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ABSTRACT

We derive the stellar population parameters of 11 quiescent ultra-diffuse galaxies (UDGs) from Keck/KCWI data. We supplement these with 14 literature UDGs, creating the largest spectroscopic sample of UDGs to date (25). We find a strong relationship between their α -enhancement and their star formation histories: UDGs that formed on very short time-scales have elevated [Mg/Fe] abundance ratios, whereas those forming over extended periods present lower values. Those forming earlier and faster are overall found in high-density environments, being mostly early infalls into the cluster. No other strong trends are found with infall times. We analyse the stellar mass–metallicity, age–metallicity, and [Mg/Fe]–metallicity relations of the UDGs, comparing them to other types of low mass galaxies. Overall, UDGs scatter around the established stellar mass–metallicity relations of classical dwarfs. We find that GC-rich UDGs have intermediate-to-old ages, but previously reported trends of galaxy metallicity and GC richness are not reproduced with this spectroscopic sample due to the existence of GC-rich UDGs with elevated metallicities. In addition, we also find that a small fraction of UDGs could be ‘failed-galaxies’, supported by their GC richness, high alpha-abundance, fast formation time-scales and that they follow the mass–metallicity relation of $z \sim 2$ galaxies. Finally, we also compare our observations to simulated UDGs. We caution that there is not a single simulation that can produce the diverse UDG properties simultaneously, in particular the low metallicity failed galaxy like UDGs.

Key words: galaxies: evolution – galaxies: formation – galaxies: kinematics and dynamics – galaxies: stellar content.

1 INTRODUCTION

Ultra-diffuse galaxies (UDGs) are a type of low surface brightness (LSB; Sandage & Binggeli 1984) galaxy, broadly defined by being extremely faint ($\mu_{0,g} > 24$ mag arcsec^{−2}), presenting a relatively large half-light radius ($R_e > 1.5$ kpc) and exhibiting dwarf-like stellar masses [$\log(M_*/M_\odot) \sim 7$ –8.5; see Van Nest et al. (2022) for an in-depth discussion on UDG definition and Watkins et al. (2023) for limitations on the use of this broad definition]. While galaxies fitting this definition had been previously discovered (e.g. Impey, Bothun & Malin 1988), limitations in the telescope facilities and data techniques at the time prevented a further characterization of their nature.

Nowadays, with instrumentation and techniques specifically developed to push the limits reachable in LSB studies, UDGs were re-discovered in 2015 by van Dokkum et al. (2015) and Koda et al. (2015). They were found by the hundreds in the nearby Coma cluster and since then they have been found to exist in all environments. They are seen in other galaxy clusters besides Coma (e.g. Mihos

et al. 2015; Mancera Piña et al. 2018; Iodice et al. 2020; Lim et al. 2020), in groups (e.g. Trujillo et al. 2017; Müller, Jerjen & Binggeli 2018; Forbes et al. 2019; 2020b; Marleau et al. 2021), and in the field (Martínez-Delgado et al. 2016; Bellazzini et al. 2017; Leisman et al. 2017; Greco et al. 2018; Román et al. 2019; Barbosa et al. 2020; Zaritsky et al. 2022). A transition in the overall properties of cluster versus field UDGs is seen, with UDGs in clusters being mostly red, spheroidal, and quenched, while those in the field tend to be bluer, more irregular, and gas-rich (e.g. van der Burg, Muzzin & Hoekstra 2016; Leisman et al. 2017; Román & Trujillo 2017). Moreover, they have been found to exist at least up to redshift $z \sim 0.8$ (e.g. Janssens et al. 2017; 2019; Lee et al. 2020; Carleton et al. 2023; Ikeda et al. 2023).

In order to better understand how UDGs formed, a large effort has been undertaken in the past few years to characterize the main properties of these LSB galaxies. Their dark matter content has been one of the most revealing properties. Some UDGs have inferred dark matter haloes similar to those found in classical dwarf galaxies at the same stellar mass [$\log(M_{\text{halo}}/M_\odot) \lesssim 10.5$; Beasley & Trujillo 2016; Amorisco 2018; Iodice et al. 2020; Gannon et al. 2021], while other UDGs show evidence of more massive dark matter haloes [$\log(M_{\text{halo}}/M_\odot) \gtrsim 10.5$; Beasley et al. 2016; Toloba et al. 2018; van

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Dokkum et al. 2019; Forbes et al. 2021; Gannon et al. 2022]. In some rare cases, they have been found to host no dark matter at all, a very disputed result that still attracts a lot of attention (e.g. van Dokkum et al. 2018; 2019; Wasserman et al. 2018; Trujillo et al. 2019; Danieli et al. 2020; Montes et al. 2020; 2021; Shen et al. 2021).

Some UDGs have been found to host rich systems of globular clusters (GCs; e.g. Beasley et al. 2016; van Dokkum et al. 2017; 2018; Amorisco et al. 2018; Lim et al. 2018; Forbes et al. 2019; 2020a; Müller et al. 2021; Saifollahi et al. 2021), whereas other UDGs can be extremely GC-poor (e.g. Forbes et al. 2020a; Saifollahi et al. 2022; Gannon et al. 2023; Jones et al. 2023). There is a tight correlation between the dark matter halo mass and the total number of GCs in normal galaxies (e.g. Spitler & Forbes 2009; Harris, Harris & Alessi 2013; Burkert & Forbes 2020), which UDGs also seem to follow (e.g. Gannon et al. 2022; Zaritsky et al. 2023). GC-poor UDGs are thus likely to reside in dark matter haloes similar to those found for classical dwarf galaxies, while GC-rich ones may reside in more massive dark matter haloes (e.g. Sifón et al. 2018; Gannon et al. 2023).

Evidence of the latter is seen in the kinematic studies of several GC-rich UDGs. Measuring their dynamical masses, they were found to be consistent with halo-profiles for massive dark matter haloes $\log(M_{\text{halo}}/M_{\odot}) \gtrsim 10.5$ (e.g. Beasley et al. 2016; van Dokkum et al. 2016; Toloba et al. 2018; Forbes et al. 2021, Gannon et al. 2020; 2022). However, much is yet to be unveiled for the GC-poor UDGs and their dark matter haloes. Altogether, from these clearly distinctive properties, two main different types of UDGs have been proposed to co-exist. Several mechanisms have been suggested to create each class of UDG, with different sets of cosmological simulations trying to reproduce them.

Primarily, many propose that UDGs are classical dwarf galaxies that are ‘puffed up’ by some mechanism into the large size observed, which we will refer to as the *puffy dwarf formation scenario*. This can be attributed to either internal feedback (e.g. Di Cintio et al. 2017; Chan et al. 2018), high halo spins (e.g. Amorisco & Loeb 2016; Rong et al. 2017; Liao et al. 2019), or external effects such as quenching, tidal forces, or ram pressure stripping (e.g. Safarzadeh & Scannapieco 2017; Ogiya 2018; Carleton et al. 2019; Sales et al. 2020; Tremmel et al. 2020; Benavides et al. 2021; Jones et al. 2021; Junais et al. 2022). Combinations of these effects may also occur, although they do not need to happen simultaneously (e.g. Jiang et al. 2019; Liao et al. 2019; Martin et al. 2019; Trujillo-Gomez, Kruijssen & Reina-Campos 2022; Watkins et al. 2023).

Another school of thought describes UDGs as a remnant dark matter halo from the early Universe. An ancient, dark matter-dominated, proto-galaxy could have been quenched very early on, preventing its subsequent expected evolution into a regular galaxy. This prevented the galaxy from acquiring the stellar mass expected given its original dark matter halo, explaining the massive dark matter halo it still resides in (e.g. Yozin & Bekki 2015; Peng & Lim 2016; Rong et al. 2017; Danieli et al. 2022). This scenario was originally named ‘failed Milky Way-galaxies’ (van Dokkum et al. 2015), although it has since been demonstrated that these galaxies cannot reside in dark matter haloes quite as massive as the Milky Way (see e.g. Sifón et al. 2018). Hence, we will use the simplified term *failed-galaxy scenario* hereafter. These UDGs, while still being LSB dwarfs in terms of their stellar masses, are expected to reside in massive dark matter haloes (e.g. Harris, Harris & Alessi 2013; Burkert & Forbes 2020), thus are expected to be GC-rich systems. However, they have been hardly reproduced by any cosmological simulations yet.

One of the best ways to investigate this presumed dichotomy in the origin of UDGs comes from the analysis of their stellar population properties, as precise and distinct predictions for each type are expected. Some works tackled this by using SED fitting to derive some stellar population properties of a large number of UDGs at once (e.g. Pandya et al. 2018; Barbosa et al. 2020; Buzzo et al. 2022). However, SED fitting is not as well constrained as spectroscopy, and only a few spectroscopic attempts have been carried to date for UDGs. This has been done for either limited samples of UDGs [mostly in clusters, e.g. Gu et al. 2018; Ferré-Mateu et al. 2018b (FM+18 hereafter); Ruiz-Lara et al. 2018a; Chilingarian et al. 2019], or for some individual cases, including UDGs in less dense environments (e.g. Fensch et al. 2019; Martín-Navarro et al. 2019; Müller et al. 2021; Villaume et al. 2022). Sometimes they have been stacked to enhance the quality of the spectra, as done for the star-forming isolated UDGs from SDSS (Rong et al. 2020b).

Under the puffy dwarf-like scenario, UDGs are expected to have mostly extended star formation histories (SFHs) and thus younger ages ($\sim 6\text{--}9$ Gyr), with metallicities similar to classical dwarfs ($[M/H] \sim -0.8$ dex; e.g. FM+18; Ruiz-Lara et al. 2018a; Chilingarian et al. 2019; Martín-Navarro et al. 2019; Rong et al. 2020a). Failed-galaxy UDGs would instead be expected to have stellar populations more similar to GCs, with old ages ($\sim 10\text{--}12$ Gyr) as they were quenched very early on. In such case, their metallicities are also expected to be much lower than the puffy dwarf-like UDGs ($[M/H] \sim -1.5$ dex; e.g. Naidu et al. 2022).

One intriguing property that has not been tackled much yet for UDGs, and that is not possible to obtain from SED fitting, is the $[\alpha/Fe]$ abundance ratio (or $[Mg/Fe]$, depending on the work and how it is measured). From the few estimates available, UDGs seem to present elevated ratios ($\gtrsim 0.3$ dex; e.g. FM+18; Ruiz-Lara et al. 2018a; Rong et al. 2020b). In the case of the isolated UDG DGSAT I, it exhibits the most extreme $[Mg/Fe]$ ever found in any type of galaxy (Martín-Navarro et al. 2019). In a canonical view, the α -enhancement has been suggested to be a cosmic-clock, as each element is formed at different time-scales (e.g. Matteucci 1994; Thomas et al. 2005; de La Rosa et al. 2011; McDermid et al. 2015). Galaxies forming on very short time-scales present elevated $[Mg/Fe]$, as the first yields to be produced by core-collapse are those of α -elements such as magnesium. As time evolves, galaxies with more extended SFHs decrease their $[Mg/Fe]$ values (typically down to solar abundance ratios) as the interstellar medium is polluted by the ejection of iron from Type Ia supernovae. Therefore galaxies with elevated $[Mg/Fe]$ but extended SFHs, as found for many UDGs, is a surprising result under the current galaxy formation paradigm.

The challenge in all these spectroscopic studies is that acquiring good quality spectra for stellar population analysis requires a large time investment with 10 m class telescopes. A balance between the quality of the spectra and the number of UDGs observed in the sample must therefore be found. Moreover, most of the studied samples are clearly biased towards cluster environments, leaving field or group UDGs mostly unexplored. To work on this, we have carried several observational campaigns to obtain moderate quality spectra ($S/N \sim 15\text{--}20$) for a large sample of UDGs in different environments, although with a heavy preference towards Perseus UDGs. As a result, adding these new UDGs (11) to previously published UDGs (14), we nearly double the number of quiescent UDGs with spectroscopic measurements. We consider UDGs to be quiescent when they have no signs of emission lines and do not exhibit ongoing star formation. We note that this does not exclude UDGs having experienced recent star formation (e.g. in the last few Gyr).

We have thus created the largest spectroscopic sample of quiescent UDGs with stellar population properties. For most of the new UDGs, a dynamical study of their masses, velocity dispersion, and dark matter content has been already published (Forbes et al. 2021; Gannon et al. 2020, 2021, 2022, 2023).

In this paper, we study the stellar populations of this large sample of spectroscopic UDGs, to further understand the formation mechanisms that regulate the life of these intriguing galaxies. The paper is organized as follows. Section 2 presents the data used, for both our observations and the compilation from the literature. Section 3 describes the methodology we use to derive the main stellar population properties of our sample, with particular attention to the different caveats of the techniques. In Section 4 we discuss our results in the context of their environment UDGs reside in (Section 4.2) and different relations related to other families of galaxies (Section 4.3). We additionally compare to predictions for UDG stellar properties coming from simulations (Section 4.4). Finally, Section 5 presents a summary of our results.

2 DATA

2.1 This work UDG sample

In this work, we present a new stellar population analysis for a sample of quiescent UDGs with data acquired using the Keck Cosmic Web Imager (KCWI), an integral field unit in the Keck II Telescope. Details about each observation can be found in the original papers presenting the dynamical and kinematic analysis of those UDGs: the Virgo cluster UDG VCC 1287 (Gannon et al. 2020), NGC 5846-UDG1 in the NGC 5846 group (Forbes et al. 2021), UDG1137+16 from the UGC 6594 group (Gannon et al. 2021), several UDGs in the Perseus cluster (i.e. PUDG-R15, PUDG-S74, PUDG-R84, and PUDG-R24; Gannon et al. 2022), and Yagi358 in the Coma cluster (Gannon et al. 2023).

We add to this sample three UDGs for which we have acquired new data. These are the Coma cluster UDGs DFX1 and DF07, along with the Perseus cluster UDG PUDG-R27. DFX1 and DF07 were observed on the nights of 2021 April 16 and 17 (Program U105, PI: Brodie) using KCWI with the medium slicer, BL grating and a 4550 Å central wavelength. PUDG-R27 ($R.A. = 49.93116$, $Dec. = 41.71326$ in J2000) was observed on the nights of 2021 October 5 (Program W024; PI: Forbes), and 2022 January 5 and 29 (Program N195, PI: Romanowsky) using KCWI with the large slicer, BL grating, and a 4550 Å central wavelength. Note that during the night of 2022 January 29, KCWI was slightly out of focus, which degraded the instrumental resolution by ~ 15 per cent, and affected ~ 20 per cent of our data for PUDG-R27. Data were reduced using the standard KCWI pipeline (Morrissey et al. 2018) with additional cropping and flat fielding corrections as described in Gannon et al. (2020). For all targets, KCWI was targeted slightly off-centre to allow on-chip sky subtraction. Final exposure times were 25 200 s for DFX1, 19 200 s for DF07, and 26 400 s for PUDG-R27.

2.2 Literature sample

We build a literature sample using moderate S/N spectroscopic results published for the stellar population properties of Coma cluster UDGs (from FM+18; Gu et al. 2018; Ruiz-Lara et al. 2018a; Chilingarian et al. 2019; and Villaume et al. 2022). We include the group UDG NGC 1052-DF2 (Fensch et al. 2019) and the isolated UDG DGSAT I (Martín-Navarro et al. 2019).

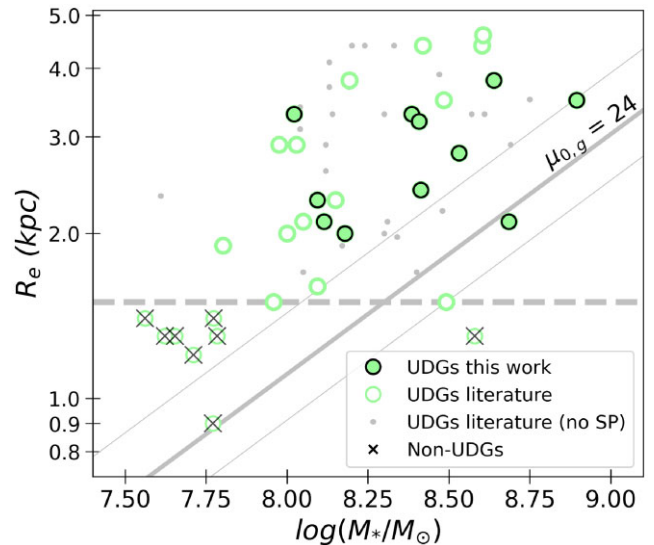


Figure 1. Mass versus size diagram to separate between true UDGs and non-UDGs. The limit of $R_e > 1.5$ kpc is shown by a dashed horizontal line, while the diagonal solid lines represent constant surface brightness limit of $\mu_{0,g} = 24$ mag arcsec $^{-2}$ with different assumptions of M/L : thickest line is for a M/L_g of ~ 1.6 (mean value of our sample, see Section 3.2), while thinner ones represent our minimum and maximum values of M/L obtained. The figure presents the UDGs in this work (filled dots) and those from the literature sample (open dots; data described in Section 2). Other UDGs in the literature that have no spectroscopic stellar population information are shown as grey dots (Buzzo et al. 2022; Gannon et al. 2022). We highlight with an X the objects that will not be considered UDGs for this work.

Various definitions have been used to define UDGs in the literature. These are based mostly on the size and surface brightness of the objects (but see Van Nest et al. 2022 for a variety of definitions). We thus need to ensure that all the considered UDGs in this work, including those from the literature, are truly UDGs. In Fig. 1 we select as bona-fide UDGs with $R_e > 1.5$ kpc (horizontal line) and $\mu_{0,g} > 24$ mag arcsec $^{-2}$ (diagonal lines, for different mass-to-light ratio, M/L , assumptions). All the UDGs from this work fulfill these criteria, with PUDG-R27 included because it lies within the surface brightness line with the highest M/L (which is in fact its own value; see Section 3.2). UDGs that lie below but close to the limits should not be considered true UDGs, although some might become a UDG with time (e.g. through tidal heating and surface brightness dimming; Grishin et al. 2021).

Adding the sample of 14 literature UDGs to the 11 UDGs from this work we have at hand the largest spectroscopic sample of UDGs to date. Table 1 summarizes the name, environment, stellar mass, half-light radius, and GC-richness of the UDGs discussed in this work. For this work, we do not require a stringent estimate on the exact number of GCs, so we adopt a boundary between ‘poor’ and ‘rich’ at 20 GCs, similar to previous works (e.g. Gannon et al. 2022). UDGs above this limit are expected to reside in a relatively massive dark matter halo under the assumption that they obey the GC number – halo mass relationship of Burkert & Forbes (2020).

The GC richness estimates are collected from a variety of sources in the literature (e.g. Forbes et al. 2020c; Danieli et al. 2022; Gannon et al. 2022). There are 10 UDGs in our sample without this information, all from the Coma cluster. For some of these we have carried out new estimates of their GC richness, as described in Appendix A, following a similar visual approach to Gannon et al. (2022). These

Table 1. Main properties of the UDGs in this work: their names (column 1); the environment they reside in (column 2): cluster, group or field, specifying to which cluster or group they belong to; their recession velocity and stellar masses (column 3 and 4); their effective radius (column 5); and their GC classification: poor or rich (column 6) according to the number of GCs. Those within parenthesis should be taken with care (see Appendix A). Column 7 cites the source of both the structural properties (velocities, stellar masses, and sizes) and stellar populations, respectively, for the sample.

Galaxy	Env.	V_r km s ⁻¹	M_* ($\times 10^8 M_\odot$)	R_e (kpc)	GC-richness classification	Data
This work						
VCC 1287	Cluster (Virgo)	1116	2.43	3.3	rich	Gannon et al. (2020), this work
NGC 5846-UDG1	Group (NGC 5846)	1014	1.30	2.1	rich	Forbes et al. (2021), this work
UDG1137+16	Group (Leo)	2167	1.05	3.3	unknown	Gannon et al. (2021), this work
PUDG-R15	Cluster (Perseus)	4762	2.59	2.4	poor	Gannon et al. (2022), this work
PUDG-S74	Cluster (Perseus)	6215	7.85	3.5	rich	Gannon et al. (2022), this work
PUDG-R84	Cluster (Perseus)	4039	2.20	2.0	rich	Gannon et al. (2022), this work
PUDG-R24	Cluster* (Perseus)	7787	3.91	3.2	poor	Gannon et al. (2022), this work
PUDG-R27	Cluster (Perseus)	4376	4.84	2.1	rich	this work
Yagi358 (GMP 3651)	Cluster (Coma)	7969	1.24	2.3	rich	Gannon et al. (2023), this work
DFX1	Cluster (Coma)	8107	3.40	2.8	rich	this work
DF07	Cluster (Coma)	6600	4.35	3.8	rich	this work
Literature						
Yagi093 (DF26, GMP 2748)	Cluster (Coma)	6611	3.05	3.5	rich	Alabi et al. (2018), FM+18
Yagi098	Cluster (Coma)	5980	1.07	2.9	rich	Alabi et al. (2018), FM+18
Yagi275 (GMP 3418)	Cluster* (Coma)	4847	0.94	2.9	unknown	Alabi et al. (2018), FM+18
Yagi276 (DF28)	Cluster (Coma)	7343	1.41	2.3	unknown	Alabi et al. (2018), FM+18
Yagi392	Cluster (Coma)	7748	0.91	1.5	poor	Alabi et al. (2018), FM+18
Yagi418	Cluster (Coma)	8335	1.24	1.6	(rich)	Alabi et al. (2018), FM+18
Yagi090	Cluster* (Coma)	9420	1.00	2.0	poor	Ruiz-Lara et al. (2018a)
OGS1	Cluster (Coma)	6367	3.10	1.5	(poor)	Ruiz-Lara et al. (2018a)
DF17	Cluster (Coma)	8311	2.63	4.4	rich	Gu et al. (2018)
DF44	Cluster* (Coma)	6324	4.03	4.6	rich	Villaume et al. (2022), Webb et al. (2022)
J130026.26+272735.2 (GMP 2673)	Cluster (Coma)	6939	1.56	3.8	(poor)	Chilingarian et al. (2019)
J130038.63+272835.3 (GMP 2552)	Cluster (Coma)	7937	0.64	1.9	(poor)	Chilingarian et al. (2019)
NGC 1052-DF2	Group (NGC 1052)	–	2.50	–	poor [†]	Fensch et al. (2019)
DGSAT I	Field	–	4.00	4.4	poor [†]	Martín-Navarro et al. (2019)

Notes. * Despite being located in clusters their local environment is considered low-density (see Section 4.2).

[†] These objects would be considered GC-rich if the defining parameter were GC system mass instead of GC numbers, e.g. Janssens et al. (2022).

estimates are primarily based on *HST* imaging, and in some cases on deep ground-based imaging from Subaru/Suprime-Cam.

Fig. 2 presents the spectra for the UDGs whose stellar populations are newly presented in this work (black). Different KCWI instrumental configurations were used. The medium slicer with a BH3 grating provides a shorter baseline ($\sim 4800\text{--}5200 \text{ \AA}$) but higher spectral resolution (0.5 \AA FWHM), crucial to derive proper velocity dispersion measurements. The BL grating provides a longer baseline ($\sim 3800\text{--}5400 \text{ \AA}$) but with lower spectral resolution (5.1 \AA FWHM), adequate for stellar population studies. PUDG-R84 was observed in both configurations as a consistency check. However, for this UDG we will use the results from the BH configuration, as it has a much higher S/N than the BL one. One caveat to have in mind for the UDGs observed with the BH grating is that a short wavelength range can introduce systematics on the derived properties. Shorter baselines (in particular when not including the age indicator H_β) tend to give younger ages (see e.g. Forbes et al. 2022). None the less, as later discussed in Appendix B2, this is not affecting our results.

3 STELLAR POPULATION ANALYSIS

Before proceeding, we need to highlight a caveat in the data. Some of the spectra have been reduced to have a pedestal removed from their continuum as a result of the sky subtraction process. This is due to an inability of the PCA sky subtraction routine described

in Gannon et al. (2020) to disentangle the absolute level of the continuum coming from the galaxy and the sky emission. This results in those spectra having negative counts at some wavelengths (see Fig. 2). While this does not have any impact on the full spectral fitting procedure detailed below, it means that we cannot perform a line index measurement for VCC 1287, NGC 5846-UDG1, UDG1137+16, PUDG-R24, and PUDG-S74.

We therefore use both the full spectral fitting code `pPXF` (Cappellari 2017) and perform a classical absorption line index measurement (when possible) to our sample of UDGs to obtain mean mass-weighted ages, metallicities, and α -enhancement ratios. We also derive the SFHs, from which one can calculate crucial time-scales such as the quenching time-scale. We employ the MILES single stellar population (SSP) models (Vazdekis et al. 2015) assuming a Kroupa initial mass function (Kroupa 2001). These models range in age from 0.03 to 14 Gyr and metallicities of $[M/H] = -2.42$ to $+0.40$ dex. We use the scaled solar base models to obtain the stellar population parameters such as ages, metallicities, and SFHs.

First, the optimal degree for the polynomials to be used in the fitting is found for each UDG, which is typically of the order of 5 for both the additive and multiplicative polynomials (see for example D’Ago et al. 2023 for a detailed discussion on how to determine these values). `pPXF` is then run to obtain a non-regularized solution (it will be the closest to a line index approach, a discrete SFH), and also a regularized one. The latter provides a more realistic

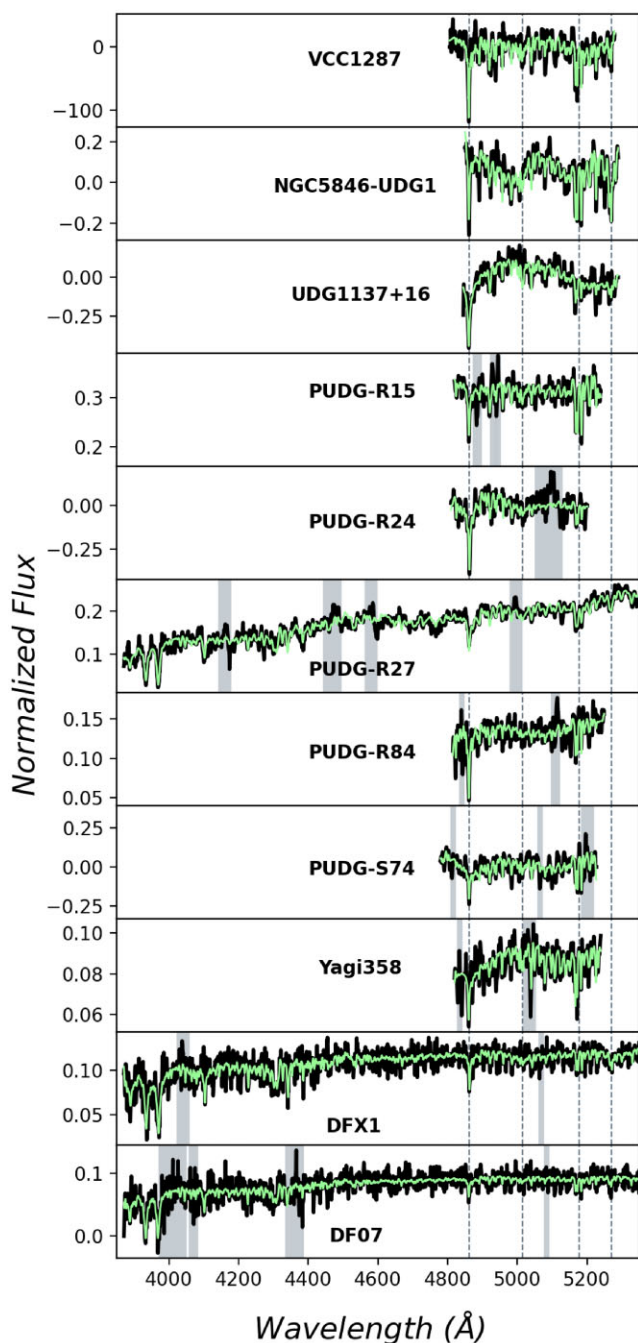


Figure 2. Spectra used in this work to derive new stellar populations for our sample of UDGs (black). The different instrumental configurations are evident: shorter wavelength and higher resolution for the BH grating, and lower resolution but larger baseline for the BL grating). The best fit from the full spectral-fitting method used in Section 3 to derive the stellar populations is shown in green. Regions masked during the fitting process are plotted as grey vertical bands. The main line indices used in the study are shown with dashed lines (from left to right): H_{β} , Fe5015, Mg_b and Fe5270.

SFH while maintaining an accurate fit given its simple Bayesian approach (see e.g. McDermid et al. 2015; Cappellari 2017; and Westfall et al. 2019 for additional information on this process). We have checked that the results are independent of the regularization choice, showing similar SFHs and stellar populations properties. We therefore use the average between the regularized and un-regularized solution for the ages, metallicities, and time-scales presented here.

All uncertainties associated with the stellar population values are calculated by performing Monte Carlo simulations of this fitting procedure for each UDG.

3.1 α -enhancement ratios

The α -enhancement is a challenging measurement that has a high dependence on the methodology used. For example, if implemented in a full spectral fitting routine as another free parameter (together with the age and metallicity), the values obtained will only span the values of the SSP models used. In the case of the models used here, those values will be between $[\alpha/Fe] = 0.0$ and $+0.4$ dex. Instead, if calculated using line index fitting, there are no limitations besides how far one can extrapolate the values outside the model grids. With this last method one obtains $[Mg/Fe]$, as those are the line indices typically used. One can convert from one to the other ($[\alpha/Fe] = 0.02 + 0.56 \times [Mg/Fe]$, if using the BaSTI Base models; Vazdekis et al. 2015). In addition, the $[\alpha/Fe]$ obtained from the full spectral fitting approach can be underestimated for galaxies with low metallicities. This is because the SSP models are based on solar-scaled stars, which in reality are Mg enhanced at the low metallicity regime ($\lesssim -1$ dex; Vazdekis et al. 2015), i.e. similar to the regime covered by UDGs.

Taking all this into consideration, it is difficult to unify the ways different works measure the ratio between α -elements and iron. Appendix B1 contains a section fully dedicated to the different ways of calculating it and their systematics. Unfortunately we cannot use this more robust line index measurement for 5 of our UDGs due to the pedestal issue discussed in Section 3. We will therefore use, when available, the $[Mg/Fe]$ value from the line index measurement. Otherwise, we will use the one from pPXF, cautioning that these may be lower limits. Moreover, we note that the abundance ratios obtained from the line-index measurement are typically above the averaged value and higher than the SSP limit marked by the full spectral fitting approach. This highlights the importance of being careful about the way one computes this value. The results from line index measurements are marked with a † in Table 2, and are shown as an open symbol in the top panel of Fig. 3.

We find that the majority of our UDGs present elevated $[Mg/Fe]$ abundance patterns, with a mean value and dispersion of $\langle [Mg/Fe] \rangle = +0.51 \pm 0.32$ dex, slightly higher than what was found in FM+18. Converting this $[Mg/Fe]$ value into $[\alpha/Fe]$ would result in an enhanced value of $[\alpha/Fe] \sim 0.32$ dex. PUDG-R24 is the only UDG that shows a scaled solar pattern.

3.2 Mean ages and metallicities

We obtain several parameters from the pPXF fitting such as the mean ages (t_M , in lookback time), mean metallicities ($[M/H]$), the stellar mass-to-light ratios (M/L in the r -band) and the predicted $g - i$ colour. The resulting stellar populations properties are shown in Fig. 3. For each, we present the mass-weighted results for both the regularized (black symbols) and non-regularized (grey symbols) solutions. Average values between regularized and non-regularized results were obtained and these are the ones that will be used for the remainder of this paper. These are summarized in Table 2.

The mean mass-weighted age of our sample of UDGs is $\langle t_M \rangle = 8.3 \pm 3.3$ Gyr. Three UDGs have very old ($t_M > 10$ Gyr) mass-weighted ages, six UDGs have old ages (between 7.5 and 10 Gyr), and two UDGs are young (< 4.5 Gyr). Due to these large age differences, the M/L values of their stellar populations vary from roughly 0.3 to 2. We remind the reader of the caveat that shorter wavelength ranges may deliver younger ages, discussed in Appendix B2.

Table 2. Stellar population properties of the UDGs in this work (top block) and literature (bottom). For this sample, the values correspond to a median between the non-regularized and regularized solution for the mean mass-weighted age (column 2); t_{50} and t_{90} (lookback times when the galaxy formed 50 and 90 per cent of its stellar mass; columns 3 and 4, respectively). Those UDGs for which this value is not stated in the original works, but the SFH is shown, have been digitized to obtain these values. t_{inf} is an estimate of the expected lookback time of cluster infall assuming quenching is due to cluster/group infall ($t_{\text{inf}} = t_{90} - 1.5$ Gyr, column 8). Δt_{50} and Δt_{90} are the time-scales it took to build 50 and 90 per cent of stellar mass since the Big Bang (columns 6 and 7, respectively). Column 8 shows the mean mass-weighted metallicity and column 9 the [Mg/Fe], with † being the values obtained from the line index approach (see Appendix B1). We note that $[\alpha/\text{Fe}]$ values from the literature have been transformed into [Mg/Fe] using the corresponding conversion from Vazdekis et al. (2015).

Galaxy	Age (t_M) (Gyr)	t_{50} (Gyr)	t_{90} (Gyr)	t_{inf} (Gyr)	Δt_{50} (Gyr)	Δt_{90} (Gyr)	[M/H] (dex)	[Mg/Fe] (dex)
VCC 1287	9.09 ± 1.07	9.2 ± 0.2	8.5 ± 1.1	10.0 ± 1.1	4.7 ± 0.2	5.9 ± 1.1	-1.06 ± 0.34	0.56 ± 0.11
NGC 5846-UDG1	8.20 ± 3.05	8.3 ± 0.1	7.5 ± 1.0	–	5.5 ± 0.1	6.8 ± 1.0	-1.15 ± 0.25	0.54 ± 0.18
UDG1137+16	2.13 ± 1.58	1.0 ± 0.3	0.1 ± 0.1	–	12.9 ± 0.3	13.7 ± 0.1	-1.52 ± 0.40	0.39 ± 0.10
PUDG-R15	11.32 ± 2.52	13.4 ± 0.8	9.5 ± 0.1	11.0 ± 0.1	0.8 ± 0.8	4.5 ± 0.1	-0.93 ± 0.32	$0.44 \pm 0.20^\dagger$
PUDG-S74	8.44 ± 2.26	8.2 ± 3.3	9.5 ± 0.1	11.0 ± 0.1	4.2 ± 3.3	7.3 ± 0.1	-0.40 ± 0.22	0.32 ± 0.11
PUDG-R84	8.99 ± 3.20	8.4 ± 0.5	6.8 ± 0.4	8.5 ± 0.4	5.7 ± 0.5	7.2 ± 0.4	-1.48 ± 0.46	0.22 ± 0.30
PUDG-R24	2.58 ± 1.80	3.1 ± 0.1	0.8 ± 0.1	2.5 ± 0.1	10.5 ± 0.1	12.9 ± 0.1	-1.00 ± 0.26	0.03 ± 0.08
PUDG-R27	11.04 ± 1.11	10.3 ± 0.6	8.8 ± 0.7	10.0 ± 0.7	3.1 ± 0.6	5.3 ± 0.7	-0.61 ± 0.14	$0.31 \pm 0.20^\dagger$
Yagi358	9.81 ± 2.46	13.2 ± 1.3	6.1 ± 0.1	7.5 ± 0.1	1.3 ± 1.3	7.4 ± 0.1	-1.56 ± 0.60	$1.40 \pm 0.20^\dagger$
DFX1	8.84 ± 1.13	9.1 ± 0.8	8.1 ± 1.4	9.5 ± 1.4	5.1 ± 0.8	6.5 ± 1.4	-1.08 ± 0.21	$0.70 \pm 0.40^\dagger$
DF07	11.18 ± 1.27	10.5 ± 1.0	9.1 ± 0.7	10.5 ± 0.7	2.8 ± 1.0	5.0 ± 0.7	-0.78 ± 0.18	$0.68 \pm 0.40^\dagger$
Yagi093	7.88 ± 1.76	11.3	4.1 ± 0.1	5.5 ± 0.1	3.2	10.0 ± 0.1	-0.56 ± 0.18	0.38 ± 0.17
Yagi098	6.72 ± 2.16	11.3	2.1 ± 2.9	3.5 ± 2.9	2.8	10.5 ± 2.9	-0.72 ± 0.20	–
Yagi275	4.63 ± 1.50	5.1	2.2 ± 0.1	4.0 ± 0.1	4.7	11.8 ± 0.1	-0.37 ± 0.17	-0.25 ± 0.38
Yagi276	4.24 ± 2.32	9.0	2.0 ± 0.1	3.5 ± 0.1	4.9	12.0 ± 0.1	-0.38 ± 0.79	–
Yagi392	7.36 ± 2.06	10.5	2.8 ± 2.7	4.5 ± 2.7	4.3	8.2 ± 2.7	-0.58 ± 0.73	–
Yagi418	7.87 ± 2.02	12.3	3.2 ± 2.5	4.5 ± 2.5	4.0	9.9 ± 2.5	-1.10 ± 0.85	0.17 ± 0.31
Yagi090	5.75 ± 1.30	9.4	5.0 ± 0.1	6.5 ± 0.1	4.4	9.6 ± 0.1	-1.35 ± 0.05	0.40
OGS1	8.50 ± 1.20	13.8	10.4 ± 0.7	11.5 ± 0.7	0.1	3.3 ± 0.7	-0.53 ± 0.06	0.35
DF17	9.11 ± 2.00	–	–	–	–	–	-0.83 ± 0.50	–
DF44	10.23 ± 1.50	0.34	3.0 ± 4.7	4.5 ± 4.7	2.8	6.3 ± 4.7	-1.33 ± 0.05	-0.10 ± 0.10
J130026.26+272735.2	1.5 ± 0.10	–	–	–	–	–	-1.04 ± 0.11	–
J130038.63+272835.3	1.7 ± 0.10	–	–	–	–	–	-0.74 ± 0.08	–
NGC 1052-DF2	8.90 ± 1.50	–	–	–	–	–	-1.07 ± 0.11	0.10 ± 0.05
DGSAT I	8.10 ± 0.40	–	–	–	5.6	13.5	-1.80 ± 0.40	1.50 ± 0.50

In this work, this systematic bias may have had an effect on the two young UDGs. However, PUDG-R24 has been identified as a very recent infall into the Perseus cluster and UDG1137+16 presents tidal features, which can in both cases explain the young stellar ages.

The UDGs in this work all have sub-solar metallicities, with a mean value of $\langle [M/H] \rangle = -1.03 \pm 0.37$ dex, similar to previous works (e.g. FM+18; Gu et al. 2018; Ruiz-Lara et al. 2018a). The mean colour of the UDGs, $\langle g - i \rangle = 0.8 \pm 0.2$, is compatible with the one found in other large samples of UDGs (e.g. Román & Trujillo 2017; Junais et al. 2022).

We note that several of the UDGs in the sample already have stellar populations results available. VCC 1287, PUDG-R24, Yagi358, DFX1, and DF07 have published SED fitting stellar populations (e.g. Pandya et al. 2018; Buzzo et al. 2022). As mentioned above, here we only consider spectroscopic results, but we refer the reader to Appendix B3 for a comparison to these SED fitting results. While stellar masses and ages do not differ significantly from the SED solution, stellar metallicities from SED fitting are systematically lower, on the order of ~ 0.25 dex. NGC 5846-UDG1 was previously studied in Müller et al. (2020) using MUSE spectroscopy with a slightly lower S/N than our new data (see also Heesters et al. 2023), and DF07 in Gu et al. (2018) from MaNGA data. These works reported a $[\text{Fe}/\text{H}] = -1.33 \pm 0.19$ dex and an age of 11.2 ± 1.5 for NGC 5846-UDG1 and $[\text{Fe}/\text{H}] = -1.03 \pm 0.31$ dex and an age of 9 ± 1.5 Gyr for DF07. These are compatible with our measurements within the uncertainties. Therefore, we will use our own values for the remaining of the paper.

3.3 Star formation histories and characteristic time-scales

From the recovered SFHs, different characteristic lookback times are calculated: when the galaxy formed half of its stellar mass (t_{50}), and when it built 90 per cent of it (t_{90}). The latter is also considered to be a good estimate for the quenching time of the galaxy (e.g. Weisz et al. 2014; 2019; Collins & Read 2022). These lookback times are translated into time-scales (i.e. the number of years it took to build that given amount of stellar mass). This way, we calculate $\Delta t_{50} = t_{\text{BB}} - t_{50}$ and $\Delta t_{90} = t_{\text{BB}} - t_{90}$, where t_{BB} is the time since the Big Bang, 13.8 Gyr. Therefore we are using here as a quenching time-scale the commonly used time-scale *since the Big Bang*, not since the onset of star formation.

Under the assumption that quenching is environmentally driven, it is thought that galaxies quench roughly 1–2 Gyr after falling into the cluster or group environments (e.g. Muzzin et al. 2008; Fillingham et al. 2015). We thus estimate a time of infall as 1.5 Gyr before quenching, i.e. $t_{\text{inf}} \sim t_{90} - 1.5$ Gyr (see also FM+18). Table 2 also quotes these relevant time-scales inferred from the SFHs.

4 DISCUSSION

4.1 SFHs and their relation with environment

The current galaxy formation paradigm predicts that the mass ratio between α elements and iron is a good proxy for the formation time-scales of galaxies (e.g. Matteucci 1994; Thomas et al. 2005; de La

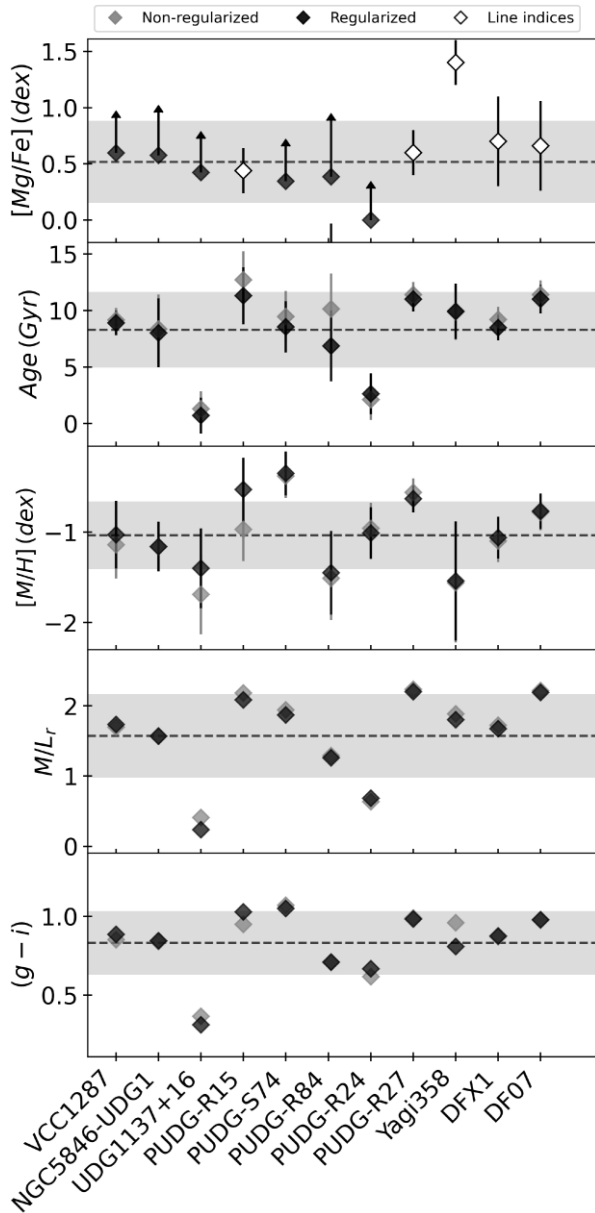


Figure 3. Stellar populations for the UDGs in this work: $[\text{Mg}/\text{Fe}]$, mass-weighted age, mass-weighted metallicity, the derived M/L in the r -band, and the $g-i$ colour (from top to bottom, respectively). For each property we present the individual measurements for each UDG (diamonds) and the average value for all UDGs (dashed line) with the $1-\sigma$ deviations (grey band). Grey symbols correspond to the non-regularized solution while black ones correspond to the regularized one from the full spectral fitting with `pPXF`. White symbols in the top panel correspond to line index measurements (see Appendix B1 for a discussion on the different ways to obtain the $[\text{Mg}/\text{Fe}]$).

Rosa et al. 2011; McDermid et al. 2015; Romero-Gómez et al. 2023). Typically, galaxies with younger ages (from more extended SFHs) tend to have near solar values, while those with very fast formation tend to have elevated abundance ratios. However, one of the most puzzling results to date for UDGs was that overall they show elevated α -enhancement, despite presenting, in many cases, extended SFHs and intermediate ages (e.g. FM+18; Martín-Navarro et al. 2019; Rong et al. 2020b).

To further investigate the relation between the SFHs and the α -enhancement, Fig. 4 presents the regularized SFHs as the cumulative

mass fraction over cosmic time of the 11 UDGs in this work and other Coma UDGs from the literature for which we have access to (FM+18 and Ruiz-Lara et al. 2018a). They are colour-coded by their level of α -enhancement: sub-solar (<0 dex), scaled-solar (0 to $+0.2$ dex), enhanced ($+0.2$ dex to $+0.4$ dex), and extremely enhanced (>0.4 dex). This has been done to unify measurements from different techniques and works (e.g. the use of $[\alpha/\text{Fe}]$ or $[\text{Mg}/\text{Fe}]$), as we are here interested in the level of enhancement rather than the absolute value of it.

Fig. 4 shows distinctly different types of SFHs. While most of the UDGs are quenched on relatively short time-scales (around 8 Gyr ago), none of them was quenched during the epoch of reionization, marked with a grey band in the figure (Fan, Carilli & Keating 2006). A clear trend with the α -enhancement is seen. Some UDGs started forming stars very early on and built their entire mass relatively fast (*early-and-fast*). Hence they show very short quenching time-scales ($\Delta t_{90} \sim 4-5$ Gyr), thus expected to be early cluster infallers, and have the oldest stellar ages ($t_M \sim 10-13$ Gyr). We note that galaxies with similar SFHs are found in the Local Group (e.g. such as And XIX; Collins & Read 2022). Moreover, these *early-and-fast* UDGs also show the highest α -enhancement values (> 0.4 dex).

Similarly high α -enhancements are seen for several UDGs that present a fast build-up but that their onset of star formation occurred later in cosmic time (~ 4 Gyr after the Big Bang). This means that while the quenching time-scales of these *late-and-fast* UDGs are as short as to the *early-and-fast* type, they will present intermediate stellar ages ($t_M \sim 6-9$ Gyr).

Some other UDGs also started forming very early on, however they did not quench for a much longer time (*early-and-slow*), sometimes even later than the *late-and-fast* UDGs. These UDGs, also mainly located in clusters, show intermediate ages due to their more extended SFHs, and their α -enhancement, despite being lower than in the previous two types, are still relatively high ($\sim 0.2-0.4$ dex).

Finally, there are three UDGs that show very late, extended SFHs, quenching only 1–3 Gyr ago (*late-and-slow*). They present extremely young stellar ages, suggestive of a late cluster infall. From these *late-and-slow* UDGs, the two located in clusters are the ones with the lowest α -enhancements ($\lesssim 0.2$ dex; PUDG-R24 and Yagi275). Such extremely late SFHs have also been reported for other classical dwarf galaxies (see e.g. Weisz et al. 2014; 2019; Ruiz-Lara et al. 2018b; Bidaran et al. 2022). In Bidaran et al. (2022) it was suggested that these recent bursts of star formation can be associated to late cluster infall. This is compatible for the two cluster UDGs, while the remaining one, the group UDG1137+16, presents a slightly higher abundance ratio than its cluster counterparts, despite sharing a similar SFH type. However, this is a UDG with tidal features and therefore this could be the result of a past interaction.

There seems to exist a clear relation between the SFHs, the ages, and the α -enhancement of UDGs, which in turn is related to the environment they reside in. Moreover, there seems to be a trend with stellar mass, with those UDGs quenching the fastest (and therefore with the most elevated $[\text{Mg}/\text{Fe}]$), being all on the massive end of the UDG population. To further investigate this, Fig. 5 shows the result of averaging the SFHs of the UDGs according to their environment. As the majority of the UDGs in our sample and in the literature are located in clusters, the left panel focuses on this environment alone. The grey region shows the average SFH and the intrinsic scatter for all cluster UDGs (from this work, FM+18 and Ruiz-Lara et al. 2018a). On average, cluster UDGs start forming early, reaching 50 per cent of their stellar masses in $\sim 4-5$ Gyr after the Big Bang. Both Coma and Perseus UDGs seem to build up their stellar mass

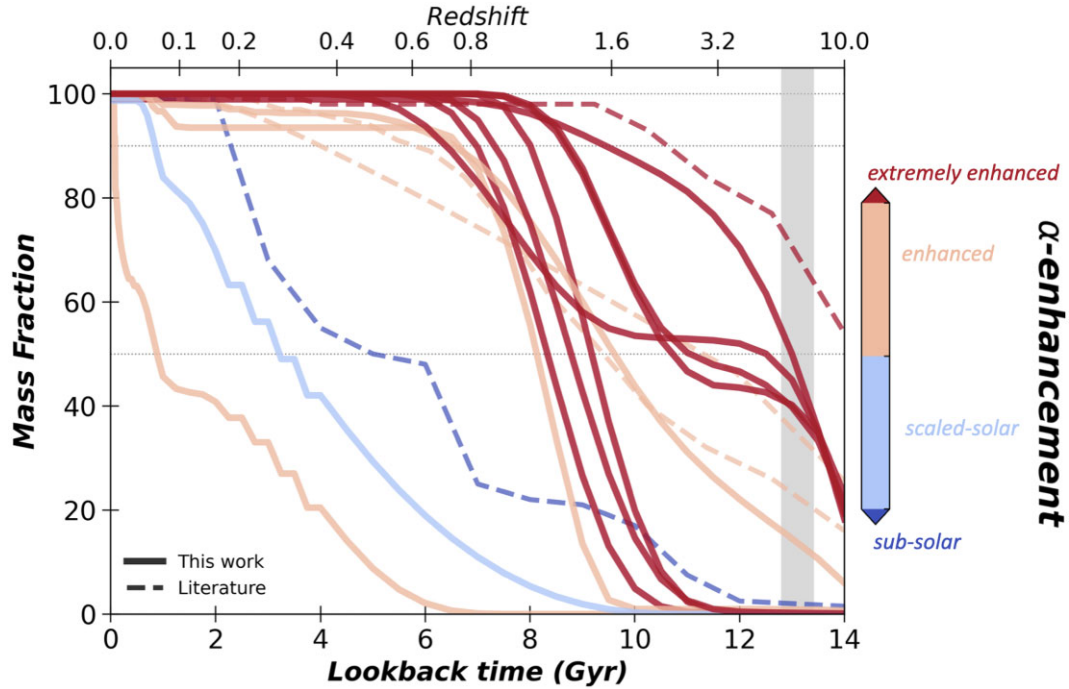


Figure 4. Cumulative mass fraction of each UDG in this work (solid line), colour-coded by their level of α -enhancement. We note that this colour classification is consistent regardless of using $[\alpha/\text{Fe}]$ or $[\text{Mg}/\text{Fe}]$, and has been chosen to unify the different results from the techniques presented in Appendix B1. We also include other Coma UDGs from FM+18 and Ruiz-Lara et al. (2018a) with existing $[\text{Mg}/\text{Fe}]$ abundance ratios (dashed lines). Horizontal dotted lines mark the 50, 90, and 100 per cent mass fraction in stars. We note that none of the UDGs was quenched by reionization, marked as a grey vertical band (Fan, Carilli & Keating 2006). A variety of SFHs are found for UDGs, with a clear relation with the α -enhancement. Rapid SFHs tend to show elevated values regardless of whether or not the galaxy started forming very early on (extremely high, >0.4 dex) or delayed in time (high, 0.2 – 0.4 dex). In contrast, those with more extended SFHs tend to be scaled solar (<0.2 dex) or even sub-solar (<0.0 dex).

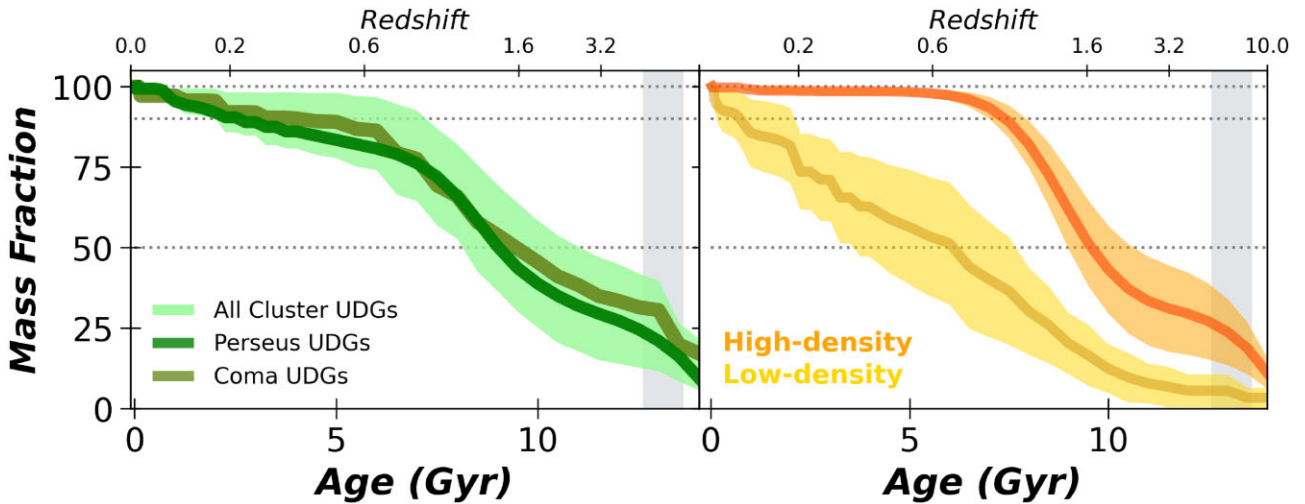


Figure 5. Similar to Fig. 4 but now UDGs are averaged by different properties. Left panel: The average of all cluster UDGs from this work, regardless of their local environment, with the intrinsic scatter shown as the shaded light green area. Perseus and Coma UDG averages are shown by solid green lines as described in the legend. It seems that UDGs build up their stellar mass in a similar way in the Perseus and Coma clusters. Right panel: all UDGs in this work separated by their local environment: low-density regions such as field, groups, and recent infalls (yellow), and high-density regions as the rest of cluster UDGs (orange). We find that UDGs in high-density environments tend to form earlier and faster than those in low-density environments.

in a similar way, although Coma UDGs quench slightly earlier than Perseus ones (5 versus 2 Gyr, respectively). However, comparing these averages with the individual SFHs in 4 it is clear that they are not a representative of the whole cluster population, as some UDGs have clearly much earlier and faster SFHs.

If UDGs fell into the cluster late, this will skew the results towards younger, more extended SFHs. This is the case of PUDG-R24, Yagi275 from FM+18 and Yagi90 from Ruiz-Lara et al. (2018a) (see also Section 4.2). In such cases, their immediate environment is better represented by a low-density one, more similar to groups. The

right panel of Fig. 5 now separates our complete sample of UDGs by (local) environment: low-density (field, group, and cluster outskirts) and high-density (the rest of cluster UDGs). With this separation we find that UDGs in high-density environments formed earlier and faster, having built the bulk of their stars $\sim 8\text{--}9$ Gyr ago. Instead, UDGs in low-density regions on average started forming with a ~ 4 Gyr delay, forming stars until almost the present day. This result supports the trends found by Buzzo et al. (2022) and Barbosa et al. (2020), where older UDGs tend to be found in denser environments than younger ones. We caution, however, that low-density UDGs conform about half of the UDGs with SFHs, and that for example there are no field galaxies in this group of UDGs.

Similar SFH trends have been found in Local Group dwarfs (e.g. Gallart et al. 2015; Weisz et al. 2015; Ruiz-Lara et al. 2018b; Collins & Read 2022). In fact, Gallart et al. (2015) separated these dwarf galaxies between fast (if quenched around 8 Gyr ago) and slow (that started forming later in time and that show more extended SFHs). Interestingly, if we split our UDGs into fast/slow according to the time-scales derived above, we find virtually the same SFHs (same shape and formation time-scales) for the fast and high-density UDGs, and for the slow and low-density ones. This further reinforces the result that environment seems to play a role in how UDGs form.

4.2 Phase-space and infall time of UDGs

Both the global (cluster, group, or field) and local (center, outskirts, filament, etc...) environment appear to influence the evolution of UDGs. For example, in FM+18 it was found that the stellar populations of UDGs correlate with cluster-centric distance. UDGs in the outskirts were younger, most likely related to a later cluster infall time. As a consequence, they had later quenching times and thus presented more extended SFHs (see also Alabi et al. 2018).

In Fig. 6 we provide a position–velocity phase-space study of our UDG sample, in order to test whether or not the different stellar populations in the cluster UDGs correspond to their location in a phase-space diagram. It shows the velocity of the galaxy compared to the mean velocity of each cluster (normalized by the velocity dispersion of the cluster) versus the projected cluster-centric distance (normalized by the virial radius, R_{200}). For each cluster analysed here we assume: $R_{200} = 2.9$ Mpc, $V_r = 6943$ km s $^{-1}$, and $\sigma = 1031$ km s $^{-1}$ for Coma; $R_{200} = 1.7$ Mpc, $V_r = 1137$ km s $^{-1}$, and $\sigma = 752$ km s $^{-1}$ for Virgo; and $R_{200} = 1.8$ Mpc, $V_r = 5366$ km s $^{-1}$, and $\sigma = 1324$ km s $^{-1}$ for Perseus (Alabi et al. 2018; Gannon et al. 2022). Overplotted are the cluster UDGs from Table 1, with symbols representing the different clusters and the colour scale in each panel corresponding to the relevant properties (when available).

In Fig. 6 there is also overlaid a scheme that correlates phase-space location with the different infall epoch from the simulations of Rhee et al. (2017). The shaded sections go from a dark to a light shade, corresponding to the transition from ancient to more recent infalls (i.e. ‘very early infall’, ‘early infall’, ‘mixed times’, ‘late infalls’, and ‘yet to infall’). We note, however, that these regions are statistically significant only with large samples of galaxies. In all regions there could be large numbers of interlopers: galaxies that do not belong to the cluster but that projection effects place them within projected cluster phase space, or cluster galaxies with a late infall time that are also placed in the centre due to projection effects.

To begin, we wish to confirm the validity of the infall regions, by comparing them to the quenching time-scale inferred from the SFHs of the UDGs. Assuming that cluster infall is the cause of quenching, the infall time can be estimated as $t_{\text{inf}} \sim t_{90} - 1.5$ Gyr. However, other UDG formation scenarios invoke self-quenching processes.

This means that some UDGs might have formed and quenched before cluster infall, in which case early quenched galaxies in late infall regions can be found. According to Rhee et al. (2017), ancient infallers are considered to be galaxies that took up to 7.25 Gyr to fall in the cluster, roughly conforming the ‘very early infall’ and ‘early infall’ regions. Nine UDGs (PUDG-R27, VCC 1287, Yagi358, Yagi093, Yagi098, Yagi276, Yagi392, Yagi418, and OGS1) are located in the ‘very early infall’ region, with other two (DF07 and PUDG-R15) in the ‘early infall’ region. They have quenching time-scales compatible with their location with the exception of Yagi093, Yagi098, Yagi276, and Yagi418, which present very long quenching time-scales that are more similar to those expected for intermediate infallers (time-scales of 7.25 to 10 Gyr). These UDGs could therefore be the result of a projection effect (i.e. interlopers). Four UDGs are located in the ‘mixed times’ regions (PUDG-S74, PUDG-R84, DFX1, DF44) all presenting quenching time-scales of $\sim 6\text{--}8$ Gyr. Finally, galaxies in the ‘late infall’ region should have taken from 10 to 13.7 Gyr to infall in the cluster according to Rhee et al. (2017). Three UDGs (PUDG-R24, Yagi275, and Yagi090) are located in this region, in agreement with their long quenching time-scales and relatively young ages. We note that for DF17, J130038.63+272835.3, and J130026.26+272735.2, we do not have an estimate of the quenching time and they are thus not included in this first panel.

Typically, low-mass galaxies that have been in the cluster environment for longer are found to be predominantly old and with higher $[\alpha/\text{Fe}]$ abundance ratios than recently accreted ones (e.g. Smith et al. 2009; Liu et al. 2016; Pasquali et al. 2019; Gallazzi et al. 2021; Bidaran et al. 2022). We find that the majority of the UDGs in the ‘very early infall’ and ‘early infall’ regions show indeed elevated α -enhancement values (following the scheme presented in Fig. 4). Interestingly, these UDGs also present some of the highest metallicities in the sample, similarly to what was found in Bidaran et al. (2022) for classical dwarfs. Conversely, low α -enhancement ratios are only found in mixed times towards late infall regions, compatible with the SFHs trends found in the previous section. There are two UDGs whose stellar populations and phase-space do not seem to match. Yagi090, one of the very late infallers, presents a very high $[\text{Mg}/\text{Fe}]$ (Ruiz-Lara et al. 2018a). According to its SFH it was quenched 5 Gyr ago, which would imply that this UDG was quenched prior to infall. DF44 instead presents a relatively low $[\text{Mg}/\text{Fe}]$ (Villaume et al. 2022) even though it is located in the ‘early infall’ region. However this UDG has been reported to belong to a group that is currently plunging through the Coma Cluster, and thus has been previously considered to be a low-density UDG. We note that this dependence of the α -enhancement with galactocentric distance (and lack of thereof for the ages and metallicities) is similar to the one recently reported by Romero-Gómez et al. (2023) for dwarfs in Fornax. No other clear trends between the phase-space infall times and the structural properties of UDGs are observed.

4.3 Stellar populations trends

We next try to elucidate the most plausible origin of the spectroscopic sample of UDGs by comparing them to other low mass systems that have been proposed to be related. For example, GCs are broadly understood to be old and metal-poor, having formed during the early stages of galaxy formation. Some works suggest that UDGs form in the same epoch as the GCs (i.e. the failed-galaxy type; Forbes et al. 2020a) or that they may even form entirely from disrupted clusters (i.e. NGC 5846-UDG1; Danieli et al. 2022), so any UDG that shares similar properties as GCs could be considered of this type. On the contrary, if UDGs are simply puffed up dwarf galaxies, one would

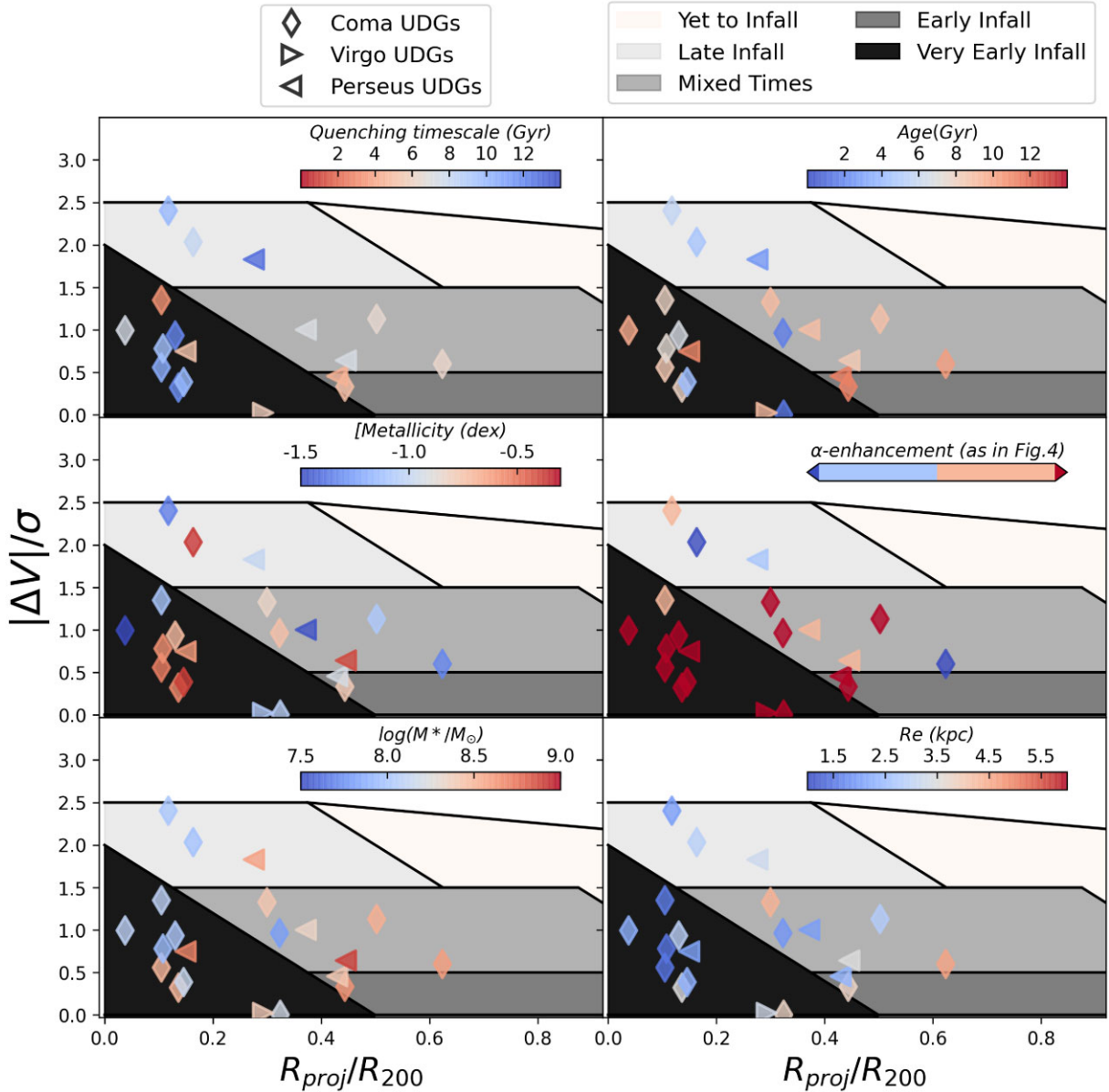


Figure 6. Phase-space diagram for the full spectroscopic sample of cluster UDGs. In each panel, the different infall regions from Rhee et al. (2017) are shown as shaded regions, going from darker to lighter to show UDGs infalling at progressively later times. In each panel, cluster UDGs are colour-coded by the respective property (i.e. quenching time, age, $[M/H]$, α -enhancement, stellar mass, and half-light radius), as shown in each colour scale. We note that quenching time and stellar age have opposite colour scales to guide the eye and that the colour scheme in the α -enhancement panel follows that presented in Fig. 4. Different symbols are used to differentiate the cluster each UDG belongs to. No clear trends are seen between infall regions and most of the UDG properties. We find that very early/early infall UDGs mostly present more elevated α -enhancement values than late infalls, compatible with their SFHs.

expect to find very similar stellar populations to classical dwarfs but with possibly different stellar masses.

Fig. 7 presents the age–metallicity relation, the $[Mg/Fe]$ –metallicity relation, and the mass–metallicity relation, some of the most relevant scaling relations regarding the stellar populations of galaxies. In the $[Mg/Fe]$ –metallicity panel the relation that was derived for UDGs is shown (Rong et al. 2020b). We include it as it was obtained using mostly quiescent UDGs, although star forming isolated UDGs were also included in the fit. In the mass–metallicity panel the scaling relation of Panter et al. (2008) is shown, which is an extension of the Gallazzi et al. (2005) relation of massive ETGs towards intermediate stellar masses (roughly covering the UDG mass regime). The theoretical mass–metallicity relation of high redshift

galaxies ($z \sim 2$) from Ma et al. (2016) is also shown. The reader is referred to Buzzo et al. (2022) for an extensive discussion on the caveats of the different mass–metallicity scaling relations at the low mass end, here we use the relation from Simon (2019). In each panel, the coloured contours correspond to different types of low-mass systems: GCs, dSphs, and dEs (spectroscopic values from Janz et al. 2016; FM+18; Recio-Blanco 2018; Naidu et al. 2022 and Romero-Gómez et al. 2023). These control galaxies cover a range in environments: from the Local Group to clusters (Fornax, Virgo, Coma). dSphs have lower stellar masses than the UDGs, and the dEs have mostly higher masses.

The age–metallicity panel of Fig. 7 (top left) shows that the majority of UDGs have ages and metallicities similar to dEs for the

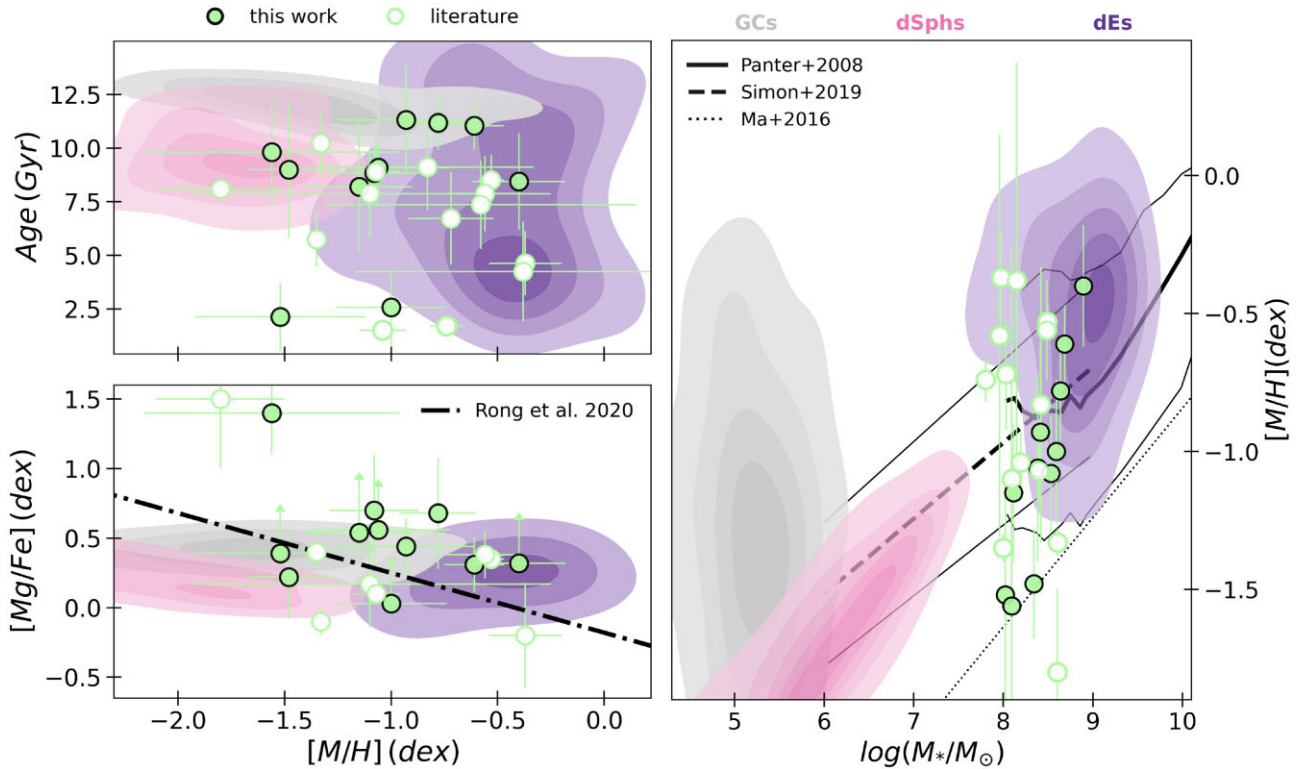


Figure 7. UDG scaling relations compared to other low mass galaxies and GCs. The figure shows the relation of the UDG metallicities with stellar age (top left), $[\text{Mg}/\text{Fe}]$ (which includes the relation for UDGs defined in Rong et al. 2020b; bottom left), and the well-known mass–metallicity relation (right panel). For the latter, the scaling relations and their intrinsic scatter of local massive and low-mass galaxies (Panter et al. 2008; Simon 2019, respectively) and the theoretical prediction at $z \sim 2$ (Ma et al. 2016) are also shown. UDGs from this work are plotted in filled symbols, and UDGs from the literature with open ones. Contours with different colours correspond to the locations of samples of GCs (grey), Local Group dSphs (pink), and Virgo/Fornax/Coma dEs (purple), from Janz et al. (2016); FM+18; Recio-Blanco (2018); Naidu et al. (2022); and Romero-Gómez et al. (2023). We have marked the UDGs with the α -enhancement obtained from pPXF as possible lower limits by arrows. UDGs show overall elevated $[\text{Mg}/\text{Fe}]$, in some cases much higher than any other galaxy type known. UDGs show a large dispersion in their stellar ages and metallicities, scattering around the local mass–metallicity relation. A handful of UDGs seem to be more compatible with high- z galaxies, suggesting these could have a different origin (i.e. ‘failed-galaxy’ like).

more metal-rich UDGs, while more metal-poor ones share similar age–metallicities to dSphs. Three UDGs (PUDG-R15, PUDG-R27, and DF07) match the GCs distribution, although in a region that also matches the dEs. A few other UDGs would be compatible with the GC distribution given the errors in metallicity. We note, moreover, that some GCs in the control sample have ages of ~ 7 – 8 Gyr and $[\text{M}/\text{H}] \sim -0.8$ dex, similar to dEs. However, they are so few compared to the bulk of GCs that are not shown by the contours. UDG1137+16 is the only UDG that is clearly different to any of the comparison samples, with a metallicity that is more typical for older galaxies. This UDG is in fact an outlier in the majority of the scaling relations, which might be a consequence of the tidal features that it shows (see e.g. Gannon et al. 2021).

In the $[\text{Mg}/\text{Fe}]$ –metallicity panel (bottom left) it is shown that most UDGs are also very α -enhanced, some of them more than any other known galaxy of the comparison distributions. Only five UDGs in the sample present scaled solar or even sub-solar $[\text{Mg}/\text{Fe}]$. In this panel we also include the $[\text{Mg}/\text{Fe}]$ –metallicity relation derived for UDGs (Rong et al. 2020b). However, we find that the bulk of our UDGs is more compatible with a flat relation, similar to what is seen for the comparison samples. We remind the reader that this is mainly caused by the limitations discussed in Section 3.1 and the Appendix B1 on the models and methods used to obtain this quantity. UDGs mostly cover the $[\text{Mg}/\text{Fe}]$ range of GCs and dwarfs at all metallicities, and only the most extreme ones (DGSAT I and Yagi358) having $[\text{Mg}/\text{Fe}]$

that do not match any galaxy in the comparison sets. In particular, DGSAT I has been shown to be an extremely bizarre galaxy in many respects (Martín-Navarro et al. 2019). Its high $[\text{Mg}/\text{Fe}]$ value has been potentially explained by a ‘depletion’ of $[\text{Fe}/\text{H}]$ rather than an over-abundance of $[\text{Mg}/\text{H}]$. From our line index analysis we also measure an extremely low $[\text{Fe}/\text{H}]$ but high $[\text{Mg}/\text{Fe}]$ for Yagi358. Given that they both have overall similar stellar population properties, we speculate that Yagi358 could have formed in a similar process to DGSAT I.

The right panel of Fig. 7 shows different mass–metallicity relations. Overall, UDGs in this sample scatter around the mass–metallicity relation expected for galaxies of $\log(M_*/M_\odot) \sim 8$ – 9 . Despite the large scatter in the metallicity distribution, UDGs seem to have on average lower metallicities than the dEs at fixed stellar mass. Five UDGs present much lower metallicities than the rest of UDGs and classical dwarfs, laying below the local mass–metallicity relations. These are Yagi358, PUDG-R84, DF44, and DGSAT I. DGSAT I has been already reported to be an unusual galaxy, being an extreme outlier of most scaling relations (e.g. Martín-Navarro et al. 2019). As the failed-galaxy formation scenario assumes early quenching, UDGs formed this way are expected to follow the mass–metallicity relation of high- z galaxies rather than local ones. We find that these other four UDGs with very low metallicities indeed match the theoretical mass–metallicity relation at $z \sim 2$ (Ma et al. 2016), which would support them as failed-galaxies.

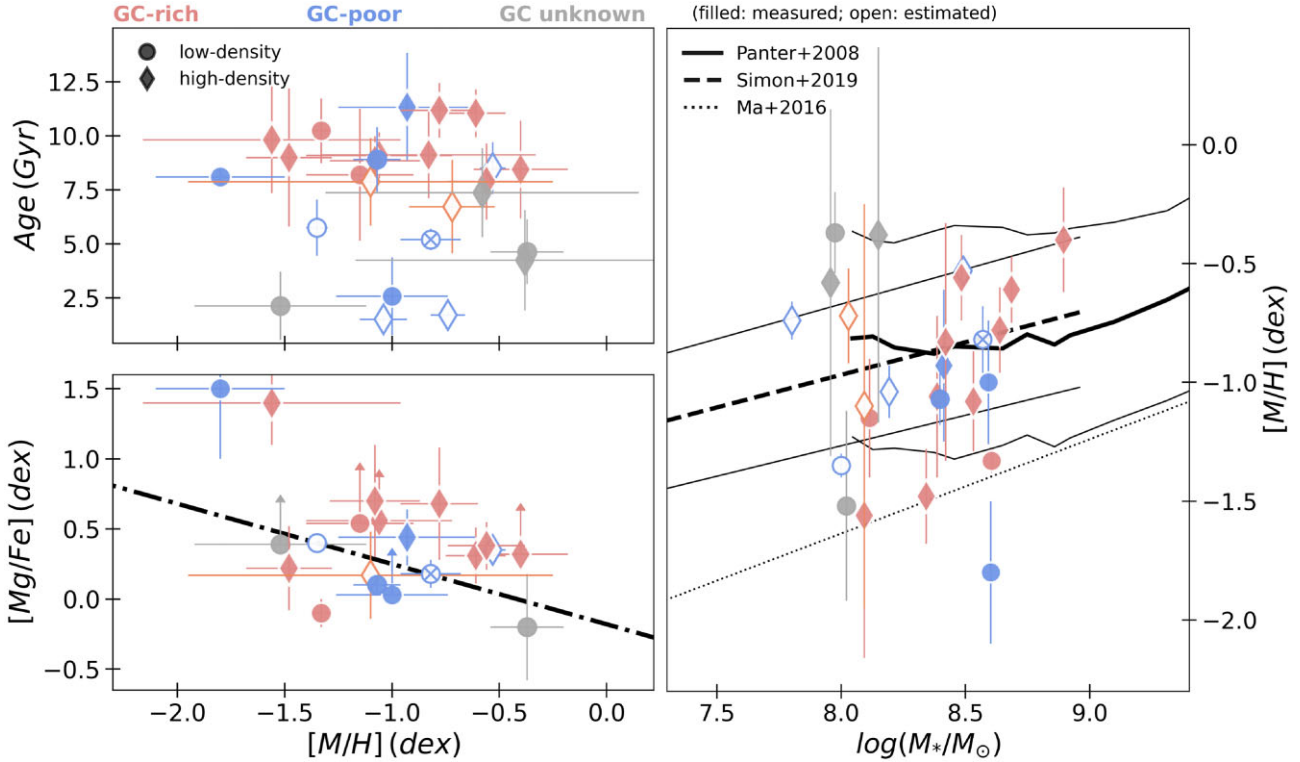


Figure 8. Main scaling relations for the total sample of UDGs with GC-richness and environment. Same as in Fig. 7 but now UDGs are colour coded by GC richness: red for GC-rich, blue for GC-poor, and grey for UDGs with no GC information. Filled symbols correspond to UDGs with proper GC counts, open symbols if the GC-richness comes from the visual estimate described in Appendix A. Circles correspond to low-density UDGs and diamonds to high-density ones, classified as in Fig. 5. The stacked value for the star forming, isolated UDGs of Rong et al. (2020b) is marked with an X in its symbol. We find that GC-rich UDGs have intermediate-to-old ages, while the young ones are all GC-poor. However we do not find any evident trends of GC richness with $[Mg/Fe]$ nor in the mass–metallicity relation.

4.3.1 Trends with environment and GC-richness

We have seen in Section 3.3 a dependence between the local environment and the stellar populations of UDGs and we have shown in Fig. 7 that the bulk of UDGs in this sample have mostly stellar population properties similar to classical dwarf galaxies. While this is compatible with a puffy dwarf origin, it is difficult to explain their populous GC systems, as reported for many cases. For example, Lim et al. (2018) and Forbes et al. (2020a) showed that Coma UDGs tend to have higher GC specific frequencies than comparable classical dwarf galaxies. Additionally, it is not clear whether there are any differences in the overall properties of GC-rich and poor UDGs. For instance, a trend with GC-richness was found by Buzzo et al. (2022), with GC-poor UDGs presenting higher metallicities than GC rich ones. From their SED fitting they also found that the oldest UDGs tend to prefer cluster environments, supported by our results in Fig. 5.

Therefore we investigate next if similar trends are recovered in this spectroscopic sample. Fig. 8 shows the sample of UDGs now separated between low density (field and group UDGs, but also UDGs considered to be recent infalls), and high density environments (the rest of cluster UDGs), as it was done in Fig. 5. To complement our comparisons within environments, we also include in this figure the stellar populations properties measured for a sample of stacked isolated, star forming UDGs (crossed circle; Rong et al. 2020b), as one formation scenario for UDGs predicts that these may evolve from star-forming in the field to quenched in the cluster environment. The colour scheme describes the GC-richness of the UDGs, where

we use both proper counts but also UDGs with a visual classification (see Appendix A).

We find that overall, most (13/16) high density UDGs show intermediate-to-old ages ($\gtrsim 7$ Gyr), confirming the results from Fig. 5, now with an expanded sample. Young UDGs (≤ 7 Gyr) are found evenly in terms of local environment. This result is thus further compatible with Buzzo et al. (2022), although their sample did not include UDGs with such young ages.

Similarly to Buzzo et al. (2022), we find that the UDGs lying on the mass–metallicity relation of high- z galaxies are all GC-rich (without considering DGSAT-I, which has been repeatedly reported as an unusual galaxy). Having a large number of GCs is one of the main expectations of the failed-like galaxy scenario, further supporting such origin for these UDGs. However, we do not find any trend between GC richness and galaxy metallicity as reported in Buzzo et al. (2022). In the latter, GC-poor UDGs have higher metallicities than GC-rich ones, while we have many GC-rich UDGs with elevated metallicities in this study.

There are many possibilities to explain these differences. For example, there seems to be a systematic effect towards lower metallicities in the SED fitting compared to spectroscopy (see Appendix B3). It is unclear if the impact would be the same for UDGs at all ages, but we find that in particular the UDGs with the largest difference in metallicity are the oldest ones (Fig. B3). There could also be a sample bias, as there are no GC-rich UDGs that are metal-rich in the work of Buzzo et al. (2022), while in this work there are several. We note that in Barbosa et al. (2020), also from SED fitting but targeting field UDGs, a large number of UDGs with such

higher metallicities were reported, but unfortunately no GC counts are yet available. We also note that our sample is biased towards both high-density environments and GC-rich UDGs. In addition, most of the GC-poor ones are coming from a visual analysis, and not from proper GC counts. None the less, we chose to use the visual GC counts as we are primarily interested in a comparative study of the properties of UDGs with/without GCs, not on their absolute numbers. We note, however, that the trend is still not found even if we exclude these galaxies from the analysis, hence further work studying the GC systems of these galaxies is necessary to understand if our findings hold.

Additionally, no particular trend is seen with the α -enhancement of the UDGs and their GC system richness. At most, it seems that GC-poor UDGs follow the Rong et al. (2020a) relation more tightly than GC-rich ones (with the exception of DGSAT-I, again). Only two out of the six low-density UDGs with a GC estimate lie above the [Mg/Fe]–metallicity relation, an area mostly populated by the high-density UDGs.

Bearing in mind the aforementioned caveats, the take away message is that until we have larger and more complete samples of UDGs of varying GC richness and environments, any trends obtained may be simply an effect of the sample selection itself and must be taken with caution for now.

4.4 Comparisons to simulated UDGs

Many different simulations have tried to reproduce the observed properties of UDGs by proposing a myriad of formation mechanisms. Here we compare some of the observed properties discussed in this work to different sets of simulated UDGs. With this exercise we wish to have a qualitative assessment of the compatibility of simulations and observations. We note, however, that most of the simulations that form UDGs do *not* form the more compact classical dE/S0s, making up the vast majority of observed dwarfs in the same mass range – an example of the ‘dwarf diversity’ challenge to current models for galaxy formation (Sales, Wetzel & Fattahi 2022).

The first simulation we compare to is the NIHAO cosmological simulation (Wang et al. 2015). The UDGs from this simulation have been studied in detail by Di Cintio et al. (2017), Jiang et al. (2019), and Cardona-Barrero et al. (2020). These simulations create UDGs in low-density environments through strong stellar feedback that puffs up dwarf galaxies. The simulated UDGs present bursty SFHs, intermediate ages, and low metallicities (Di Cintio et al. 2017; Cardona-Barrero et al. 2020).

We also use UDGs from the FIRE simulations of Chan et al. (2018). In this case, the simulation mimics the creation of UDGs in a cluster environment by quenching the simulated UDG at a given time. Two different scenarios were investigated: one where the galaxy is artificially quenched very early on (roughly ~ 2 Gyr after the Big Bang), while the other scenario allows UDGs to evolve for longer times, quenching only about 4 Gyr ago to simulate late cluster infalls. Similar to NIHAO, FIRE primarily creates UDGs through stellar feedback before their artificial quenching.

We include simulated UDGs from the Romulus cosmological simulation (Tremmel et al. 2017), which are all considered to be like puffed-up dwarfs. The sample consists of a set of cluster UDGs (from RomulusC; Tremmel et al. 2020) and isolated UDGs (from Romulus25; Tremmel et al. 2017; Wright et al. 2021). Those in clusters present different infall times, although the authors found that UDGs seem to fall into the cluster earlier than regular dwarfs. Those in the field are the result of major mergers occurring at earlier stages than non-UDGs, increasing their spin and redistributing the star

formation towards the outskirts. Therefore for both cluster and field UDGs in the Romulus simulations, their nature is finally shaped by the passive evolution of their stellar populations (Román & Trujillo 2017).

We also include two different sets of Illustris-TNG simulations. We employ the simulated UDGs of Illustris-TNG50 (Benavides et al. 2023), which are a mixture of high and low density environments (sampling cluster, groups, and field UDGs). These also include the backplash UDGs presented in Benavides et al. (2021). Overall, these authors found that cluster UDGs tend to be smaller and less massive than field UDGs or recent infallers, with slightly lower metallicities. We also include the sample of Sales et al. (2020), which presented UDGs from the Illustris-TNG100 simulation. These are classified as either ‘born’ UDGs (already an UDG before cluster infall), or those that ‘become’ a UDG due to tidal effects within the cluster after infall (‘tidal UDGs’). The latter are typically accreted earlier in the cluster (in 1 to 6 Gyr after the Big Bang) and have larger stellar masses at the time of the infall. We caution the reader that all TNG simulations have some difficulties in reproducing the metallicities of galaxies (e.g. Nelson et al. 2018; Sales et al. 2020; Benavides et al. 2021). For instance, TNG50 simulations were originally re-scaled to match the observed scaling relations at $\log(M_*/M_\odot) = 9$ (see Benavides et al. 2023). The shift applied (~ 0.5 dex) is comparable to the reported discrepancy in Nelson et al. (2018) for all galaxies in TNG100. Therefore, we apply the same re-scaling to the TNG100 sample of UDGs (Benavides et al., private communication).

Fig. 9 shows the stellar mass–size relation (top) and the mass–metallicity relation (bottom), two of the main scaling relations studied in this work. We plot the corresponding simulated and observed UDGs according to their local environment: high-density environments (clusters; left panels) and low-density ones (groups and field; right panels). UDGs that are located in the ‘late-infall’ region of Fig. 6 or that are known to be in an infalling group, such as DF44, are shown in all panels. The mass–size panels are simply shown to have all set of simulations together with the observed UDGs. For all sets of simulations we have only considered UDGs that match the observational criteria of $R_e \geq 1.5$ kpc. Many of the simulations produce UDGs with lower stellar masses than in our observed sample, in particular in low-density environments. This is most likely the result of spectroscopic observations being biased towards more massive UDGs.

In the stellar mass–metallicity relations of Fig. 9, most simulations predict a relation that changes only very weakly with environment – similar to observations of a near-universal red sequence for classical dwarfs. Most of the observed UDGs in high-density environments share location with the simulated UDGs (lower left panel), while UDGs in low-density environments, as well as late infallers, are on average offset to lower metallicities (lower right).

NIHAO has a mass–metallicity relationship in their simulations systematically lower than that of other simulations and what is observed in this stellar mass regime. It is not clear that NIHAO can produce non-UDGs following the mass–metallicity relationship in this mass regime. This is a result of their strong internal feedback, which expands the galaxies and can expel the metals. Despite the systematically lower metallicities, NIHAO does not produce UDGs as low as some of the observed UDGs (e.g. that match the theoretically $z \sim 2$ mass–metallicity relation). Furthermore, this stellar feedback mechanism was rejected to explain the formation of the star forming isolated UDGs in Rong et al. (2020a) as these match the expected scaling relations and are therefore better represented by other simulations.

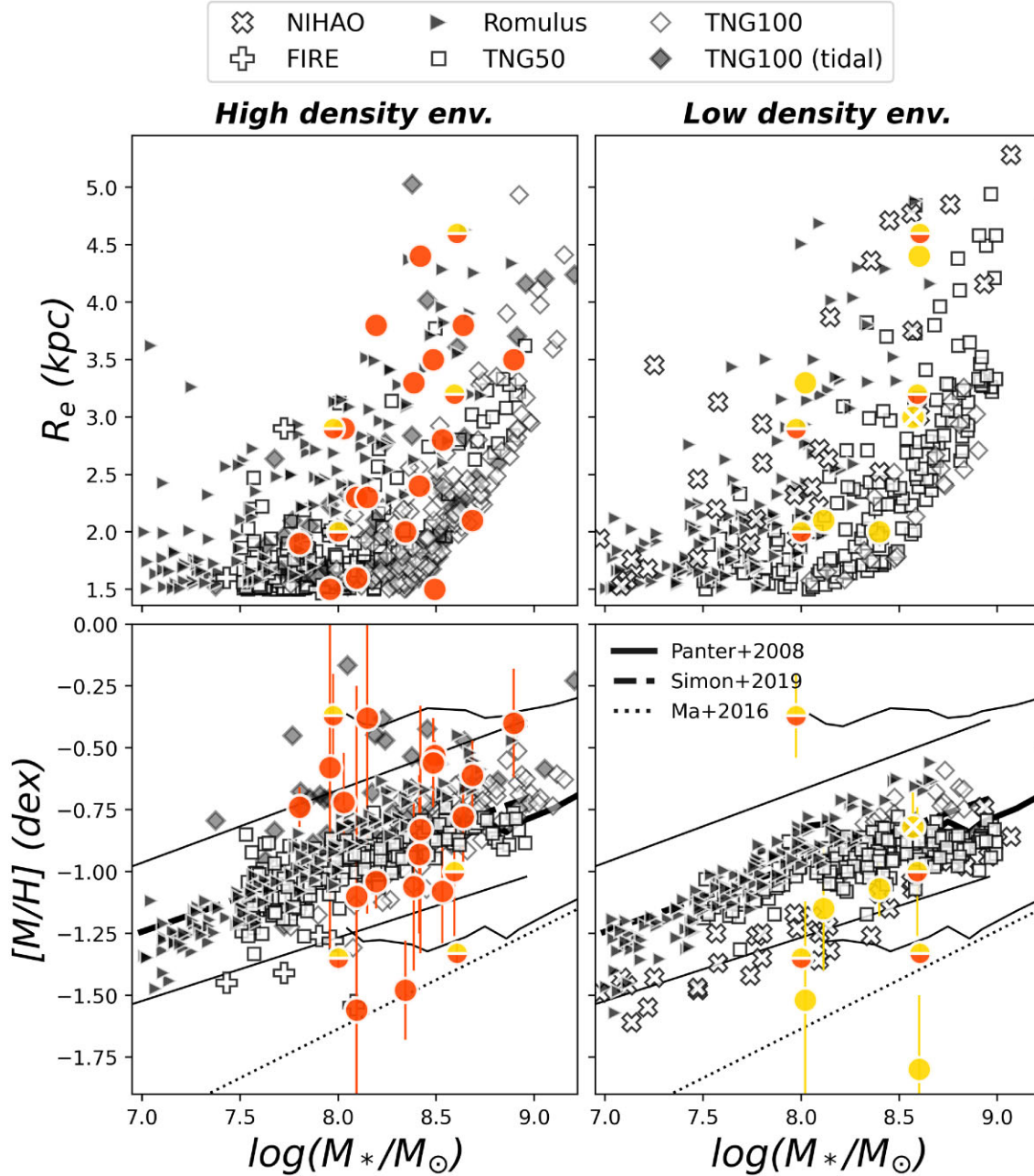


Figure 9. Comparison of observed and simulated UDGs separated by environment: high-density (cluster, left column) and low-density (field, group; right column). The following sets of simulations cover high density environments: FIRE (white crosses; Chan et al. 2018), Romulus (black triangles; Tremmel et al. 2020), TNG50 (white squares; Benavides, private communication) and TNG100 (white diamonds for cluster UDGs, filled in grey for those that are considered to be tidally stripped galaxies from Sales et al. 2020). Romulus, TNG50 and TNG100 are also shown in the low-density section, together with the UDGs from NIHAO (white x-symbols; Di Cintio et al. 2017). The top panels present the mass–size relations, where simulated UDGs are limited to $R_e \geq 1.5$ kpc to match the observations, and TNG50 metallicities have been re-scaled as in Sales et al. (2020). Local scaling relations (solid lines) at the high mass (Panter et al. 2008) and low-mass ends (Simon 2019), and their scatter (thin solid line), are shown. The high redshift theoretical mass–metallicity relation of Ma et al. (2016) is also shown in the bottom panels (dotted line). The spectroscopic sample of UDGs is represented by circles, coloured by their local environment: orange for UDGs in high density environments, yellow for low density ones. Circles with half-and-half colours denote UDGs thought to be very recent cluster infalls or located in filaments. These are shown in all panels, as their properties could be similar to both field UDGs (before infall) and cluster ones. Like in the previous figure, the properties of the isolated, star forming UDGs from Rong et al. (2020a) are marked by a white X. It is clear that there is not a single simulation set that reproduces all type of UDGs at the same time. UDGs with the lowest metallicities and that lay below the local mass–metallicity relation (i.e. most likely failed-galaxy type), are the only UDGs not reproduced by any set of simulations.

It is important to note that there is not one single formation scenario proposed for UDGs and that there is evidence of UDGs forming via alternate pathways in Fig. 9. For example, Yagi275 and Yagi276 are two of the UDGs with the highest metallicities for their

stellar mass. One possibility to explain galaxies above the relation is that such elevated metallicities are indicative of a tidally stripped galaxy (e.g. Gallazzi et al. 2005; Ferré-Mateu et al. 2018a; 2021; Du et al. 2019). Interestingly, these two UDGs are located where the

simulated UDGs resulting from tidal effects are found (in the high-density environment panel; Sales et al. 2020). This scenario would be plausible for Yagi276, as it has long resided in the cluster (owing to its infall diagram), which is one of the requirements of this scenario. However, it would only work for Yagi275 if this UDG was already a stripped galaxy before infall, as it is considered to be a late infall from both its phase-space location and SFH.

The most interesting cases are the UDGs that present the lowest metallicities, laying below the local mass–metallicity relation but follow instead the theoretical relation expected for high- z galaxies. From those, PUDG-R84 and Yagi358 (in high-density environments), and DF44 (as infall galaxy), present early and fast formation and are GC-rich, indicative of a large dark matter halo. These three UDGs are therefore the best candidates to failed-galaxy like UDGs considering all their inferred properties (see also previous sections). Interestingly, these three UDGs are not reproduced by any simulation in the high-density environment. The only simulated galaxy compatible with this scenario is one of the FIRE UDGs that was forced to quench very early on (Chan et al. 2018). This further supports the failed-galaxy origin for these subset of UDGs.

DGSAT-I and UDG1137+16 (both in low-density environments), although they also present low metallicities and lay below the relations, do not met the rest of expectations for being considered failed-galaxy like UDGs. The uniqueness of the exotic chemistry of DGSAT-I has already been discussed (e.g. Martín-Navarro et al. 2019) but not reproduced yet in simulations. UDG1137+16 has a complicated formation as shown by its tidal signatures (Gannon et al. 2021), but it has metallicity similar to the NIHAO UDGs of similar stellar mass.

We have at hand several simulations that attempt to create UDGs following different prescriptions, although there is not a single simulation that can reproduce the diversity of observed UDGs. In addition, many present caveats (e.g. cannot reproduce metallicities and stellar masses of observed UDGs). Finally, galaxies with properties similar to UDGs of failed-galaxy origin have not successfully been simulated as of yet.

5 SUMMARY

In this work we have analysed the stellar populations of the largest spectroscopic sample of observed UDGs to date (25), covering different environments and GC richness. This sample comprises 11 UDGs for which new spectroscopy from KCWI on the Keck Telescope was obtained, along with 14 literature UDGs. The new sample is mostly comprised of UDGs in Perseus, Coma, Virgo and two UDGs in groups of galaxies. For all UDGs in our sample, we obtain their structural properties and number of GCs from published works (Gannon et al. 2020; 2021; 2022; 2023 and Forbes et al. 2021). For some of the UDGs without previous GC counts, we have performed a visual inspection to increase the sample, although we caution the reader about using these estimates at face-value. We have carried a new stellar population analysis for the UDGs and discussed different relations and dependencies in order to elucidate their most plausible origin (see Table C1 for a summary of all properties and suggested origins).

(i) We have derived the main stellar population properties (such as mean ages, metallicities, and α -enhancement) of these 11 new UDGs via both a full spectral fitting approach (using pPXF) and using classical line index measurements. These UDGs have a mean age $\langle t_M \rangle = 8.3 \pm 3.3$ Gyr, metallicity $\langle [M/H] \rangle = -1.03 \pm 0.37$ dex, and mean α -enhancement of $\langle [Mg/Fe] \rangle = +0.51 \pm 0.32$ dex.

(ii) We find a clear trend between the SFH and the α -enhancement. UDGs that formed earlier and faster present the highest enhancements, while those forming over a longer period of time show scaled solar values. UDGs with intermediate ages also show relatively high abundance patterns because they formed relatively quick despite beginning their star formation later in time.

(iii) Local environment seems to be relevant in the way UDGs are formed. UDGs in high density environments show very early onset of star formation, quenching within time-scales of ~ 6 Gyr. All these UDGs are at the high mass end of the UDG population. On the contrary, UDGs in low density environments (groups and cluster outskirts), started forming their stars roughly 10 Gyr ago, and did not quench until very recently.

(iv) We thus propose different formation scenarios for the UDGs according to their environment and stellar populations. UDGs that formed early and quickly, thus quenching at early ages, can be considered ancient infalls into the cluster. These can be UDGs with a failed-galaxy origin if their metallicities are low, or tidally stripped UDGs if their metallicities are high. The rest of UDGs in the sample are more compatible with having a puffy dwarf origin.

(v) We study the phase-space for cluster UDGs related to the main stellar population properties such as age, metallicity, α -enhancement and estimated quenching times, while looking for trends with the stellar mass and size of the UDGs. We find that the location in the phase-space diagram only hints for a relation with $[Mg/Fe]$ but not with any other property. We therefore caution about using such phase-space diagrams to determine the true locations in the cluster.

(vi) Many UDGs share similar properties (ages and metallicities) to dEs, suggesting a puffy dwarf-like origin for them. However, a subset of UDGs could have been formed as failed galaxies, as they present the lowest metallicities and lay below the local scaling relations but follow the high- z ones instead.

(vii) We include the information of the GC richness to investigate previous trends reported. We find that all GC-rich UDGs in our sample have intermediate-to-old ages and that all the young UDGs are GC-poor. Moreover, we find that the UDGs suggested as failed-galaxy UDGs are all GC-rich, further supporting this formation scenario for such UDGs.

(viii) UDGs have a large number of simulations trying to reproduce their properties. These simulations are successful in producing the relevant stellar population properties for a large fraction of observed UDGs in this work. However, there is no single simulation that can recreate the observed diversity of UDG properties and many of the simulations are unable to correctly reproduce some of the relevant properties for the non-UDG population. Moreover, the observed UDGs that match expectations of $z \sim 2$ galaxies are not reproduced by any simulation as of yet.

We thus find that there exists a wide range of formation pathways to create a UDG. Failed-galaxies UDGs are one of these pathways, although they appear more rare than those following a puffy-dwarf formation scenario in our data. In order to seek conclusive trends that can help differentiate the various formation pathways of UDGs, we are in need of new, more complete spectroscopic studies targeting UDGs (e.g. GC-poor UDGs, low-density environment UDGs).

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DATA AVAILABILITY

The KCWI data presented are available via the Keck Observatory Archive, 18 months after observations are taken.

REFERENCES

- Alabi A. et al., 2018, *MNRAS*, 479, 3308
 Amorisco N. C., 2018, *MNRAS*, 475, L116
 Amorisco N. C., Loeb A., 2016, *MNRAS*, 459, L51
 Amorisco N. C., Monachesi A., Agnello A., White S. D. M., 2018, *MNRAS*, 475, 4235
 Barbosa C. E. et al., 2020, *ApJ*, 247, 46
 Beasley M. A., Trujillo I., 2016, *ApJ*, 830, 23
 Beasley M. A., Romanowsky A. J., Pota V., Navarro I. M., Martínez Delgado D., Neyer F., Deich A. L., 2016, *ApJ*, 819, L20
 Bellazzini M., Belokurov V., Magrini L., Fraternali F., Testa V., Beccari G., Marchetti A., Carini R., 2017, *MNRAS*, 467, 3751
 Benavides J. A. et al., 2021, *Nat. Astron.*, 5, 1255
 Benavides J. A., Sales L. V., Abadi M. G., Marinacci F., Vogelsberger M., Hernquist L., 2023, *MNRAS*, 522, 1033
 Bidaran B. et al., 2022, *MNRAS*, 515, 4622
 Burkert A., Forbes D. A., 2020, *AJ*, 159, 56
 Buzzo M. L. et al., 2022, *MNRAS*, 517, 2231
 Cappellari M., 2017, *MNRAS*, 466, 798
 Cardona-Barrero S., Di Cintio A., Brook C. B. A., Ruiz-Lara T., Beasley M. A., Falcón-Barroso J., Macciò A. V., 2020, *MNRAS*, 497, 4282
 Carleton T., Errani R., Cooper M., Kaplinghat M., Peñarrubia J., Guo Y., 2019, *MNRAS*, 485, 382
 Carleton T. et al., 2023, *ApJ*, 953, 83
 Chan T. K., Kereš D., Wetzel A., Hopkins P. F., Faucher-Giguère C. A., El-Badry K., Garrison-Kimmel S., Boylan-Kolchin M., 2018, *MNRAS*, 478, 906
 Chilingarian I. V., Afanasiev A. V., Grishin K. A., Fabricant D., Moran S., 2019, *ApJ*, 884, 79
 Collins M. L. M., Read J. I., 2022, *Nat. Astron.*, 6, 647
 D’Ago G. et al., 2023, *A&A*, 672, A17
 Danieli S., van Dokkum P., Abraham R., Conroy C., Dolphin A. E., Romanowsky A. J., 2020, *ApJ*, 895, L4
 Danieli S. et al., 2022, *ApJ*, 927, L28
 de La Rosa I. G., La Barbera F., Ferreras I., de Carvalho R. R., 2011, *MNRAS*, 418, L74
 Di Cintio A., Brook C. B., Dutton A. A., Macciò A. V., Obreja A., Dekel A., 2017, *MNRAS*, 466, L1
 Du M. et al., 2019, *ApJ*, 875, 58
 Fan X., Carilli C. L., Keating B., 2006, *ARA&A*, 44, 415
 Fensch J. et al., 2019, *A&A*, 625, A77
 Ferré-Mateu A., Forbes D. A., Romanowsky A. J., Janz J., Dixon C., 2018a, *MNRAS*, 473, 1819
 Ferré-Mateu A. et al., 2018b, *MNRAS*, 479, 4891 (FM + 18)
 Ferré-Mateu A., Durré M., Forbes D. A., Romanowsky A. J., Alabi A., Brodie J. P., McDermid R. M., 2021, *MNRAS*, 503, 5455
 Fillingham S. P., Cooper M. C., Wheeler C., Garrison-Kimmel S., Boylan-Kolchin M., Bullock J. S., 2015, *MNRAS*, 454, 2039
 Forbes D. A., Alabi A., Brodie J. P., Romanowsky A. J., 2019, *MNRAS*, 489, 3665
 Forbes D. A., Alabi A., Romanowsky A. J., Brodie J. P., Arimoto N., 2020a, *MNRAS*, 492, 4874
 Forbes D. A., Dullo B. T., Gannon J., Couch W. J., Iodice E., Spavone M., Cantiello M., Schipani P., 2020b, *MNRAS*, 494, 5293
 Forbes D. A., Ferré-Mateu A., Durré M., Brodie J. P., Romanowsky A. J., 2020c, *MNRAS*, 497, 765
 Forbes D. A., Gannon J. S., Romanowsky A. J., Alabi A., Brodie J. P., Couch W. J., Ferré-Mateu A., 2021, *MNRAS*, 500, 1279
 Forbes D. A., Ferré-Mateu A., Gannon J. S., Romanowsky A. J., Carlin J. L., Brodie J. P., Day J., 2022, *MNRAS*, 512, 802
 Gallart C. et al., 2015, *ApJ*, 811, L18
 Gallazzi A., Charlot S., Brinchmann J., White S. D. M., Tremonti C. A., 2005, *MNRAS*, 362, 41
 Gallazzi A. R., Pasquali A., Zibetti S., Barbera F. L., 2021, *MNRAS*, 502, 4457
 Gannon J. S., Forbes D. A., Romanowsky A. J., Ferré-Mateu A., Couch W. J., Brodie J. P., 2020, *MNRAS*, 495, 2582
 Gannon J. S. et al., 2021, *MNRAS*, 502, 3144
 Gannon J. S. et al., 2022, *MNRAS*, 510, 946
 Gannon J. S., Forbes D. A., Brodie J. P., Romanowsky A. J., Couch W. J., Ferré-Mateu A., 2023, *MNRAS*, 518, 3653
 Greco J. P., Goulding A. D., Greene J. E., Strauss M. A., Huang S., Kim J. H., Komiyama Y., 2018, *ApJ*, 866, 112
 Grishin K. A., Chilingarian I. V., Afanasiev A. V., Fabricant D., Katkov I. Y., Moran S., Yagi M., 2021, *Nat. Astron.*, 5, 1308
 Gu M. et al., 2018, *ApJ*, 859, 37
 Harris W. E., Harris G. L. H., Alessi M., 2013, *ApJ*, 772, 82
 Heesters N. et al., 2023, *A&A*, 676, A33
 Ikeda R. et al., 2023, *MNRAS*, 523, 6310
 Impey C., Bothun G., Malin D., 1988, *ApJ*, 330, 634
 Iodice E. et al., 2020, *A&A*, 642, A48
 Janssens S., Abraham R., Brodie J., Forbes D., Romanowsky A. J., van Dokkum P., 2017, *ApJ*, 839, L17
 Janssens S. R., Abraham R., Brodie J., Forbes D. A., Romanowsky A. J., 2019, *ApJ*, 887, 92
 Janssens S. R. et al., 2022, *MNRAS*, 517, 858
 Janz J. et al., 2016, *MNRAS*, 456, 617
 Jiang F., Dekel A., Freundlich J., Romanowsky A. J., Dutton A. A., Macciò A. V., Di Cintio A., 2019, *MNRAS*, 487, 5272
 Jones M. G., Bennet P., Mutlu-Pakdil B., Sand D. J., Spekkens K., Crnojević D., Karunakaran A., Zaritsky D., 2021, *ApJ*, 919, 72
 Jones M. G. et al., 2023, *ApJ*, 942, L5
 Junais et al., 2022, *A&A*, 667, A76
 Koda J., Yagi M., Yamanoi H., Komiyama Y., 2015, *ApJ*, 807, L2
 Kroupa P., 2001, *MNRAS*, 322, 231
 Lee J. H., Kang J., Lee M. G., Jang I. S., 2020, *ApJ*, 894, 75
 Leisman L. et al., 2017, *ApJ*, 842, 133

- Liao S. et al., 2019, *MNRAS*, 490, 5182
- Lim S., Peng E. W., Côté P., Sales L. V., den Brok M., Blakeslee J. P., Guhathakurta P., 2018, *ApJ*, 862, 82
- Lim S. et al., 2020, *ApJ*, 899, 69
- Liu Y. et al., 2016, *ApJ*, 818, 179
- Ma X., Hopkins P. F., Faucher-Giguère C.-A., Zolman N., Muratov A. L., Kereš D., Quataert E., 2016, *MNRAS*, 456, 2140
- Mancera Piña P. E., Peletier R. F., Aguerri J. A. L., Venhola A., Trager S., Choque Challapa N., 2018, *MNRAS*, 481, 4381
- Marleau F. R. et al., 2021, *A&A*, 654, A105
- Martin G. et al., 2019, *MNRAS*, 485, 796
- Martín-Navarro I. et al., 2019, *MNRAS*, 484, 3425
- Martínez-Delgado D. et al., 2016, *AJ*, 151, 96
- Matteucci F., 1994, *A&A*, 288, 57
- McDermid R. M. et al., 2015, *MNRAS*, 448, 3484
- Mihos J. C. et al., 2015, *ApJ*, 809, L21
- Montes M., Infante-Sainz R., Madrigal-Aguado A., Román J., Monelli M., Borlaff A. S., Trujillo I., 2020, *ApJ*, 904, 114
- Montes M., Trujillo I., Infante-Sainz R., Monelli M., Borlaff A. S., 2021, *ApJ*, 919, 56
- Morrissey P. et al., 2018, *ApJ*, 864, 93
- Müller O., Jerjen H., Binggeli B., 2018, *A&A*, 615, A105
- Müller O. et al., 2020, *A&A*, 640, A106
- Müller O. et al., 2021, *ApJ*, 923, 9
- Muzzin A., Wilson G., Lacy M., Yee H. K. C., Stanford S. A., 2008, *ApJ*, 686, 966
- Naidu R. P. et al., 2022, preprint (arXiv:2204.09057)
- Nelson D. et al., 2018, *MNRAS*, 475, 624
- Ogiya G., 2018, *MNRAS*, 480, L106
- Pandya V. et al., 2018, *ApJ*, 858, 29
- Panter B., Jimenez R., Heavens A. F., Charlot S., 2008, *MNRAS*, 391, 1117
- Pasquali A., Smith R., Gallazzi A., De Lucia G., Zibetti S., Hirschmann M., Yi S. K., 2019, *MNRAS*, 484, 1702
- Peng E. W., Lim S., 2016, *ApJ*, 822, L31
- Recio-Blanco A., 2018, *A&A*, 620, A194
- Rhee J., Smith R., Choi H., Yi S. K., Jaffé Y., Candlish G., Sánchez-Jánsen R., 2017, *ApJ*, 843, 128
- Román J., Trujillo I., 2017, *MNRAS*, 468, 703
- Román J., Beasley M. A., Ruiz-Lara T., Valls-Gabaud D., 2019, *MNRAS*, 486, 823
- Romero-Gómez J. et al., 2023, *MNRAS*, 522, 130
- Rong Y., Guo Q., Gao L., Liao S., Xie L., Puzia T. H., Sun S., Pan J., 2017, *MNRAS*, 470, 4231
- Rong Y., Mancera Piña P. E., Tempel E., Puzia T. H., De Rijcke S., 2020a, *MNRAS*, 498, L72
- Rong Y., Zhu K., Johnston E. J., Zhang H.-X., Cao T., Puzia T. H., Galaz G., 2020b, *ApJ*, 899, L12
- Ruiz-Lara T. et al., 2018a, *MNRAS*, 478, 2034
- Ruiz-Lara T. et al., 2018b, *A&A*, 617, A18
- Safarzadeh M., Scannapieco E., 2017, *ApJ*, 850, 99
- Saifollahi T., Trujillo I., Beasley M. A., Peletier R. F., Knapen J. H., 2021, *MNRAS*, 502, 5921
- Saifollahi T., Zaritsky D., Trujillo I., Peletier R. F., Knapen J. H., Amorisco N., Beasley M. A., Donnerstein R., 2022, *MNRAS*, 511, 4633
- Sales L. V., Navarro J. F., Peñafiel L., Peng E. W., Lim S., Hernquist L., 2020, *MNRAS*, 494, 1848
- Sales L. V., Wetzel A., Fattahi A., 2022, *Nat. Astron.*, 6, 897
- Sandage A., Binggeli B., 1984, *AJ*, 89, 919
- Shen Z. et al., 2021, *ApJ*, 914, L12
- Sifón C., van der Burg R. F. J., Hoekstra H., Muzzin A., Herbonnet R., 2018, *MNRAS*, 473, 3747
- Simon J. D., 2019, *ARA&A*, 57, 375
- Smith R. J., Lucey J. R., Hudson M. J., Allanson S. P., Bridges T. J., Hornschemeier A. E., Marzke R. O., Miller N. A., 2009, *MNRAS*, 392, 1265
- Spitler L. R., Forbes D. A., 2009, *MNRAS*, 392, L1
- Thomas D., Maraston C., Bender R., Mendes de Oliveira C., 2005, *ApJ*, 621, 673
- Toloba E. et al., 2018, *ApJ*, 856, L31
- Tremmel M., Karcher M., Governato F., Volonteri M., Quinn T. R., Pontzen A., Anderson L., Bellovary J., 2017, *MNRAS*, 470, 1121
- Tremmel M., Wright A. C., Brooks A. M., Munshi F., Nagai D., Quinn T. R., 2020, *MNRAS*, 497, 2786
- Trujillo I., Roman J., Filho M., Sánchez Almeida J., 2017, *ApJ*, 836, 191
- Trujillo I. et al., 2019, *MNRAS*, 486, 1192
- Trujillo-Gomez S., Kruijssen J. M. D., Reina-Campos M., 2022, *MNRAS*, 510, 3356
- van der Burg R. F. J., Muzzin A., Hoekstra H., 2016, *A&A*, 590, A20
- van Dokkum P. G., Abraham R., Merritt A., Zhang J., Geha M., Conroy C., 2015, *ApJ*, 798, L45
- van Dokkum P. et al., 2016, *ApJ*, 828, L6
- van Dokkum P. et al., 2017, *ApJ*, 844, L11
- van Dokkum P., Danieli S., Cohen Y., Romanowsky A. J., Conroy C., 2018, *ApJ*, 864, L18
- van Dokkum P. et al., 2019, *ApJ*, 880, 91
- Van Nest J. D., Munshi F., Wright A. C., Tremmel M., Brooks A. M., Nagai D., Quinn T., 2022, *ApJ*, 926, 92
- Vazdekis A. et al., 2015, *MNRAS*, 449, 1177
- Villaume A. et al., 2022, *ApJ*, 924, 32
- Wang L., Dutton A. A., Stinson G. S., Macciò A. V., Penzo C., Kang X., Keller B. W., Wadsley J., 2015, *MNRAS*, 454, 83
- Wasserman A., Romanowsky A. J., Brodie J., van Dokkum P., Conroy C., Abraham R., Cohen Y., Danieli S., 2018, *ApJ*, 863, L15
- Watkins A. E., Salo H., Kaviraj S., Collins C. A., Knapen J. H., Venhola A., Román J., 2023, *MNRAS*, 521, 2012
- Webb K. A. et al., 2022, *MNRAS*, 516, 3318
- Weisz D. R., Dolphin A. E., Skillman E. D., Holtzman J., Gilbert K. M., Dalcanton J. J., Williams B. F., 2014, *ApJ*, 789, 147
- Weisz D. R., Dolphin A. E., Skillman E. D., Holtzman J., Gilbert K. M., Dalcanton J. J., Williams B. F., 2015, *ApJ*, 804, 136
- Weisz D. R. et al., 2019, *ApJ*, 885, L8
- Westfall K. B. et al., 2019, *AJ*, 158, 231
- Wright A. C., Tremmel M., Brooks A. M., Munshi F., Nagai D., Sharma R. S., Quinn T. R., 2021, *MNRAS*, 502, 5370
- Yozin C., Bekki K., 2015, *MNRAS*, 452, 937
- Zaritsky D., Donnerstein R., Karunakaran A., Barbosa C. E., Dey A., Kadowaki J., Spekkens K., Zhang H., 2022, *ApJS*, 261, 11
- Zaritsky D., Donnerstein R., Dey A., Karunakaran A., Kadowaki J., Khim D. J., Spekkens K., Zhang H., 2023, *ApJS*, 267, 27

APPENDIX A: NEW GLOBULAR CLUSTER RICHNESS ESTIMATES

In this paper we collect published GC counts and classify the UDGs between GC-poor and GC-rich following previous works (e.g. Gannon et al. 2022). This broad and *arbitrary* limit, which is set at ~ 20 GCs, aims at separating UDGs with dwarf-like dark matter haloes from UDGs with more massive ones (Berkert & Forbes 2020). We note that for the UDGs with already published values, none of them would be classified differently if accounting for the errors in the number count or if considering the different values published.

However, measuring the number of GCs in a galaxy ideally requires deep, multiband imaging with excellent image quality. Both *HST* and ground-based imaging are used for distances out to ~ 20 Mpc, while greater distances normally require *HST*. Even so, the brighter GCs around a galaxy can be detected out to ~ 100 Mpc given particularly high-quality ground-based imaging, which can be enough for a crude assessment of how populous a GC system is. Firm, quantitative measurements require careful photometry, completeness tests, and detailed modelling of the GC spatial distribution and contamination fraction (e.g. Janssens et al. 2022).

None the less, in many cases it is possible to obtain a quick visual evaluation of GC richness by assessing the excess number of point-like sources associated with a low surface brightness galaxy. This approach was taken by Gannon et al. (2022) for UDGs in the Perseus cluster using *gri* imaging from Subaru / Hyper Suprime-Cam, and validated from detailed analysis of a subset of the galaxies using *HST*. We obtained *HST/ACS* images in either the F475W or F606W filter from the Hubble Legacy Archive. For training purposes, we used three Coma UDGs (DF02, Yagi358, Yagi275) with *HST*-based GC numbers in Forbes et al. (2020c). The visual inspection and votes have been carried independently by several co-authors, reaching a consensus for Yagi098, Yagi392, and Yagi090. Yagi418, OGS1, J130026.26+272735.2, and J130038.63+272835.3 are not as clear and are thus placed within parenthesis in Table 1. We therefore caution the reader about using these estimates until better imaging and a proper process is carried out. We note, however, that the results presented and discussed in this paper stand if we do not include the visual classification.

APPENDIX B: STELLAR POPULATION MEASUREMENTS

In this appendix, we discuss the main systematics related to the measurement of the most critical properties in this work, namely mean ages, metallicities, and α -enhancement ratios.

B1 α -enhancement

We have estimated the α -enhancement with three different approaches, as done in other works for faint galaxies with moderate quality spectra (e.g. FM+18, Ferré-Mateu et al. 2021; Forbes et al. 2022). The first method uses the full spectral fitting method (applied to all galaxies), while the other two employ an absorption line index analysis (only for those spectra with no pedestal issue, as discussed in Section 3.1). We therefore obtain three independent estimates:

(i) using the SSP model predictions with $[\alpha/\text{Fe}]=0.0$ and $[\alpha/\text{Fe}] = +0.4$ dex we obtain the $[\text{Mg}/\text{Fe}]$ from the comparison of the metal indices Mg_b versus Fe5015. While it is easy to interpolate within the two model grids, extrapolations outside of it make it more difficult. We therefore set as limits for a safe extrapolation of $[\alpha/\text{Fe}] = [-0.1, +0.6]$ dex.

(ii) using two metallic indices, Mg_b and Fe5015 separately and fixing the ages of the galaxies to those obtained from pPXF. This provides a more orthogonal grid that does not typically need any extrapolation. From this we obtain the relative metallicities, Z_{Mg_b} and Z_{Fe} , which are used to obtain the proxy $[\text{Mg}/\text{Fe}] = Z_{\text{Mg}_b} - Z_{\text{Fe}}$.

(iii) using the full spectral fitting approach, we run pPXF with a set of SSP models that range from scaled solar to $[\alpha/\text{Fe}] = +0.4$ dex. This method will always deliver values only within the SSP models used. We note that because this is applied to the entire spectra, in principle will contain several α -elements. Therefore this value is closer to a ‘true’ $[\alpha/\text{Fe}]$ and can be converted into $[\text{Mg}/\text{Fe}]$ using the formula from Vazdekis et al. (2015) $[\alpha/\text{Fe}] = 0.02 + 0.56 \times [\text{Mg}/\text{Fe}]$.

Fig. B1 shows a comparison of the different $[\text{Mg}/\text{Fe}]$ obtained from the three methods. It shows the results from method (i) compared to method (ii) in the top panel, and from method (i) compared to method (iii) in the bottom one. The grey regions in this panel show the parameter space that cannot be obtained in the full spectral fitting

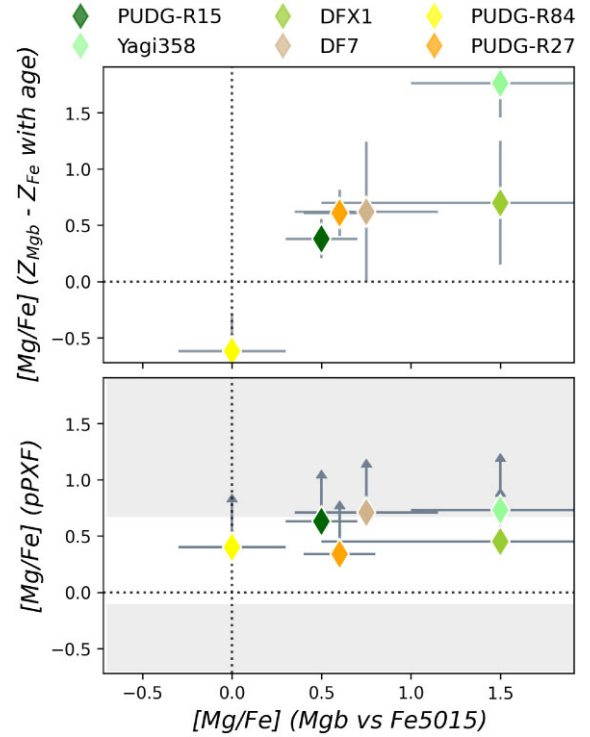


Figure B1. α -enhancement measured by different methods for the UDGs with no pedestal issue. Top: comparison of the $[\text{Mg}/\text{Fe}]$ derived from method (i) versus method (ii): the Mg_b versus Fe5015 indices for scaled solar and $[\alpha/\text{Fe}] = +0.4$ dex SSPs as opposed to age-SSP grids of Mg_b and Fe5015 to obtain the proxy $[Z_{\text{Mg}_b}/Z_{\text{Fe}}]$. Bottom: comparison of method (i) with the values obtained from the pPXF routine after converting into $[\text{Mg}/\text{Fe}]$ (method iii). This last method has the limitation on the SSP models, as shown by the shaded grey areas. The arrows show how much the values could move if a shift is applied, see the text.

method. We have marked as in Fig. 7 the lower limits for the pPXF estimates, if we were to apply a shift to compensate for this method. This shift has been obtained as the mean difference between the two estimates, and it is roughly $+0.25$ dex. Although the results from the line indices are compatible at a first order, all this highlights the need to obtain the $[\text{Mg}/\text{Fe}]$ abundance ratios from a line index analysis rather than using full spectral fitting, when possible. For the work presented here we will use the values that are obtained from method (i) for those UDGs with no pedestal issue and the results from pPXF for the remaining (marked in Table 2).

B2 Mean ages and metallicities

In Forbes et al. (2022) it was shown that for a sample of GCs observed with a short baseline that did not include H_β (the main age indicator in the optical range), tend to bias the results towards younger ages. Although we have this line in all our spectra, we still want to test the impact on using a short (4800 to 5300 Å) as opposed to a long (3800 to 5500 Å) baseline from the different gratings. Therefore we repeat the fitting for PUDG-R27, PUDG-R84, DFX1, and DF07, limiting now the fit to the short spectral range. The results are shown in Fig. B2. We also include the GCs from Forbes et al. (2022), where this was first studied for the KCWI gratings. We find that as long as H_β is included, the recovered parameters are not much affected.

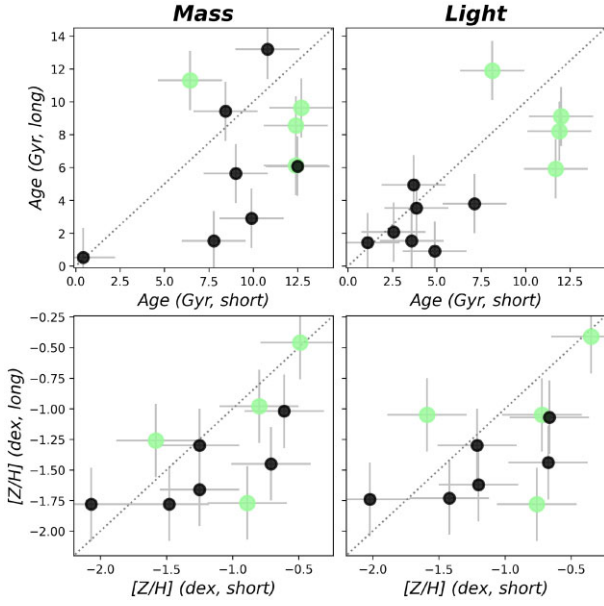


Figure B2. Comparison of using a long spectral range (as the KCWI BL grating) as opposed to a shorter one (BH) to test possible systematics in the derivation of the main stellar populations. Age (top) and metallicity (bottom) are shown, both for their mass-weighted (left panels) and light-weighted (right panels) values. Green points correspond to the UDGs observed with the BL configuration, which had their analysis repeated with the short baseline simulating the BH grating. Black symbols correspond to the GCs of Forbes et al. (2022), where it was shown that missing the H_{β} could bias the results towards younger ages. Overall, the biggest impact is seen only for the GCs that did not include the H_{β} .

B3 Comparison to other literature and SED fitting results

UDGs studies based on spectroscopy are still in their infancy given the challenges in obtaining large samples with high quality spectra. An alternative way to study some of the stellar populations is to obtain them from SED fitting, which tends to go deeper and is less time consuming. We next compare in Fig. B3 the spectroscopic results of the UDGs in this work that have been previously analysed using SED fitting by Buzzo et al. (2022). In the latter it was shown that the solution that includes a small amount of dust tends to be more similar to the spectroscopic results, hence those are the values shown in the figure. Overall, both stellar masses and ages are compatible, but the stellar metallicities derived using the SED fitting provide systematically lower metallicities, by about 0.25 dex (note however, that this is within the uncertainties). In fact, we find that the UDGs deviating the most in their metallicity tend to be the ones with very old stellar ages.

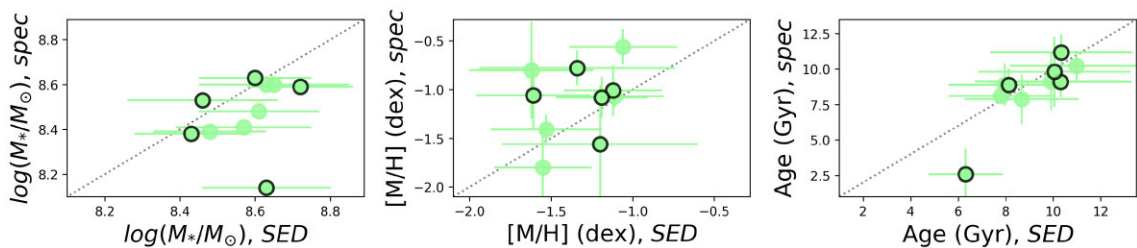


Figure B3. Comparison of the obtained properties with those from the work of Buzzo et al. (2022) based on SED fitting. Circles with black outline are the new UDGs from this work, while those without outline are from the literature. We present the derived stellar mass (left), metallicity (middle), and mass-weighted age (right). Overall, the stellar masses and ages are generally consistent. However, metallicities from SED fitting tend to be lower than those from spectroscopy, by roughly ~ 0.25 dex.

total halo masses, and proper GC counts, it is hard to point to a particular origin from the stellar populations alone.

APPENDIX C: PLAUSIBLE FORMATION SCENARIOS SUMMARY

We summarize in Table C1 all the properties discussed in this work. However, without further information on e.g. the dark matter content,

Table C1. Summary of properties for each UDG.

Galaxy	Phase-space	SFHs	Stellar populations Age / [M/H] / [Mg/Fe]	Relations Age-[M/H] / [Mg/Fe]-[M/H] / MZR	GC richness
VCC 1287	VEI	late + fast	O / L / EE	dE-dSph / GC / within	R
NGC 5846-UDG1	–	late + fast	O / L / EE	dE-dSph / GC / within	R
UDG1137+16	–	late + slow	Y / L / SS	? / GC-dSph / below	–
PUDG-R15	EI	late + fast	EO / L / EE	GC-dE / GC-dE / within	P
PUDG-S74	mix	early + slow	O / H / E	dE / dE / within	R
PUDG-R84	mix	late + fast	O / L / E	dSph / dSph-GC / below	R
PUDG-R24	LI	late + slow	Y / L / US	dE / dE / within	P
PUDG-R27	VEI	early + fast	EO / L / E	dE-GC / dE / within	R
Yagi358	VEI	early + fast	O / L / EE	dSph / ? / below	R
DFX1	mix	late + fast	O / L / EE	dE-dSp / GC / within	R
DF07	EI	early + fast	EO / L / EE	GC-dE / ? / within	R
Yagi093	VEI	early + slow	O / H / E	dE / dE / within	R
Yagi098	VEI	early + slow	I / H / -	dE / - / within	(R)
Yagi275	LI	late + slow	I / H / US	dE / dE / (above)	–
Yagi276	VEI	late + slow	I / H / -	dE / - / (above)	–
Yagi392	VEI	late + fast	I / H / -	dE / - / within	–
Yagi418	VEI	late + fast	O / L / SS	dE-dSph / dE-dSph-GC / within	(R)
Yagi090	LI	early + slow	I / L / E	dE / GC / within	(P)
OGS1	VEI	early + fast	O / H / E	dE / dE / within	(P)
DF17	VEI	–	O / L / -	dE / dE / within	R
DF44	EI	early + fast	EO / L / US	dSph-GC / dSph / below	R
J130026.26+272735.2	VEI	–	Y / H / -	? / - / within	(P)
J130038.63+272835.3	VEI	–	Y / H / -	dE / - / within	(P)
NGC 1052-DF2	–	–	O / L / SS	dE-dSph / dE-dSph / within	P
DGSAT I	–	late + slow	O / L / EE	dSph / ? / below	(P)

Notes: Column 2: location information from phase-space (VEI - very early infall; EI - early infall; mix - mix regions; LI - late infall)

Column 3: the type of SFH

Column 4: main stellar populations properties:

- Stellar age (EO - extremely old, >10 Gyr; O - old, 7.5 to 10 Gyr; I - intermediate, 4.5 to 7.5 Gyr; Y - young, <4.5 Gyr);

- Metallicities (L - low, $\lesssim -0.75$ dex; H, $\gtrsim -0.75$ dex);

- α -enhancement (EE - extremely enhanced, > 0.4 dex; E - enhanced, 0.2 to 0.4 dex; SS - scaled solar, 0.0 to 0.2 dex; US - under solar, <0.0 dex)

Column 5: type of object they resemble in the scaling relations and whether they follow or are located above or below the local MZR.

Column 6: GC richness (in parenthesis for the visual estimated)

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