothed Voronoi-rebinned version, as derived in Section 3.1. Here the small-scale features are visible. Panels on the right column show the 2D maps smoothed by averaging the Voronoi values over a grid of 15×15 macro bins and then masked with an ellipse mimicking the orientation and the axial ratio of the outermost isophote enclosing the VIMOS mosaic. Here the large-scale pattern of the kinematical quantities is more appreciable.

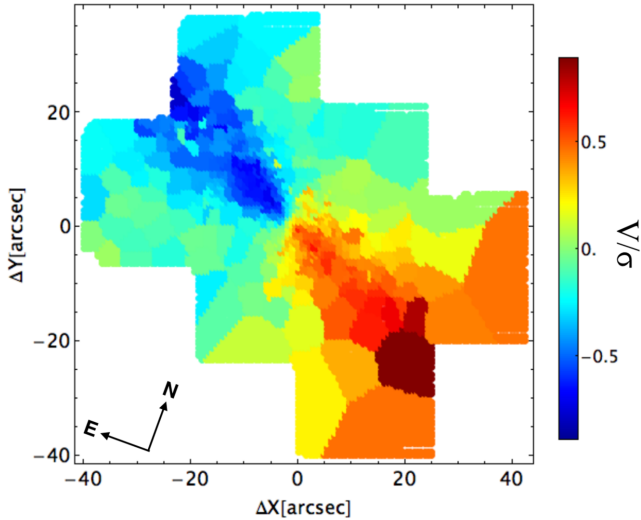


Figure 3. Top panel: V/σ 2D Voronoi map, quantifying the ratio of the ordered to random motion of NGC 4697.

Voronoi-binning technique discussed in Section 3.1. Here the small-scale features are visible, mainly in the centre, while the large-scale structure at larger radii looks noisy. In order to make more evident the large-scale pattern of the kinematical quantities, we show on the right column the 2D quantities smoothed by averaging the Voronoi values over a grid of 15×15 macro bins and then interpolated. The final average maps have been masked with an ellipse mimicking the orientation and the axial ratio of the outermost isophote enclosing the VIMOS mosaic.

The galaxy shows a rotation velocity (upper panel) and h_3 (third panel from above) patterns typical of a rotation-supported structure along the major axis of the system. In particular, the h_3 indicates the presence of a disc-like structure elongated along the major axis.

The velocity dispersion map (second panel), on the other hand, is rather flat over the full area covered by our observations, with minimal, statistically insignificant fluctuations.

Finally the h_4 parameter map (bottom panel) shows a clear elongated area along the line of nodes with positive values ($h_4 \sim 0.01$), which tend to become negative in the regions off the plane, which are not common for disc-like structure. The presence of negative h_4 off the main rotation axis in the regions where the rotation field is almost zero indicates tangential orbits, possibly belonging to the bulge component that starts to dominate outwards (see Section 4.1). The presence of two components along the line of sight might be responsible for the positive h_4 parameter aligned along the galaxy major axis by producing an overall peaked distribution due to the superposition of a cold disc and hot bulge.

IFS provides a complete and HR 2D map of the LOSVD, allowing the definition and estimation of a number of important parameters that can be used to better understand the detailed internal structure of the galaxy, a crucial ingredient for investigating possible formation scenarios. An interesting quantity that is possible to derive in 2D is the ratio of the ordered to random motion in a galaxy (V/σ).

In Fig. 3, we present the first V/σ 2D map of NGC 4697, up to $\sim 0.7R_{\text{eff}}$, which gives a more quantitative view of the dynamical state of the galaxy and further confirms the presence of the galaxy sub-components. One expects that the regions which are bulge dominated have smaller V/σ values, while the rotation-supported regions where the disc is dominant have larger V/σ val-

ues. This is indeed what we observe in the 2D map and it is further confirmed in the next section, where we compare our h_3-V/σ and h_4-V/σ correlations with those inferred from hydrodynamical simulations.

We note that none of the previous 3D measurements of galaxies have performed the same kinematical diagnostic as the one above. However, more standard rotation and dispersion 2D maps of the LOSVD of NGC 4697, like the ones derived from the ATLAS^{3D} survey, are fully consistent with the ones presented here, although with a lower spatial resolution and a smaller FOV (ATLAS^{3D} covers 40×60 arcsec; Krajnović et al. 2011). We do not directly compare our 2D maps to those obtained by ATLAS^{3D} because of the different spatial and spectral resolution but in Section 3.5 we extract 1D profiles along the major axis and compare our kinematics profiles against the ATLAS^{3D} ones.

Finally, extended 2D kinematical maps up to $3.5R_{\text{eff}}$ of this galaxy were recently presented as part of the SLUGGS Survey (Arnold et al. 2014). Overall there is fair agreement on the features observed in the data, both in the velocity field and in the higher order moments of the LOSVD. In fact, based on their more spatially extended kinematical map, the SLUGGS team concluded that the fast rotation is due to an embedded disc with a rotation amplitude clearly declining with radius. The presence of a disc is also found from our detailed analysis of the h_3-V/σ and h_4-V/σ diagrams.

3.4 Comparison with simulations

h_3-V/σ and h_4-V/σ correlations have also been predicted in simulations of disc merging (Naab, Jesseit & Burkert 2006a; H+09; N+14). For instance, N+14 linked the present-day central kinematic properties of galaxies to their individual cosmological formation histories and showed that galaxy mergers (major and minor) have a significant influence on the rotation properties, resulting in both a spin-down and a spin-up of the merger remnants, and in a clear anti-correlation between h_3 and V/σ .

N+14 have defined six classes (named with letters from A to F) with different evolutionary paths and histories of mass assembly that can be directly linked to their 2D kinematical maps and the present-day shapes. The galaxy groupings are based on the 2D kinematical maps and on the so-called spin parameter (λ_R), a proxy for the projected angular momentum per unit mass (Emsellem et al. 2007, 2011).

The λ_R parameter has been used to show that, within the effective radius, a dichotomy exists between ‘fast’ ($\lambda_R > 0.1$) and ‘slow’ ($\lambda_R < 0.1$) rotators (Emsellem et al. 2007). The second class of galaxies shows little or no rotation, significant misalignment between the photometric and the kinematic axes, and contains kinematically decoupled components, whereas the first class exhibits regular stellar velocity fields, consistent with disc-like rotation and sometimes bars. This version of the ‘spin’ parameter takes the form:

$$\lambda_R = \frac{\sum_{i=1}^{N_p} F_i R_i |V_i|}{\sum_{i=1}^{N_p} F_i R_i \sqrt{V_i^2 + \sigma_i^2}}, \quad (1)$$

where F_i , R_i , V_i and σ_i are the flux, circular radius, velocity and velocity dispersion of the i th spatial bin, with the sum running over the N_p bins.

A revised and more accurate separation of these two classes of galaxies has been proposed by Emsellem et al. (2011) to include the associated apparent ellipticity: a higher value of the specific stellar

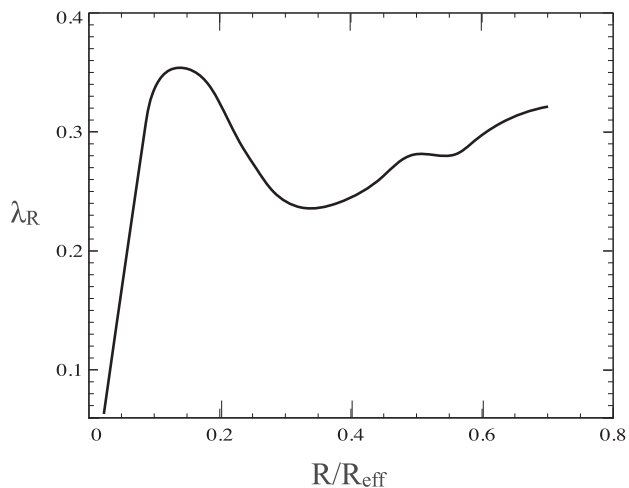


Figure 4. The λ_R parameter calculated using the formula of Emsellem et al. (2007) as a function of radius.

angular momentum is expected for galaxies which are more flattened or closer to edge-on if these are all intrinsically fast rotators.⁵

The threshold for λ_R was fixed to be proportional to the square root of ε with a scaling parameter that depends on the aperture (see appendix B in Emsellem et al. 2011 for further details). For instance, at $R_{\text{eff}}/2$

$$\lambda_{R_{\text{eff}}/2} = (0.265 \pm 0.01) \times \sqrt{\varepsilon_{\text{eff}}/2}. \quad (2)$$

Following this classification, NGC 4697 was previously classified as a ‘fast rotator’. This means that, given an ellipticity of $\varepsilon = 0.36$ (Peletier et al. 1990), one expects for this parameter a value of $\lambda_R > 0.16$ (for $r < R_{\text{eff}}/2$). We find $\lambda_{R_{\text{eff}}/2} = 0.35 \pm 0.05$, confirming the classification of this galaxy as a ‘fast rotator’ and in perfect agreement with the ATLAS^{3D} result ($\lambda_{R_{\text{eff}}/2} = 0.322$; Emsellem et al. 2011).

Fast rotators belong to class A, B and D in the N+14 classifications. However, only the first two groups of galaxies will result in h_3 and h_4 distributions similar to what we observe in NGC 4697. Both groups (A and B) of galaxies are rotationally supported, had a significant amount of central *in situ*, dissipative, star formation but, while class A have late assembly histories dominated by minor mergers, class B galaxies have gone through a late gas-rich major merger. These two formation scenarios will produce the same h_3 - V/σ anti-correlation that we observe in our maps.

One way to distinguish between the two classes is to look at the λ_R profiles: class A galaxies show peaked λ_R profiles within the effective radius, originating from the fast-rotating central stellar disc, whereas class B galaxies have λ_R profiles constantly rising beyond the effective radius.

In Fig. 4, we plot the λ_R profile as a function of radius calculated from circular apertures. The profile has a strong peak around $0.2R_{\text{eff}}$ and then it rises until $\sim R_{\text{eff}}$, which suggests that the mass assembly of NGC 4697 has involved significant *in situ* star formation, that NGC 4697 has experienced several gas-rich minor mergers (like simulated galaxies of class A) and possibly a late gas-rich major

⁵ A galaxy with a relatively low value of λ_R , e.g. of 0.2, may be consistent with a simple spheroidal axisymmetric system viewed at a high inclination (near face-on), but this is true only if its ellipticity is correspondingly low (e.g. $\varepsilon \leq 0.3$). A large ellipticity would instead imply a more extreme object (in terms of orbital structure or anisotropy).

merger leaving a previously rapidly rotating system unchanged (like simulated galaxies of class B). However, given the poor resolution of the simulations, this interpretation should be taken with caution [fig. 5 in N+15 shows the λ_R profiles of the simulated galaxies in the six classes and only one galaxy of class A shows a peak, which is at $\sim 0.95R_{\text{eff}}$. On the other hand, we observe that in fig. 5 of Emsellem et al. (2011), showing λ_R for the full ATLAS^{3D} sample, several fast rotators show a profile similar to what we derive for NGC 4697]. Moreover, Coccato et al. (2009) derived the λ_R profile for NGC 4697 out to $\sim 3.5R_{\text{eff}}$ based on PNe and showed that it flattens out at $\sim 1R_{\text{eff}}$ which supports the idea that NGC 4697 should be compared with simulated galaxies of class A. We also note that Arnold et al. (2014) concluded that galaxies like NGC 4697 formed through rapid, dissipative processes, which formed the central regions, followed by a prolonged period of dry minor mergers. We cannot exclude the scenario of gas-rich late major mergers, but independent lines of evidence point towards a scenario where the mass assembly of NGC 4697 is dominated by minor mergers and gradual dissipation.

In Fig. 5, we show the h_3 - V/σ and h_4 - V/σ correlations derived from the 2D kinematical maps (left-hand panels), and we compare them to luminosity contours derived by H+09 for a 1:1 disc–disc merger with 30 per cent gas fraction (middle panels) and from the cosmological hydrodynamical simulations of N+14 (class A and class B, right-hand panels). Good agreement, at least on a qualitative level, is found between data and simulations, although we caution that H+09 only considered the case of equal-mass mergers.

VIMOS data points are colour-coded in accord with their value of $|\text{Cos}[\theta]|$, calculated with respect to the galaxy minor kinematical axis. We find a clear correlation between the $|\text{Cos}[\theta]|$ values and the regions where the disc component is elongated. Points oriented near the major axis (i.e. $|\text{Cos}[\theta]| \sim 0$) have also higher absolute values of h_3 and therefore are associated with the rotation-supported structure. Similarly, points with higher values of $|\text{Cos}[\theta]|$ have lower h_3 and roughly correspond to the bulge region. The h_4 does not show a similar (anti-)correlation with V/σ .

The h_3 - V/σ anti-correlation generically traces cold discs. It is a feature that naturally emerges in disc–disc gas-rich mergers, and tends to produce large asymmetric deviations from Gaussianity because of the presence of sub-components that retain a memory of the disc-star streaming and because of gas that settles down through dissipation to form a thin embedded disc. This h_3 - V/σ correlation is generally negative when the dominant orbit population tends to stream (e.g. z -tube orbits; see also Röttgers, Naab & Oser 2014), and positive when the rotation is caused by sub-dominant streaming population (e.g. x -tube orbits) superimposed on the tail of the dominant streaming orbits (e.g. z -tubes which still contribute some rotation along the minor axis; see also the discussion in H+09). The same features are not found in second generations of major (re-)mergers (mergers at lower redshift, where less cold gas is present and with elliptical-morphology progenitors), in which case the features are more mixed and look more concentrated around the centre of the diagram.

The h_4 - V/σ correlation of NGC 4697 in Fig. 5 shows a V-shaped distribution with a double peak at $|V/\sigma| \sim 0.5$. This distribution is very similar to the predictions from simulations of N+14 (right, bottom panel) and it also agrees with H+09 (middle, bottom panel), although there is a significant offset between the data and the H+09 predictions. The peaked population could correspond to a late-formed disc stars (note the correspondence to the extreme h_3 values and $|\text{Cos}[\theta]| \sim 0$ indicating the alignment along the major axis) which tend to produce a positive h_4 in superposition with the

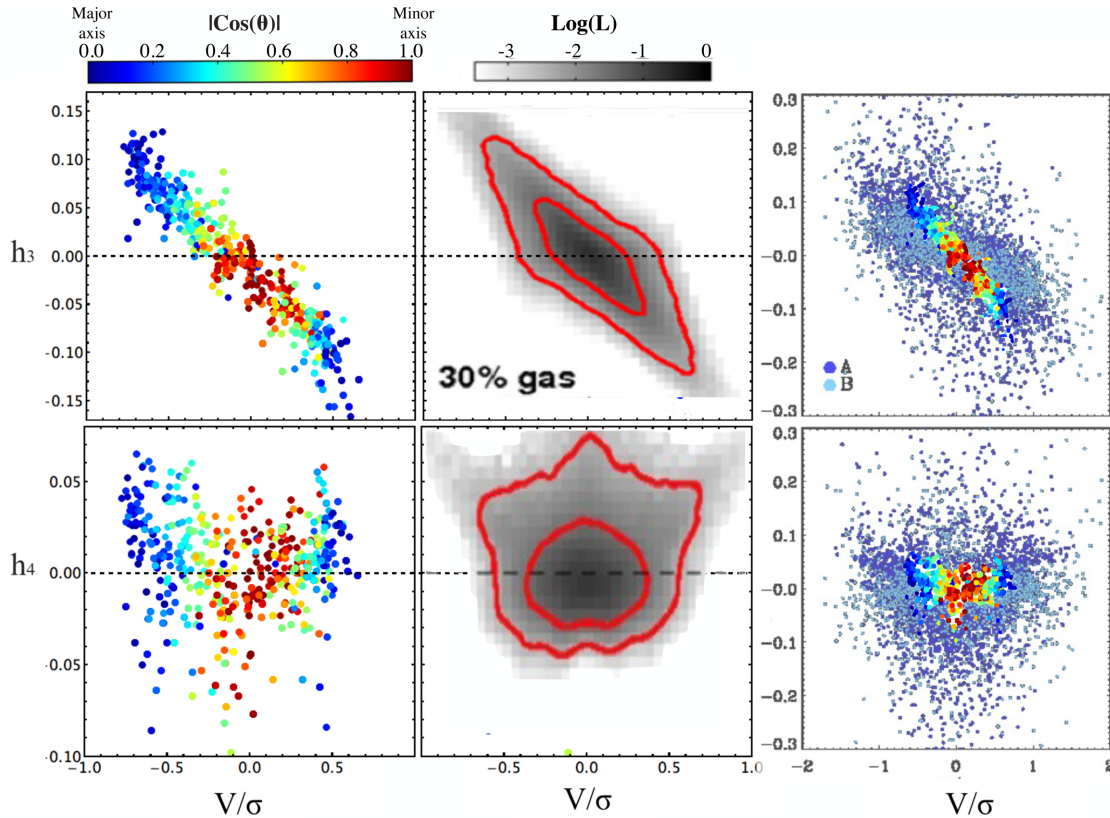


Figure 5. Left-hand panels: the h_3-V/σ (top) and the h_4-V/σ (bottom) distributions from the VIMOS kinematical maps. Points are colour-coded according to the azimuthal deviation from the galaxy minor axis. Middle panels: comparison with predictions from 1:1 disc–disc merger simulations from H+09 with 30 per cent gas fraction. The shading is proportional to the logarithm of the luminosity, whose scale is plotted in the bar on the top and the red contours contain 68 and 95 per cent of the luminosity. Right-hand panels: comparison with predictions of N+14 simulations for galaxies of class A and B.

more isotropic component (differences in ages of the two components are discussed in Section 4). In the simulations of H+09, the re-mergers would look more concentrated around the origin and would tend to have LOSVDs with positive h_4 , similarly to expectations from violent relaxation, which generally results in a radially biased distribution function (Binney & Tremaine 2008).

To conclude, the high-order moments of the LOSVD provide kinematical evidence for a two-component system: the largest absolute values of h_3 and h_4 indicate an area of the galaxy with the dominance of x -tube over z -tube orbits, which are produced in collisionless mergers and trace the presence of a (later) streaming structure; these are also the regions around $|\text{Cos}[\theta]| \sim 0$ (minor axis). Values with $h_3 \sim 0$ and $h_4 \sim 0$ indicate instead the presence of a non-rotating subsystem that dominates in the regions where the values of $|\text{Cos}[\theta]|$ are higher. The h_3-V/σ and the peaked λ_R profile suggest that gas-rich minor mergers dominate the mass assembly of NGC 4697, although we cannot completely exclude the possibility of a major merger event that happened at redshift $z \geq 1$.

We make use of this evidence in Section 4.1 where, based on the photometry of NGC 4697, we separate the disc from the bulge and study the ages and metal abundances of the stellar populations of these two components.

The details of the gas accretion leading to the formation of the central disc may be somewhat different from the one investigated in the simulation of disc merging and it might be useful to have LOSVD predictions from simulations of cold accretions as an alternative high dissipative process forming the disc. It seems evident, though, from H+09 and N+14 that dry mergers are not able to

produce the kinematical substructures we see in the h_3-V/σ and h_4-V/σ diagrams. In fact, in a merger without a dissipative gas component, the remnant galaxies are axisymmetric and do not show a population of stars with high angular momentum tube orbits nor the re-growth of a significant disc component (e.g. class D in N+14).

3.5 1D spatially resolved kinematics profiles

To compare our results against published results based on long-slit data and IFS with different FOV and resolution, we extract from the 2D kinematical maps the information along the major axis. Specifically, we extract and sum all the spectra positioned in a rectangle with width of 1.34 arcsec (2 pixels) oriented along the major axis of the galaxy and centred on our kinematic centre (to be as close as possible to the configuration of slit-based studies). We perform the same extraction for the IFU ATLAS^{3D} observations (publicly available on the survey webpage) in order to perform a simple comparison of the two sets of data with different spatial and spectral sampling. In Fig. 6, we compare our profiles (black points) with the ATLAS^{3D} ones (green points) and with the ones published in de Lorenzi et al. (2008, hereafter DL08), derived from high-S/N integrated absorption line spectra obtained by Méndez et al. (2005) with FORS2 at the VLT, a slit width of 1 arcsec and seeing of 1–1.5 arcsec (blue points). For the major axis inclination, we assume the value PA = 63 deg from Binney et al. (1990), slightly different from the one assumed by DL08 but providing the best symmetry and maximum rotation for our extracted 1D profiles. We do not worry about this small PA discrepancy that could be caused by the fact that

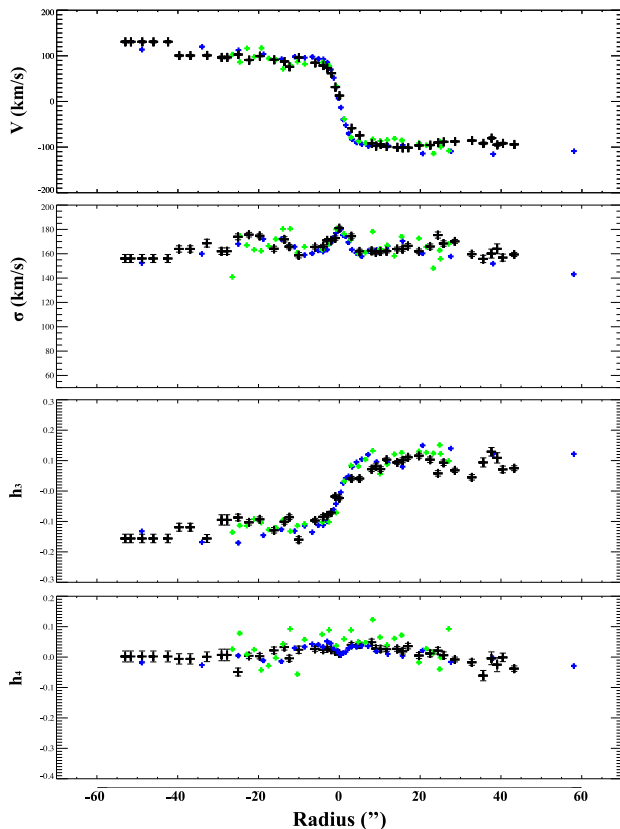


Figure 6. NGC 4697 1D kinematics profiles along the major axis. Black points are data from this work, blue points are long-slit data from de Lorenzi et al. (2008) and green points are IFU data from the ATLAS^{3D} survey, extracted along the major axis with the same procedure as used to extract our points. Good agreement is found between the different studies. Uncertainties on the rotation and velocity dispersion are of the order of 10 km s^{-1} , while on the h_3 and h_4 parameters are of the order of 10 per cent.

the slit in DL08 was not exactly on-centre or could be related to the offset we found between kinematic and photometric centre, most probably caused by a small effect of atmospheric refraction. For the ATLAS^{3D} data, we use the same inclination of $\text{PA} = 63$ deg and extract the major axis profiles with the same procedure applied to our VIMOS data. We found overall good agreement between the three 1D kinematical profiles which nicely demonstrate the robustness of our mosaic approach and the validity of our results.

4 2D STELLAR POPULATION ANALYSIS

In this section, we study the stellar population of the disc and of the bulge of NGC 4697 separately, to put constraints on the age, $[\alpha/\text{Fe}]$, metallicity and IMF normalization of the stellar population of the different morphological components of the galaxy.

To constrain the stellar population parameters, we focus on line-index measurements rather than full spectral fitting, which avoids issues with spectral calibration when comparing to observations that might have been poorly calibrated and we assume a simple stellar population (SSP), rather than an extended star formation history, which is still a strong assumption and will be properly addressed in future work.

We measure equivalent widths (EWs) of several stellar absorption features in each of the Voronoi-rebinned 1D spectra, after convolving them to a common resolution of 350 km s^{-1} to correct for

kinematic broadening.⁶ We compared these EWs to those measured from an extended version of the SSP models of Conroy & van Dokkum (2012, hereafter CvD12) with the same common resolution, obtained in Spiniello et al. (2015b).

CvD12 optimized models over the wavelength interval $0.35 \mu\text{m} < \lambda < 2.4 \mu\text{m}$ at a resolving power of $R \sim 2000$, specifically for the purpose of studying old, metal-rich stellar populations. The models are built combining two empirical stellar libraries (MILES, Sánchez-Blázquez et al. 2006; and IRTF, Cushing, Rayner & Vacca 2005) and three sets of isochrones to measure the IMF slope down to $\sim 0.1 M_{\odot}$. The CvD12 models explore variations in age in the range 3–13.5 Gyr, α -enhancement of 0–0.4 dex and four different single-slope IMFs: a bottom-light single-slope power law with $x = 1.8$ (where x is the IMF slope, using $dN/dm \propto m^{-x}$) that produces a stellar mass-to-light ratio (M^*/L) comparable to the one obtained from a standard Chabrier IMF (Chabrier 2003), a Salpeter (1955) IMF with a slope of $x = 2.35$, and two bottom-heavy (dwarf-rich) IMFs with slopes of $x = 3.0$ and 3.5 .

The CvD12 models also allow for variations in the $[\alpha/\text{Fe}]$ ratio as well as variations in the abundance pattern of 11 different single elements. This is enough to decouple IMF variations from abundance variations. However, all of the CvD12 models use solar metallicity isochrones, even when synthesizing with different abundance patterns or different $[\alpha/\text{Fe}]$. The abundance variations of single elements are implemented at fixed $[\text{Fe}/\text{H}]$, which implies that the total metallicity Z cannot be easily controlled. This is restrictive, if one aims to disentangle elemental enhancements, metallicity changes and IMF variations, especially in the case of massive ETGs, which have star formation histories different from the solar neighbourhood (e.g. Peterson 1976; Peletier 1989; Worthey 1992; Trager et al. 2000a; Arrigoni et al. 2010).

To resolve this issue, we follow the approach presented in Spiniello et al. (2015b, hereafter S+15) of extending the parameter space of the SSP models using response functions. We construct a metallicity response function for each given age and IMF slope taking the ratio between two spectra of the MIUSCAT SSP models (Vazdekis et al. 2012)⁷ with the same age and IMF slope but different total Z . We then multiply this response function with the spectrum of a CvD12 model with fixed age and IMF to build a new model (SSP) that extrapolates the latter model to a new part of parameter space (i.e. covering super-solar Z). These modified SSP models combine the flexibility of both the MIUSCAT and CvD12 models to predict spectra in a part of parameter space that neither of them reaches separately (MIUSCAT models use different total metallicities but do not allow one to change the abundances which are fixed to solar). As pointed out in S+14 and S+15, selecting IMF-dependent features that are age and metallicity independent and combining them with indices that depend mainly on age or mainly on element abundance is important for breaking the age–metallicity–IMF degeneracy when using SSP models to infer the stellar populations from unresolved galaxies spectra.

⁶ This final resolution has been chosen for practical reasons: the version of the SSP models used in this paper was created and used in S+14 with a resolution of 350 km s^{-1} . To save computer power, we did not perform again the EW measurements and simply smoothed our data to the resolution of the models. We are confident that this did not cause any loss of information, and we check (by eye-inspection) that contamination from sky line residuals does not affect the spectral features.

⁷ The MIUSCAT SSP has instead a fixed total metallicity and covers six metallicity bins $Z = 0.0004, 0.001, 0.004, 0.008, 0.019$ and 0.03 , where 0.019 represents the solar value.

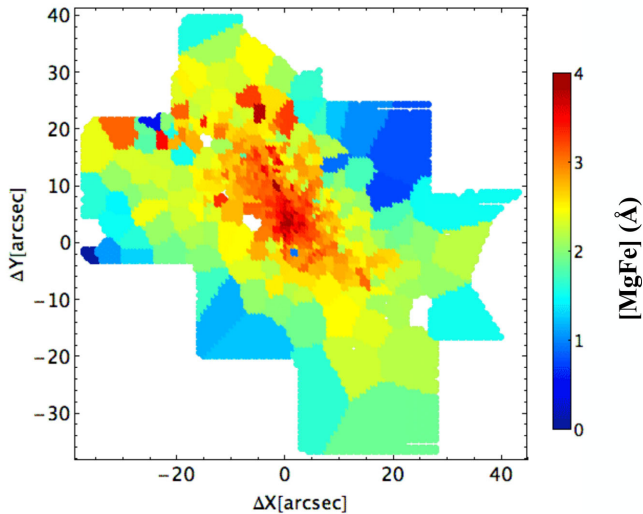


Figure 7. NGC 4697 2D maps of the [MgFe] stellar absorption feature measured from the spectra at a final resolution of $\sigma = 350 \text{ km s}^{-1}$. The map clearly shows that the stellar population of the disc is different from the stellar population of the bulge component of the galaxy. Areas within the disc are in fact more metal rich (higher [MgFe] EWs).

We measure some of the classical blue Lick indices (e.g. $H\beta$, Mgb , $Fe5270$, $Fe5335$, [MgFe]; González 1993; Worthey 1994; Trager et al. 2000b) that have been extensively and successfully used in the literature to constrain the age and metallicity of galaxies, together with some of the bluer IMF-sensitive features defined and used in S+14 and S+15. Unfortunately, this paper makes use of data taken with the HR-blue grism, and therefore only covers the bluer IMF features of S+14 (specifically bTiO, aTiO, TiO1) while in the future we plan to obtain new data using the HR-red grism which will cover also TiO2, CaH1, TiO3 and TiO4, allowing a more detailed analysis and a more secure inference on the low-mass end of the IMF slope.

Before measuring EWs, we convolved galaxy and model spectra to an effective velocity dispersion of $\sigma = 350 \text{ km s}^{-1}$ to correct for kinematic broadening. Indices in both the galaxy and the model spectra have been measured with the same definitions and methods.⁸ We do not place our indices on the zero-point system of the Lick indices and quote them as index strengths in units of angstrom (if atomic) or magnitudes (if molecular, like the TiO lines).

The great advantage of integral field spectroscopy is the simultaneous availability of spatial and spectral information. This allows us to generate 2D spatial maps of spectral features of interest, like the one in Fig. 7, showing the [MgFe]9 index that has been widely used to constrain the metallicity. The 2D map suggests different stellar populations in the bulge and in the disc: specifically, the [MgFe] EWs are systematically higher in the Voronoi cells within the disc. Therefore, the 2D element maps indicate that the stars in the disc are more metal-rich than the stars in the bulge. However, we caution the reader that it is not possible to directly infer the total metallicity by looking at a single-element map, because of the degeneracies present between the stellar population parameters. For instance, in this case the [MgFe] values could be higher due to an increase of age and/or could increase of total metallicity. This is an example of the so-called age-metallicity degeneracy which can be broken by

using more than one indicator. This motivates us to decompose the bulge and disc components and to extract a 1D spectrum of each of them in order to study their stellar populations separately through a more quantitative, proper SSP analysis.

4.1 Disc–bulge decomposition

We perform a two-component fit to the *HST* V-band image of NGC 4697 obtained by Tal et al. (2009), driven by the correlations in the h_3 and h_4 parameters versus V/σ that clearly indicate the presence of a disc. We use the latest version of GALFIT (Peng et al. 2010) to fit two ellipsoid models (‘Sérsic’ for the bulge and ‘expdisk’ for the disc) to the light profiles of the galaxy image. We obtain a similar position angle (PA ~ 64 deg, confirming previous findings) but very different axial ratios for the two components: $(b/a)_{\text{disc}} = 0.297$, $(b/a)_{\text{bulge}} = 0.7512$. We use those to spectroscopically separate the bulge from the disc. We note that a more rigorous spectroscopic separation, following the approach of Coccato et al. (2011) and Coccato, Iodice & Arnaboldi (2014), will be performed in future studies to fully separate these two components.

First, we extract and sum together all the spectra in the data cube with $r < R_{\text{eff}}/8$. In this central region, we expect bulge and disc to both contribute substantially. Secondly, we select and sum the spectra belonging to an elliptical annulus centred on the kinematics centre with the axial ratio and PA found by GALFIT for the disc: $(b/a)_{\text{disc}} = 0.297$, PA = 65.47 deg and with $r > R_{\text{eff}}/8$ (to exclude the central region). Finally, we sum up the remaining spectra where the contribution of the bulge is dominant. The resulting 1D spectra have very high S/N (S/N = [450, 320, 280] for centre, disc and bulge, respectively).

In Fig. 8, we show the extracted 1D spectra of the disc and bulge, as well as of the mixed (disc+bulge) population of the central region of NGC 4697 for comparison. Shaded yellow bands show some of the line indices used in this work, along with continuum side bands (blue and red coloured bands). Some small differences are visible already from the spectra: the disc spectrum shows stronger Mg and Fe lines, which confirms what we have already seen from the [MgFe] 2D map and gives hints of a more metal rich population. The Na D absorption line, the narrow strong line around 5900 Å, is stronger in the bulge than in the disc, consistent with the hypothesis that bulges could be Na-enhanced (e.g. in the bulge of the Milky Way, on average [Na/Fe] ~ 0.2 ; Fulbright, McWilliam & Rich 2007). However, we do not use this particular line in the following analysis because its blue band is heavily contaminated from emission sky lines and because of the recent finding presented in S+15 showing that the Na D lines are model dependent and therefore must be interpreted with caution. The aTiO index has also been excluded from the analysis because it is contaminated from emission sky lines.

In Fig. 9, we show index–index plots of the selected optical indicators to provide qualitative constraints on the stellar population parameters. In panels (a) and (b) the IMF dependence is minimal and the age and abundances of the galaxy can be inferred. The IMF normalization of the galaxy can be instead inferred from panels (c) and (d), using SSP models with the metallicity and age previously determined.

4.2 Age and metallicity of the stellar population

In the upper panel of Fig. 9, we show index–index diagrams of some of the classical Lick indices. These diagrams are widely used to constrain age and metallicity of integrated stellar populations.

⁸ We use the code SPINDEX from Trager, Faber & Dressler (2008).

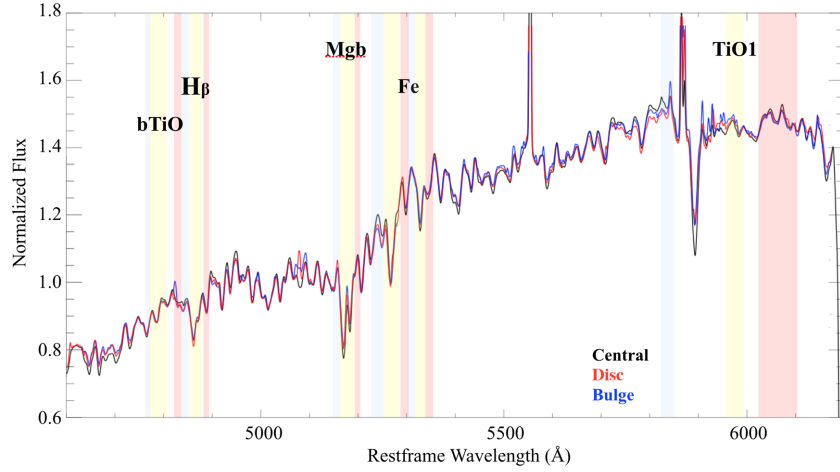


Figure 8. 1D spectra extracted from the bulge (blue), disc (red) and the central $\sim R_{\text{eff}}/8$ region (10 arcsec, black). Yellow boxes show the bandpasses of the line indices used to constrain the stellar population parameters, while blue and red bands represent the blue and red continuum side bands, respectively.

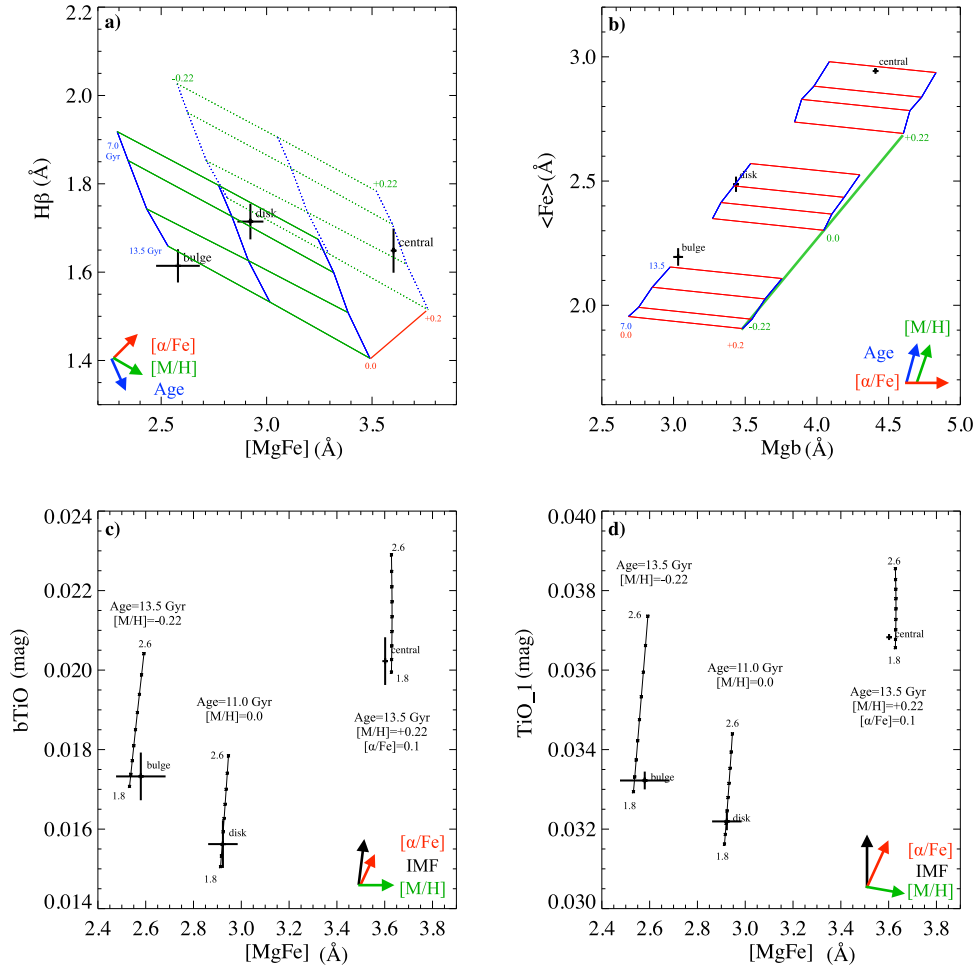


Figure 9. Index–index plots of some of the classical stellar absorption features and of the bluer IMF-sensitive optical indicators defined in S+14 in the optical regime. Grids are CvD12 extended SSP models obtained in S+15 via the response function approach (see the text for further details). Blue lines show models with varying ages with values [7,9,11,13.5] Gyr, red lines are SSP models with varying $[\alpha/\text{Fe}]$: [0.0,0.2,0.4] dex, green lines are SSP models with varying total metallicity ($[M/\text{H}] = [-0.22, 0.0, +0.22]$) and, finally, black lines are models with varying IMF slope from $x = 1.8$ (Milky Way-like) to $x = 2.6$ (being Salpeter $x = 2.35$). Coloured arrows in the corner of each panel show the direction of the variation of these three single parameters. Points with error bars are high-S/N galaxy spectra extracted from the different regions by summing the Voronoi-rebinned spectra. Galaxies and models are convolved to a final common resolution of 350 km s^{-1} before measuring index strengths. In the plots on the upper row (panels a and b), IMF dependence is minimal (all models have $x = 1.8$ here) and the age and abundances of the galaxy can be inferred. The IMF normalization of the galaxy can instead be inferred from panels (c) and (d), using SSP models with the metallicity and age constraints previously determined. Here, lines are CvD12 models with varying IMF slope in the range 1.8–2.6.

Table 1. Best-fitting stellar population models for the different components obtained via a χ^2 fitted with CvD12 SSP models with a $x = 1.8$ IMF slope.

Component	Age (Gyr)	[M/H]	[α /Fe]
Central	$12.9^{+1.5}_{-1.8}$	$+0.23^{+0.08}_{-0.05}$	$+0.09^{+0.03}_{-0.01}$
Bulge	$13.5^{+1.4}_{-1.4}$	$-0.17^{+0.12}_{-0.1}$	$+0.03^{+0.03}_{-0.01}$
Disc	$10.5^{+1.6}_{-2.0}$	$-0.03^{+0.02}_{-0.1}$	$+0.04^{+0.02}_{-0.01}$

The average indices extracted from the different regions are plotted on top of grids of CvD12 SSP models with varying ages (blue lines, [7,9,11,13.5] Gyr), [α /Fe] (red lines, [0.0, 0.2, 0.4] dex) and total metallicity (green lines, [-0.22, 0.0, +0.22]). All plotted models have a Chabrier-like IMF, but for these indicators ($H\beta$, Mgb , Fe5270, Fe5335) the dependence on IMF is minimal.

Panel (a) of Fig. 9 shows the $H\beta$ –[MgFe] diagram. The $H\beta$ index is one of the most used age indicators in the optical, and allows one to break the age–metallicity degeneracy, when combined with the commonly used [MgFe] index. However, this diagram alone does not allow one to also break degeneracies with [α /Fe]. This is clear for instance in the case of the disc component. From panel (a) it is inferred that the bulge has an old population with solar [α /Fe] and slightly sub-solar metallicity. The central region is also old but has super-solar abundances. For the disc, the interpretation is somewhat more difficult. In fact, the disc spectrum matches a model with age $=11 \pm 2$ Gyr, slightly super-solar metallicity and solar [α /Fe] but also a model with older age, lower metallicity and α -enhanced. To discriminate between these two possibilities, the reader should refer to the results plotted in panel (b). Panel (b) shows Mgb versus $\langle Fe \rangle$, which is the mean of Fe5270 and Fe5335. Here the disc matches only with a model with age $=11 \pm 1$ Gyr, solar metallicity and solar [α /Fe].

Index–index plots are a useful tool to give qualitative inference on the stellar population parameters, but it is sometimes hard to break degeneracies and quantitatively constrain age, metal abundance and IMF slope at the same time. Therefore, in order to give more quantitative constraints on the stellar population parameters of the three components, we perform a χ^2 minimization comparing the 1D spectrum of each component with grids of linearly interpolated SSPs spanning a range of ages (between 7 and 13.5 Gyr with steps of 1 Gyr.), [α /Fe] (between 0.2 and +0.4 dex, with a step of 0.1 dex) and total metallicity ([M/H] = [-0.22, 0.0, +0.22]) for a Chabrier IMF (we note that a moderate change in the IMF slope has only a negligible effect on these blue features). We find that the disc is slightly younger and more metal rich than the bulge, consistent with an SSP model of $\sim 10.5^{+1.6}_{-2.0}$ Gyr, solar metallicity ([M/H] = $-0.03^{+0.02}_{-0.1}$) and solar [α /Fe] ([α /Fe] = $+0.04^{+0.02}_{-0.01}$). The bulge fits better with a $13.5^{+1.4}_{-1.4}$ Gyr model with sub-solar metallicity ([M/H] = $-0.17^{+0.12}_{-0.1}$) and solar [α /Fe] ([α /Fe] = $+0.03^{+0.03}_{-0.01}$). The central region shows stellar population properties in between those of the disc and the bulge, as expected. It is as old as the bulge ($12.9^{+1.5}_{-1.8}$ Gyr), shows a solar or slightly super-solar [α /Fe] ([α /Fe] = $+0.09^{+0.03}_{-0.01}$) and it is even more metal enriched ([M/H] = $+0.23^{+0.08}_{-0.05}$). These SSP results nicely confirm the qualitative picture obtained via index–index plots in the upper row of Fig. 9, and are summarized in Table 1. However, we note that since the 1D spectra of bulge and disc are composite spectra where the contributions of the two components are not perfectly disentangled, the true stellar population differences are likely to be even greater.

Finally, we consider NGC 4697 as whole and calculate gradients in its stellar populations. Following the approach of Tortora et al. (2010a), we define the stellar population gradient as the dimensionless coefficient of the relation

$$\log X - \log R/R_{\text{eff}} \left(\nabla_X = \frac{\delta(\log X)}{\delta \log(R/R_{\text{eff}})} \right), \quad (3)$$

where X represents the age or the total metallicity expressed in terms of [M/H]. We find a non-negligible but very small age gradient and a negative metallicity gradient in the sense that NGC 4697 is more metal rich and older in the centre. These gradients are consistent with the measured gradients in old ETGs (Kuntschner et al. 2010; Tortora et al. 2010a) and with remnants of gas-rich major mergers from simulations (e.g. H+09; Hopkins et al. 2009).

4.3 The IMF normalization

Using the metallicity and age constraints obtained from the Lick blue indices via the χ^2 minimization, we focus now on two of the IMF-sensitive optical indicators that fall in the VIMOS wavelength range, the bTiO and the TiO1. We plot them against the [Mg/Fe] index in the lower panels of Fig. 9. The combination of the latter indicator with IMF-sensitive features breaks the degeneracies between the [α /Fe], total metallicity and IMF slope variations.

For each component, we use SSP models with varying IMF slope (x between 1.8 and 2.5) and [α /Fe], and with the other stellar population parameters as close as possible to the ones derived from the χ^2 fit to the blue indices. For the bulge, we use SSP models with 13.5 Gyr and [M/H] = -0.22, and for the disc we use SSP models with 11.0 Gyr and [M/H] = 0.0 (solar). Finally, for the mixed, central population, we use SSP models with 13.5 Gyr and [M/H] = +0.22 (super-solar).

In Fig. 9, black lines show models with different IMF slopes and red lines show models with varying [α /Fe]. No detectable variation in the IMF normalization between bulge and disc is visible from the two panels. Both components of NGC 4697, as well as the central region, are consistent with a normalization of $x = 2.0 \pm 0.2$, slightly dwarf richer than a Chabrier IMF ($x = 1.8$, still consistent within 1σ) but shallower than a Salpeter ($x = 2.35$), which is perfectly consistent with the slope predicted from the IMF– σ_* relation (e.g. La Barbera et al. 2013; Tortora et al. 2013; S+14) for an ETG with $\sigma_* = 160 \text{ km s}^{-1}$. IMF normalization does not change with radius for this galaxy is consistent with the results found in Martín-Navarro et al. (2015). The lack of IMF spatial gradient is also consistent with the results on the low-mass system in Martín-Navarro et al. (2015) that shows an IMF normalization constant with radius and consistent with a Kroupa-like profile.

4.4 Linking kinematics and stellar populations

The main conclusion from the stellar population analysis is that the inner region (typically $\sim R_{\text{eff}}/8$) looks coeval (or a bit older) and more metal rich than the regions farther out, which suggests that this central region has been formed early via rapid collapse. However, from the stellar population analysis alone, we cannot exclude minor or major merging of late-type disc galaxies, which have destroyed spiral arm features and left a younger disc component. It is therefore necessary to connect and correlate the stellar population results with those on the spatially resolved kinematics.

The SAURON team (Kuntschner et al. 2006; Krajnović et al. 2008) have found that the fast-rotating component almost always features a higher metallicity compared to the galaxy as a whole.

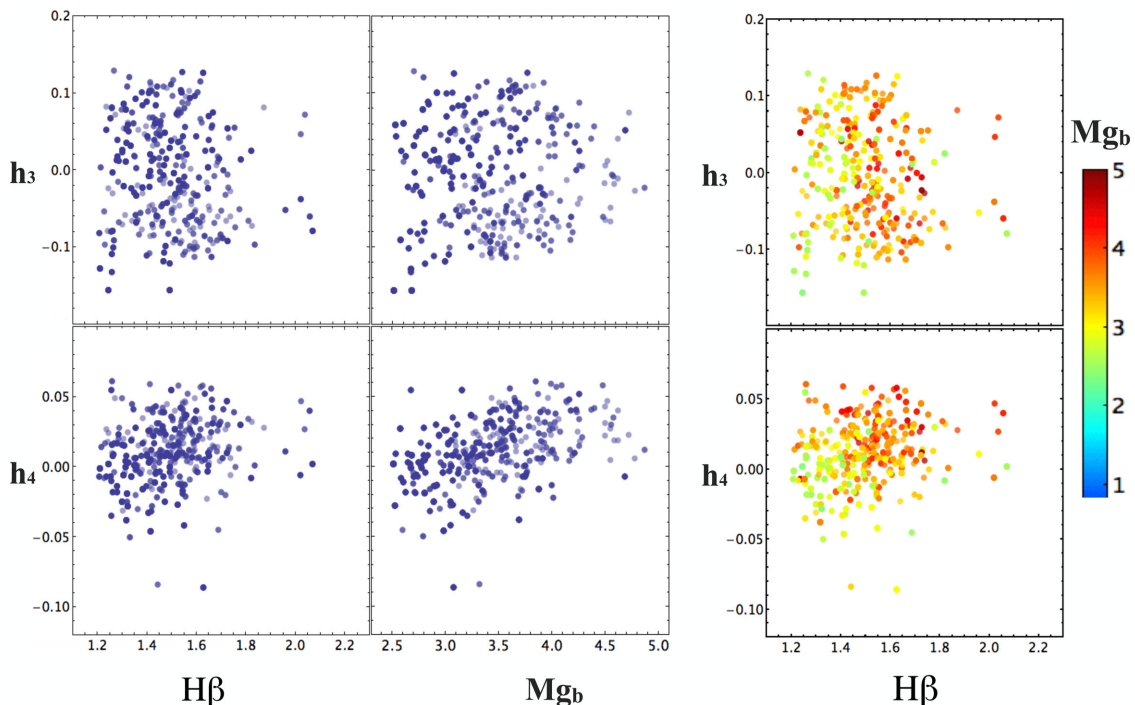


Figure 10. Left-hand panels: $H\beta$ EWs versus the h_3 (top) and the h_4 (bottom) parameter. Middle panels: Mgb EWs versus the h_3 (top) and the h_4 (bottom) parameter. In all panels, points represent values obtained from the macro bins of the smoothed maps and show a correlation between LOSVD properties and stellar population features. Larger $H\beta$ values roughly indicate younger ages, while larger Mgb values indicate higher total metallicity (although the well-known age–metallicity degeneracy makes things a bit more complicated than that; e.g. Worthey (1994)). Right-hand panels: $H\beta$ EWs versus the h_3 (top) and the h_4 (bottom) parameter, colour-coded by Mgb EW values. Points with larger, positive h_4 , aligned with the major axis (see Fig. 5), have on average higher $H\beta$ and Mgb . This is true also for points with larger h_3 values (in absolute value) although the plot looks a bit noisier. This figure highlights the agreement between kinematics and stellar population analysis both confirming that stars aligned with the major axis (where the disc dominates) are younger and more metal rich.

They have attributed this to the presence of a later-forming stellar population in a rotating structure with disc-like kinematics (Kuntschner et al. 2010). From the NGC 4697 2D maps as well as from the stellar population analysis presented in this study, we infer the same scenario. This is demonstrated in Fig. 10, which shows the correlation between the h_3 (top panels) and the h_4 (bottom panels) parameters and the $H\beta$ (left) and Mgb (middle) values obtained from the smoothed 2D maps. We run a Spearman test on these relations to better quantify the statistical dependence between the high-order moments of the LOSVD and $H\beta$ and Mgb , respectively. The resulting Spearman rank correlation coefficients are $\rho = 0.36$ (0.2) and $\rho = 0.32$ (0.36) for h_4 (h_3) and $H\beta$ and Mgb , respectively. Given the big sample size (more than 100 points), these values of ρ all exceed the critical value ($\rho_{\text{crit}}(n > 100) = 0.165$) and therefore the relations are statistically significant, although not strong.

The right-hand panels of Fig. 10 show the $H\beta$ – h_3 (top panel) and $H\beta$ – h_4 (bottom panel) diagrams with the points colour-coded according to the Mgb values. Although the discussion must be kept on a qualitative level because of the well-known age–metallicity degeneracy (Worthey 1994), we can simplify the situation and say that larger $H\beta$ values indicate younger ages and larger Mgb values indicate higher total metallicity. Thus, we can conclude that points with $h_4 > 0$, tracing the later-formed disc, are slightly younger and more metal rich, while points with $h_4 < 0$, roughly tracing the bulge component because they are along the minor axis of NGC 4697, all have generally older age and lower metallicity. The same holds also for points with larger $|h_3|$, although in this case the relations are less robust.

Remarkably, the spatially resolved stellar population results perfectly agree (on a qualitative level at least) with the spatially resolved kinematics results. Unfortunately, none of the hydrodynamical simulations derive 2D detailed stellar population properties. In the future, it might be extremely useful to have predictions from simulations on the ages and metal abundances of the remnant of different mergers scenarios.

In the SAURON ETGs of low to intermediate mass ($\sigma = 100$ – 160 km s^{-1}), the region of increased metallicity is connected to a mild decrease of the $[\alpha/\text{Fe}]$, whereas in more massive galaxies ($\sigma > 160 \text{ km s}^{-1}$) the disc formation and increased metallicity are typically constrained to a central location involving only a minor fraction of the total mass of the galaxy and only in about 25 per cent of the cases also a central depression in $[\alpha/\text{Fe}]$ can be seen.

In the case of NGC 4697, we find that the central region is slightly α -enhanced with respect to the bulge, whereas the disc shows roughly the same $[\alpha/\text{Fe}]$ as the bulge. A larger metallicity associated with the fast-rotating component seems to point towards a formation scenario where fast rotator galaxies have undergone one or multiple periods of secondary star formation at a later stage and over longer time-scales and therefore resulted in an increased metallicity and a decrease of the abundance ratio in a disc-like structure (Kuntschner et al. 2010).

A younger and more metal rich disc is also well consistent with the scenario suggested by the gas-rich merger (e.g. H+09), as the younger disc would arise when the gas, retaining its angular momentum during the merger, is fuelled in the centre of the remnant and eventually turns into stars. In Section 3.3, we have

demonstrated that the same scenario is also perfectly able to explain the kinematical substructures.

We note here that the disc merging scenario does not necessarily imply major merging, as also merging with lower mass fraction can produce similar kinematical features. In fact, simulations in N+14 have shown that correlations similar to the ones we observe in the h_3 and h_4 parameters of NGC 4697 can originate from two galaxy formation scenarios: a late gas-rich major merger or *in situ* star formation plus minor merging.

In conclusion, the combination of the kinematical features and the stellar population properties seems to point to the scenario where some fast-rotating and discy ellipticals can have formed from dissipative (wet) mergers of disc galaxies.

5 DISCUSSION AND CONCLUSIONS

In this paper, we have presented high-S/N integral-field kinematics and stellar population combined analysis of the elliptical galaxy NGC 4697, an almost edge-on, intermediate-luminosity (L_*) isolated elliptical galaxy. We used a mosaic approach, the first of its kind with VIMOS IFU data, obtaining 2D maps of stellar kinematics and absorption lines up to $0.7R_{\text{eff}}$. The final effective FOV of $\sim 80 \text{ arcsec} \times 60 \text{ arcsec}$ allowed us to obtain the largest 2D kinematical maps ever measured for a single ETG (in terms of arcsec) providing a unique view on the internal structure of the system, as well as crucial morphological and dynamical information.

Our main scientific findings are the following.

(i) We have measured the 2D velocity moments of the galaxy's stellar population finding clear evidence for a rotation-supported structure along the major axis from the rotation velocity and the third Gauss–Hermite coefficient h_3 (representing the skewness). We have measured the ‘spin parameter’ (λ_R) confirming the classification of this system as a ‘fast rotator’.

(ii) We have analysed the correlations between the h_3 and h_4 parameters with V/σ (Fig. 3), together with the λ_R spatial profile, demonstrating how these characterize the different kinematical components (bulge and disc) of the galaxy.

(iii) We have directly compared these correlations with those predicted by hydrodynamical simulations of H+09 from gas-rich disc–disc major mergers with 30 per cent gas fraction and of N+14 for gas-rich minor and major mergers (class A and B). The correlations we presented closely resemble the theoretical predictions and indicate that the mass assembly of NGC 4697 is dominated by gas-rich minor mergers, gradual dissipation and possibly a late gas-rich major merger.

(iv) A direct comparison with 1D long-slit spectroscopy (DL08) shows overall good agreement between the different kinematics results along the major axis.

(v) We extracted 1D spectra of the disc and the bulge of NGC 4697, using a two-component fit to a deep V-band image.

(vi) We studied the stellar populations of these two components using SSP models. We find indications that the stars in the disc are $\sim 2\text{--}3$ Gyr younger and more metal rich than the stars in the bulge.

(vii) We linked the stellar population analysis and the internal kinematics of the system by correlating age and metallicity of the stellar populations of the two components with the h_3 and the h_4 kinematic parameters, finding further evidence for a significant *in situ* formation of stars since $z \sim 2$ with additional gas-rich minor mergers (although we do not exclude the possibility of a late major merger that left a previously rapidly rotating system unchanged).

(viii) From the two bluer IMF-sensitive optical indicators of S+14, which fall in the VIMOS wavelength range, we do not detect any difference in the IMF slope between the bulge and the disc. Both components are consistent with an IMF with a single power slope of $x = 2.0$, consistent with a Kroupa IMF and slightly lighter than a Salpeter-like IMF. The inferred IMF slope is perfectly consistent with the one predicted from the IMF– σ_* relation for an ETG with $\sigma_* = 160 \text{ km s}^{-1}$, such as NGC 4697 (Treu et al. 2010; CvdD12; La Barbera et al. 2013; Tortora et al. 2013, 2014a; S+14).

The absolute novelty of this work resides (i) in the depth and width of the data acquired as we have very high S/N spectra over an effective area of the sky of $80 \text{ arcsec} \times 60 \text{ arcsec}$ (e.g. larger than MUSE FOV); (ii) in the presentation of a detailed spatially resolved kinematical and population 2D analysis within the central effective radius, which has allowed us to cross-correlate LOSVD information (i.e. h_3 and h_4) with relevant line indices (e.g. h_4 versus $H\beta$ and Mgb in Fig. 10), showing that points with $h_4 > 0$, associated mainly with the later-formed disc, are younger and more metal rich, while points with $h_4 < 0$, roughly tracing the bulge component as aligned with the minor axis of NGC 4697, all have generally a smaller $H\beta$ and smaller $[M/H]$.

The emerging picture is that a later-formed disc has been built from pre-enriched gas that formed stars after having dissipatively sunk into the centre. This is well consistent with the scenario suggested by gas-rich disc merging simulations (e.g. H+09), as the younger disc would arise when the gas, retaining its angular momentum during the merger, is fuelled in the centre of the remnant and eventually turns into stars. It is remarkable that within this scenario, photometry, stellar population features and kinematical features like the $h_3\text{--}V/\sigma$ and $h_4\text{--}V/\sigma$ correlations are all qualitatively correctly predicted. We stress again that H+09 only consider 1:1 disc–disc merger but that also merging with lower mass fraction (minor mergers) can produce similar kinematical features (e.g. N+14). Furthermore, other mechanisms involving a rapid dissipative collapse of enriched gas (clumpy discs, violent relaxation, cold accretion) might have produced a similar galaxy component.

This paper is intended as a pilot programme for a large project that aims at the study of internal structure, mass distribution and stellar population of a sample of nearby L_* ETGs. With medium- to high-S/N spatially resolved spectra, we will perform a detailed 2D study of the stellar population via line-index measurements and try to further test possible variation in the IMF normalization. We will attempt to characterize the low-mass end of IMF normalization in the bulge and in the disc components by measuring EWs of the IMF-sensitive features defined in S+14. In order to achieve this result, we plan to change the configuration of VIMOS in the new observations, choosing a redder filter (HR-red) to cover the proper wavelength range.⁹ In addition to a direct measurement of the M^*/L , thanks to the 2D integral-field coverage, we will be able to study age and metallicity radial gradients for our systems. Moreover, where spatial gradients in the stellar population parameters are present, the assumption of a spatially constant M^*/L clearly does not work. Consequently, the 2D coverage, allowing a more accurate measure

⁹ We are aware that moving to the red filter will cause the loss of some of the blue indices covered in this paper, making age determination more complex. We therefore plan to select ETGs where measurements of age and metallicity are available from the literature and mainly focus on measuring gravity-sensitive features to study the possible spatially variation of the IMF slope, with particular regard on bulge- versus disc-dominated zones.

of the M^*/L , also permits a more accurate and complete dynamical analysis (e.g. Tortora et al. 2010b, 2011).

A quantitative dynamical model to fit the whole 2D maps for the rotation and the velocity dispersion is out of the purpose of this paper but will be presented in future work. Only with a complete 2D kinematical map of a galaxy (up to at least $1R_{\text{eff}}$) as well as a robust dynamical model, it will be possible to measure the DM fraction, trace its radial distribution as a function of the mass of the halo and test plausible environment dependences. In future work, we therefore plan to perform a detailed study of the stellar content of ETGs to combine spatially resolved stellar information on the central region with PN data on the outskirts (up to ~ 5 arcmin) with the final purpose of obtaining the complete rotation curve and distribution of DM (e.g. Napolitano et al. 2009, 2011).

Moreover, thanks to a complete and detailed dynamical analysis, we will have a completely independent measurement of the IMF slope that we will compare to the one obtained by stellar population modelling. If both methods agree on the resulting IMF, this will be a strong test of the absence of hidden systematics.

Finally, another possible interesting follow-up will be studying the dynamics and the IMF normalization of this galaxy in the phenomenological framework of the Modified Newtonian Dynamics theory (Milgrom 1983a,b, following the approach presented in Tortora et al. 2014a).

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